

# Astroparticle Physics

## Lectures:

05.02.2019 [1. Historical introduction, basic properties of cosmic rays](#)

07.02.2019 [2. Hadronic interactions and accelerator data](#)

19.02.2019 [3. Cascade equations](#)

21.02.2019 [4. Electromagnetic cascades](#)

26.02.2019 [5. Extensive air showers](#)

28.02.2019 6. Detectors for extensive air showers

09.04.2019 7. High energy cosmic rays and the knee in the energy spectrum of cosmic rays

16.04.2019 8. Radio detection of extensive air showers

23.04.2019 9. Acceleration, astrophysical accelerators and beam dumps

07.05.2019 10. Extragalactic propagation of cosmic rays

14.05.2019 11. Ultra high energy cosmic rays

21.05.2019 12. Astrophysical gamma rays and neutrinos

28.05.2019 13. Neutrino astronomy

04.06.2019 14. Gamma-ray astronomy

<http://particle.astro.ru.nl/goto.html?astropart1819>

# lecture 6

# Detectors for extensive air showers

## *Gaisser chapter 16*

### 16 Extensive air showers

- 16.1 Basic features of air showers
- 16.2 The Heitler–Matthews splitting model
- 16.3 Muons in air showers
- 16.4 Nuclei and the superposition model
- 16.5 Elongation rate theorem
- 16.6 Shower universality and cross section measurement
- 16.7 Particle detector arrays
- 16.8 Atmospheric Cherenkov light detectors
- 16.9 Fluorescence telescopes
- 16.10 Radio signal detection

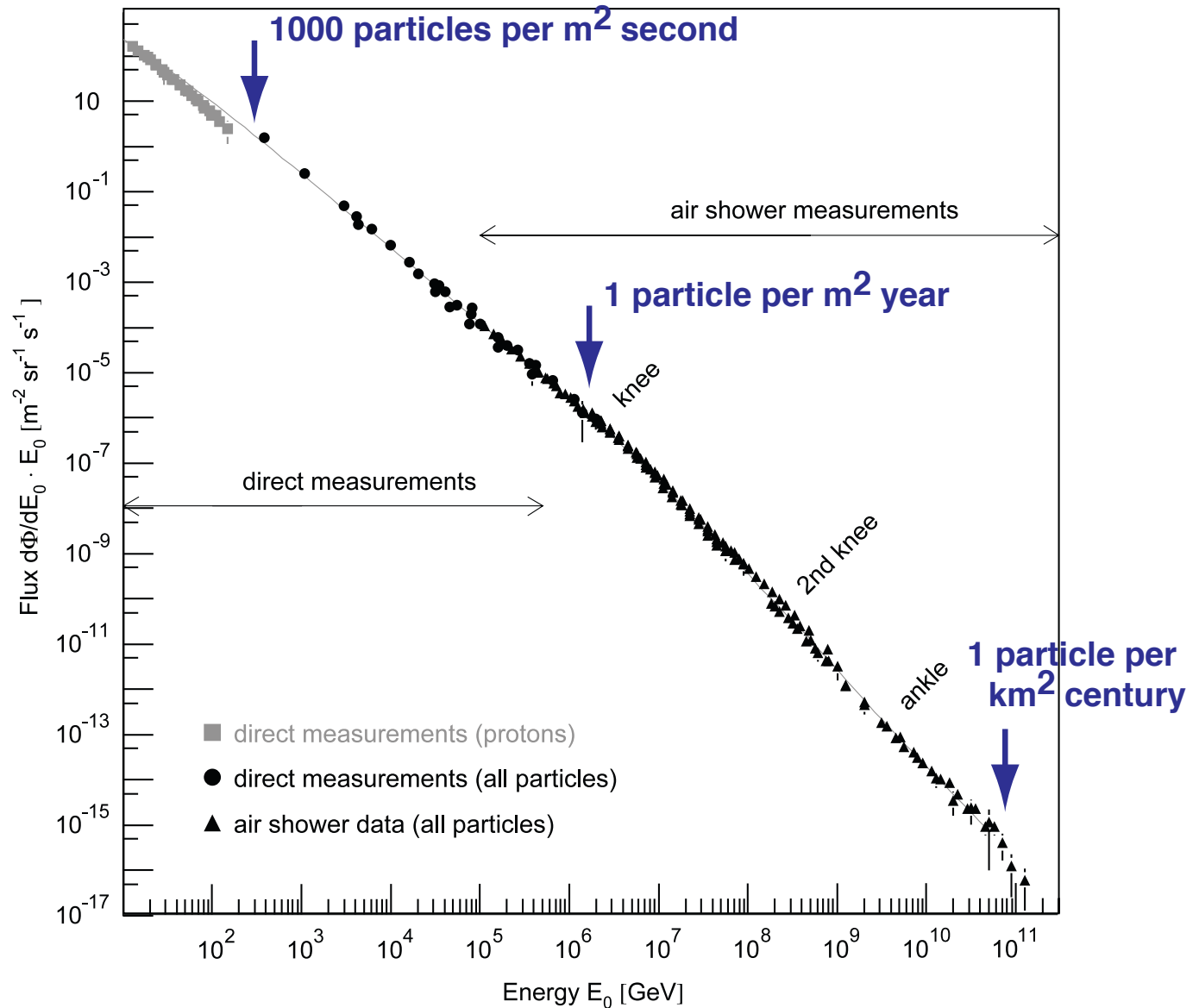


Fig. 1. All-particle energy spectrum of cosmic rays as measured directly with detectors above the atmosphere and with air shower detectors. At low energies, the flux of primary protons is shown.

# Measurement of Cosmic Rays at $E > 10$ TeV

## Extensive Air Showers

flux falls  $\sim E^{-2.7}$   $\rightarrow E = 10^{15}$  eV: 1 particle/m<sup>2</sup>/a

atmosphere acts as an absorber/calorimeter

vertical atmospheric depth  $1035 \frac{\text{g}}{\text{cm}^2}$   
 $\approx 12\lambda_I$   
 $\approx 30X_0$

$\rightarrow$  calorimetric measurement

# Extensive Air Shower

Proton  $10^{15}$  eV:

on ground

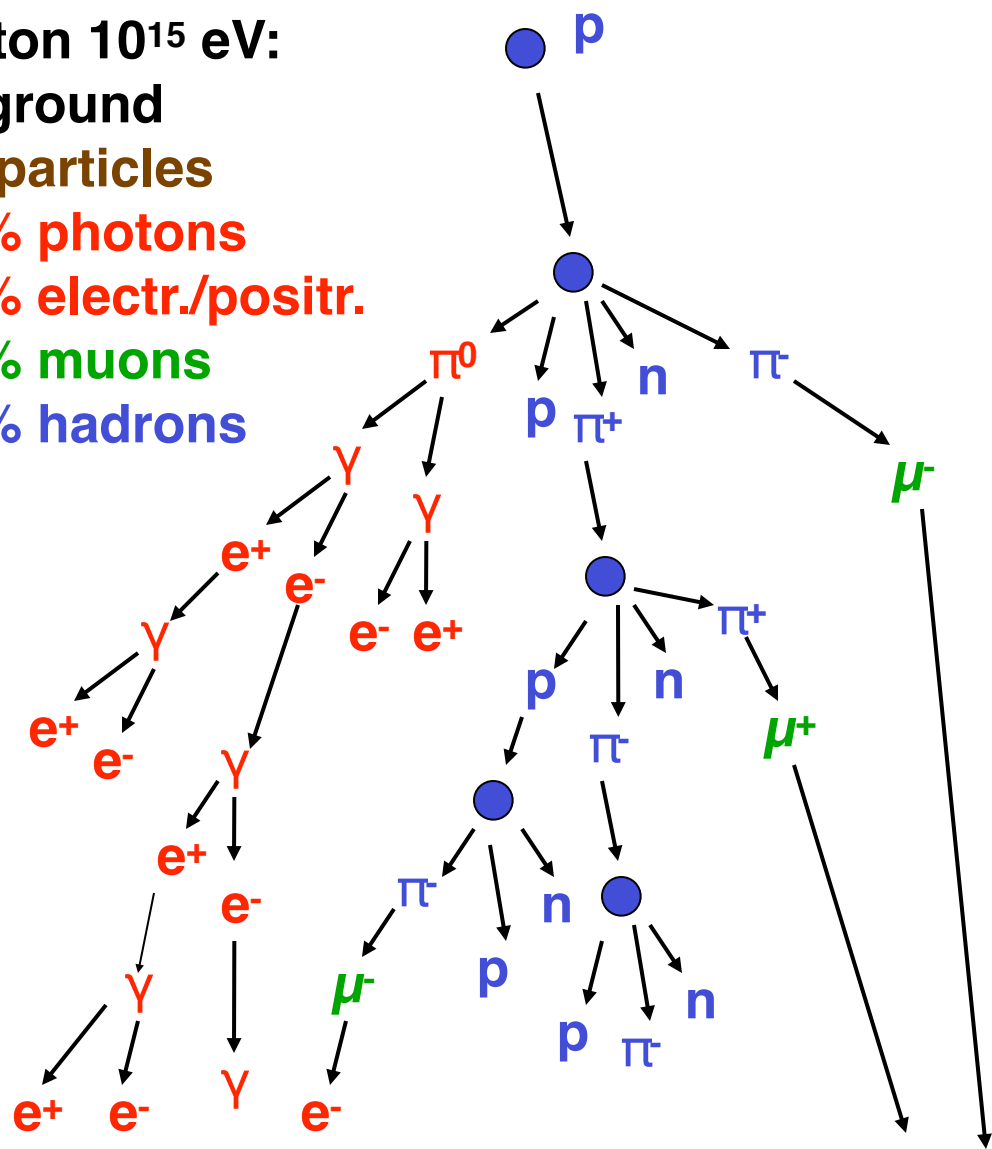
$10^6$  particles

80% photons

18% electr./positr.

1.7% muons

0.3% hadrons

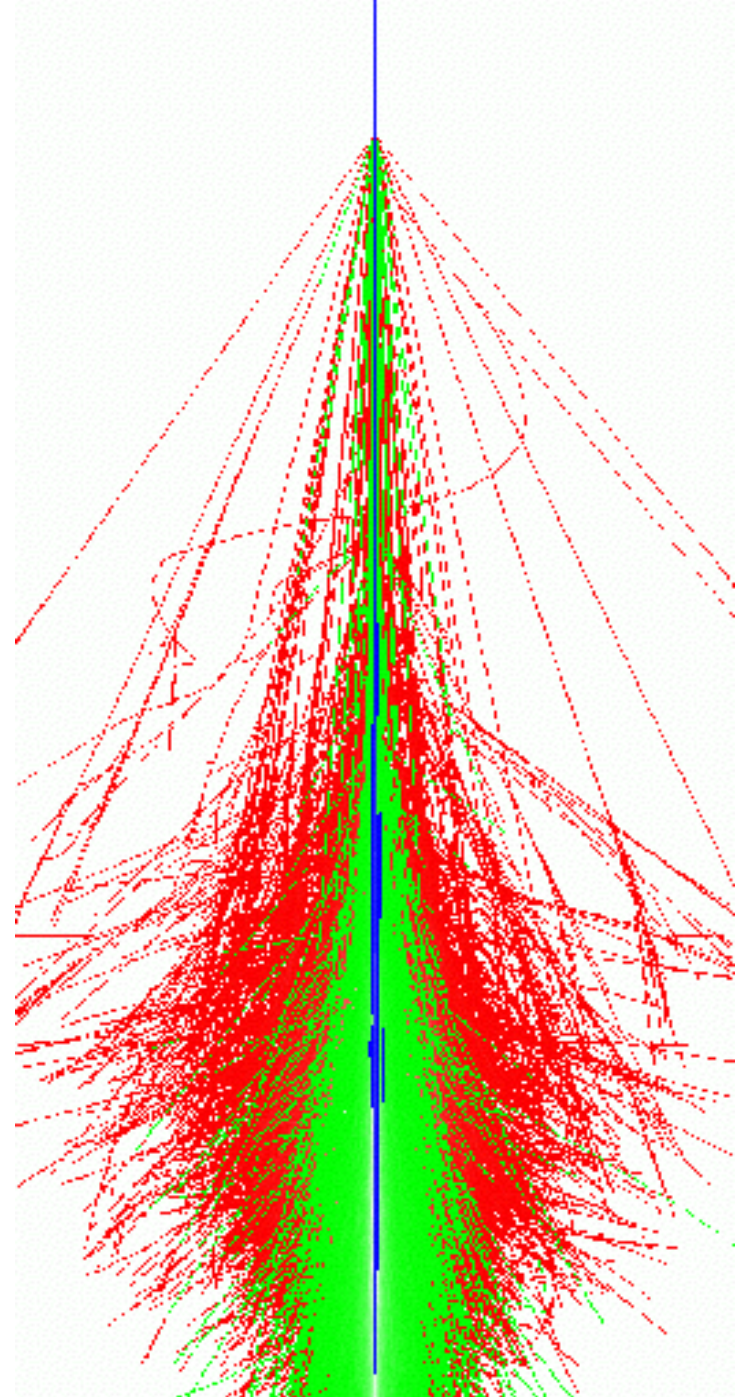


electromagnetic

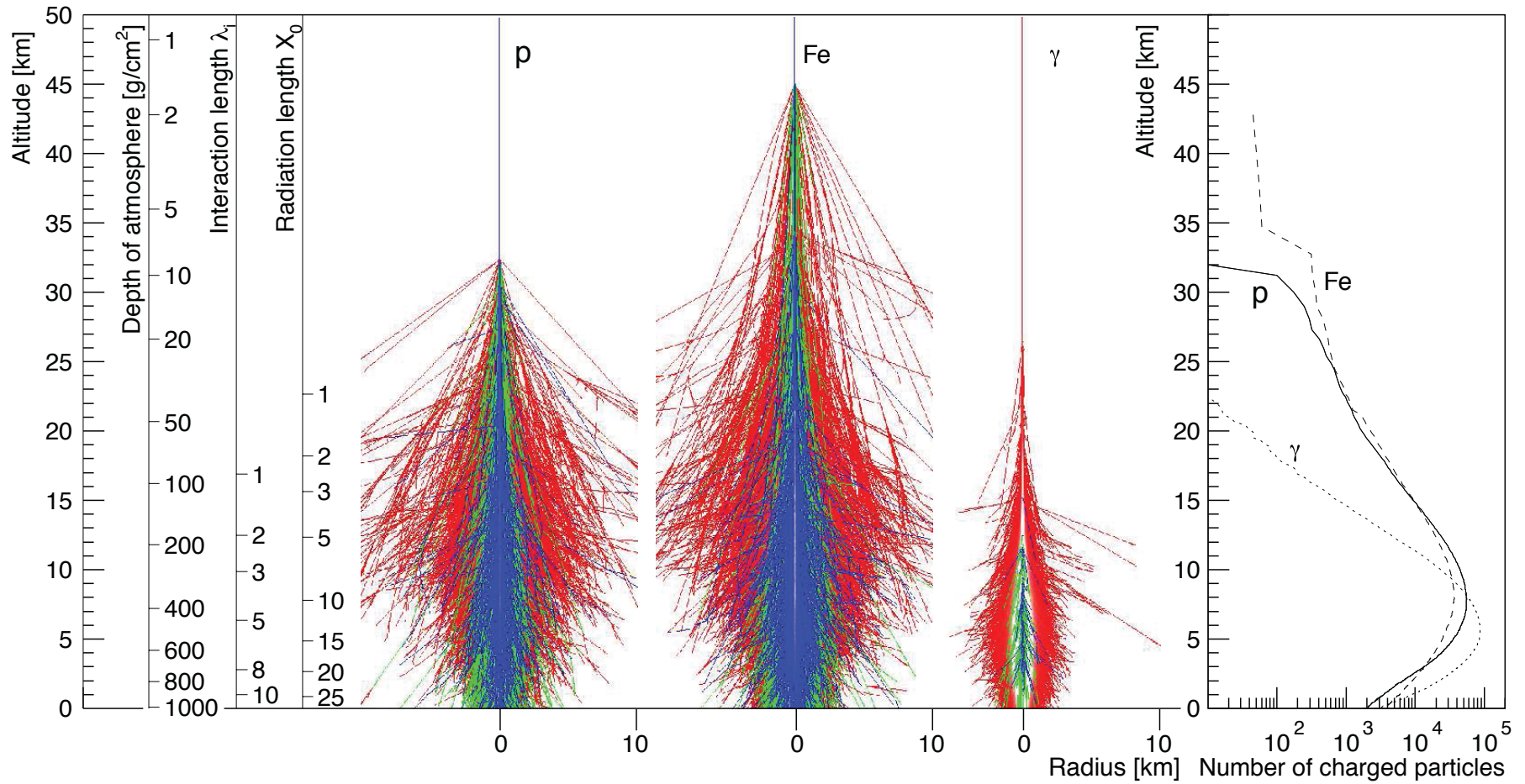
hadronic

muonic

shower component



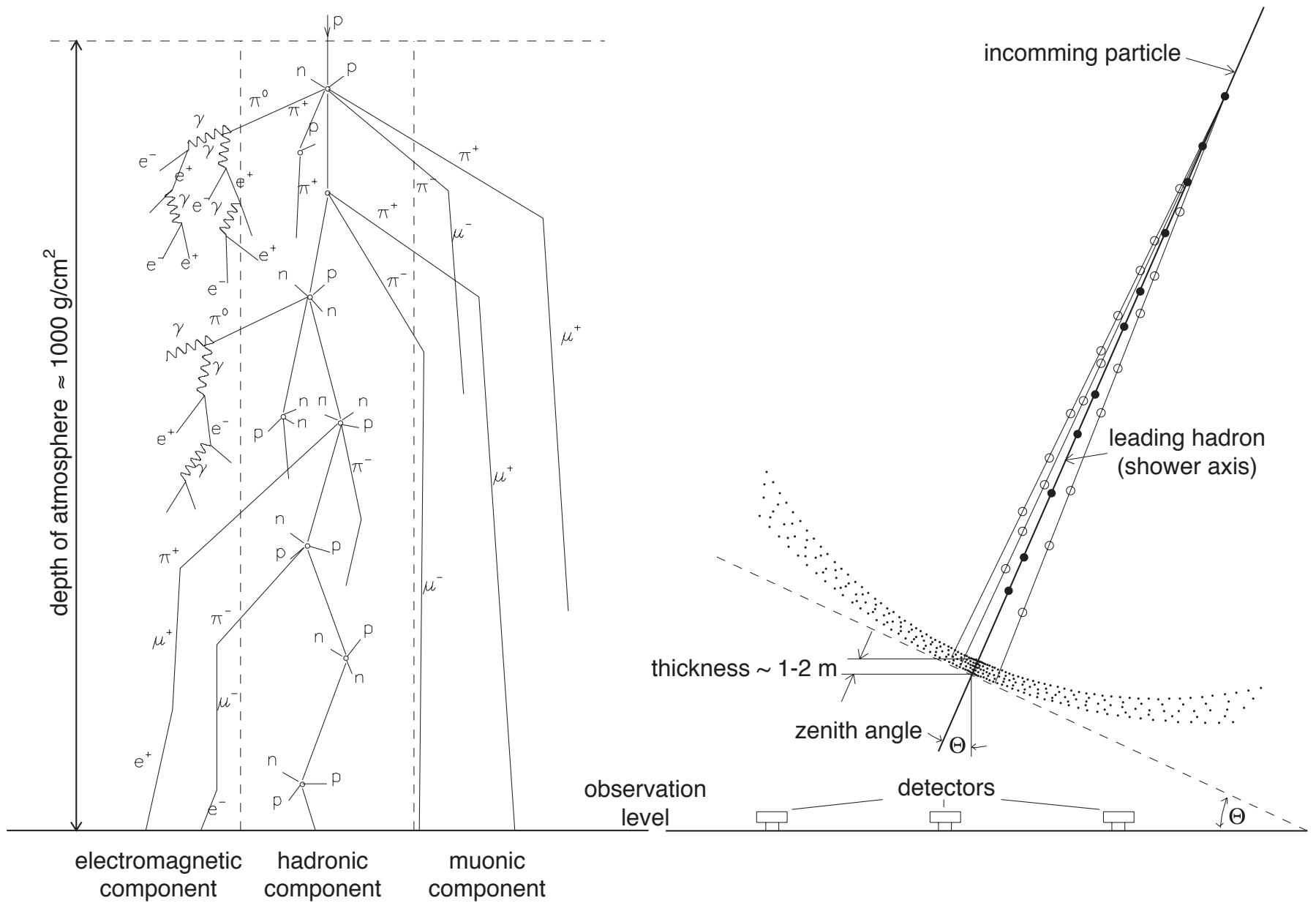
# examples of showers with same (total) energy



## 16.7 Particle detector arrays

Surface detector arrays consist of a set of particle detectors that are typically arranged in a regular pattern. Depending on the energy range the experiment is optimized for, the distance between the detector stations can vary from  $\sim 15$  m (KASCADE [536], Tibet AS- $\gamma$  [537]) up to more than 1000 m (Telescope Array [30], Auger Observatory [29]).

Showers are detected by searching for time coincidences of signals in neighboring detector stations. The arrival direction can then be determined from the time delay of the shower front reaching the different detectors. The shower comprises a disk of particles that is a few meters thick in the center, increasing up to a few hundred meters at large lateral distances. Only at small lateral distances can the curvature of the shower front be approximated as a sphere. The angular resolution of the reconstructed arrival direction depends on the distance and accuracy of time synchronization between the detector stations and the number of particles detected per station (for defining the arrival time of the shower front). Air shower arrays reach angular resolutions of typically  $1 - 2^\circ$  for low-energy showers and better than  $0.5^\circ$  for large showers.

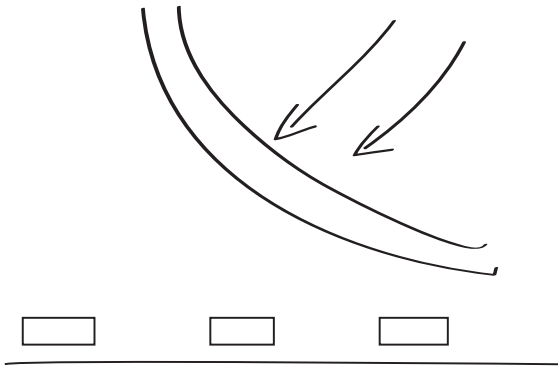




# Measurement methods

## particles at ground

**cascade of secondary particles allows to sample the shower at specific points**



**detector coverage**

**$10^{15}$  eV: 1% (15 m)**

**$10^{20}$  eV:  $10^{-8}$  (1.5 km)**

**example: KASCADE 40000 m<sup>2</sup> total area**

**500 m<sup>2</sup> e/m detectors (~1.2%), detector distance 13 m**

# KARlsruhe Shower Core and Array DETector

**Simultaneous measurement of  
electromagnetic,  
muonic,  
hadronic  
shower components**

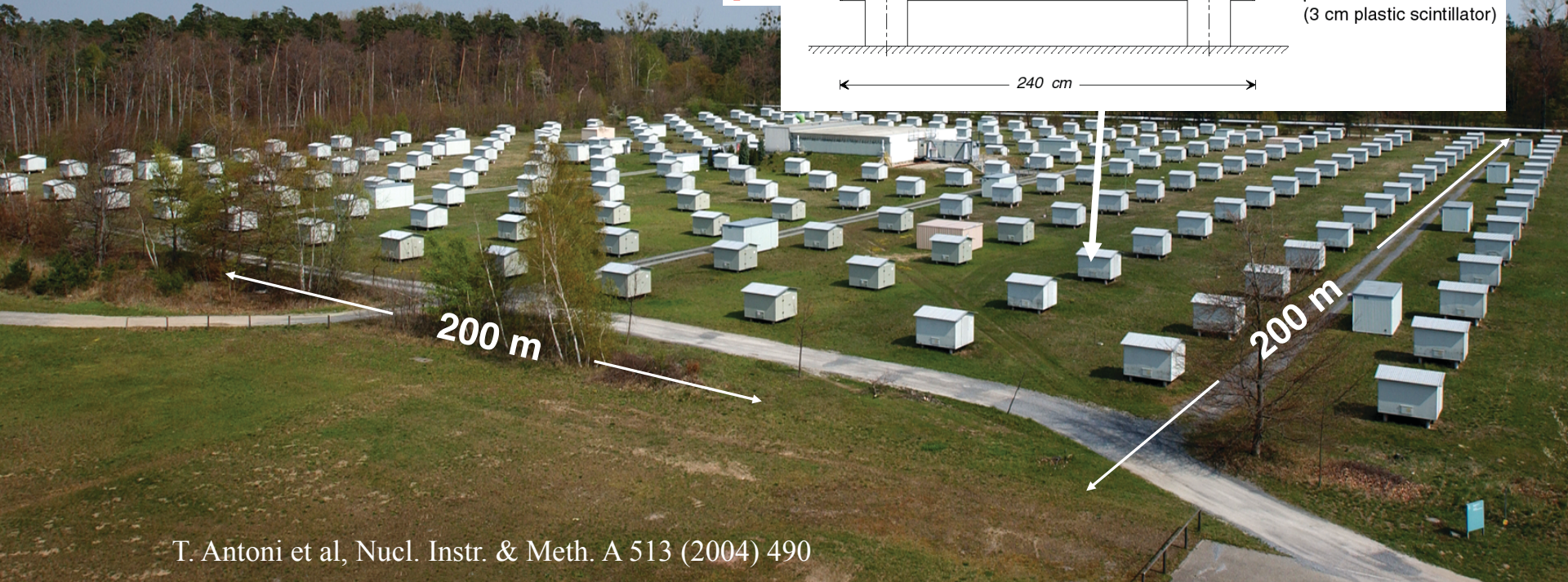
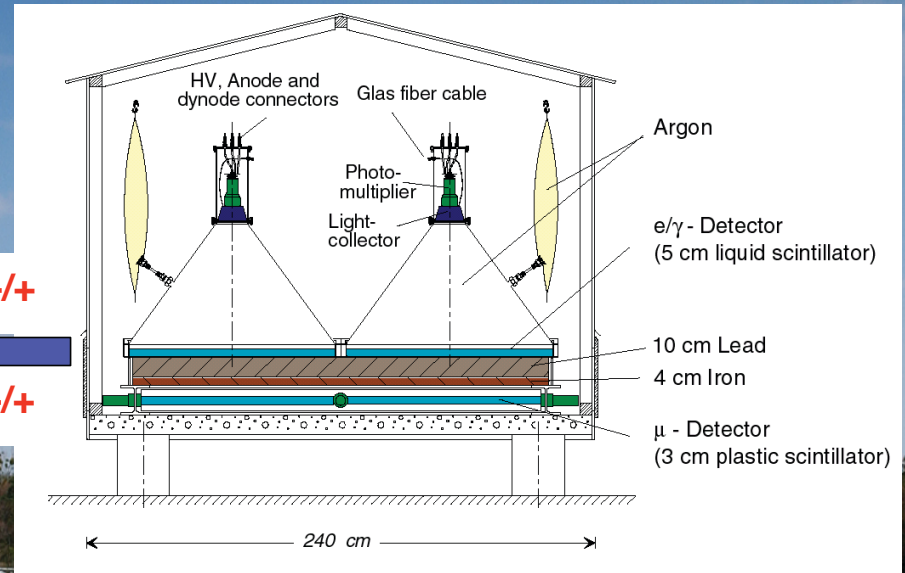


T. Antoni et al, Nucl. Instr. & Meth. A 513 (2004) 490

# KARlsruhe Shower Core and Array DETector

Simultaneous measurement of electromagnetic, muonic, hadronic shower components

$e^{-/+}$   
 $\mu^{-/+}$



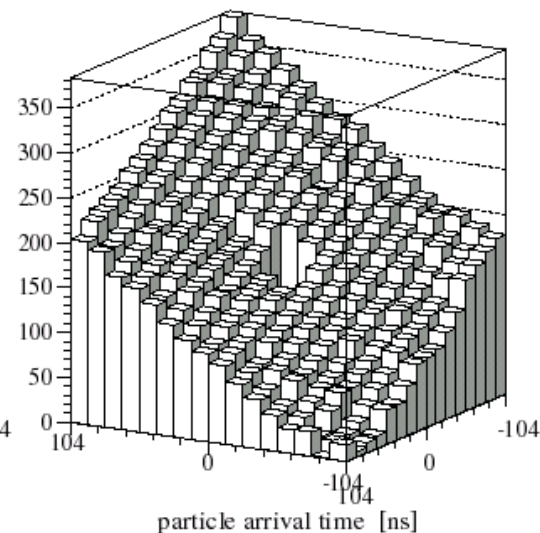
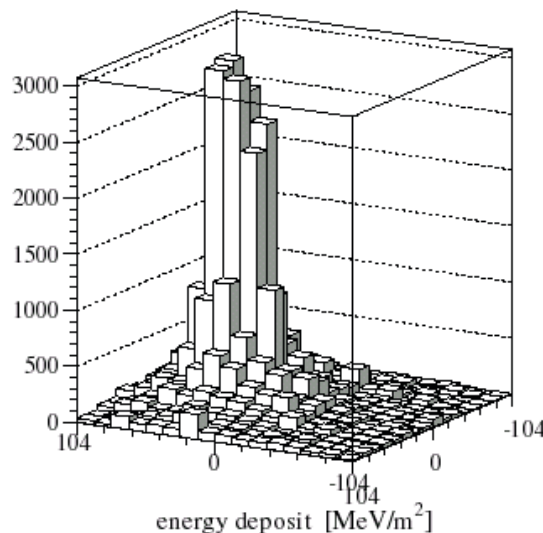
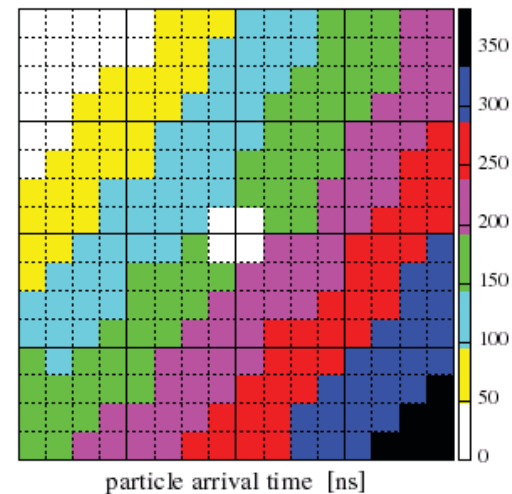
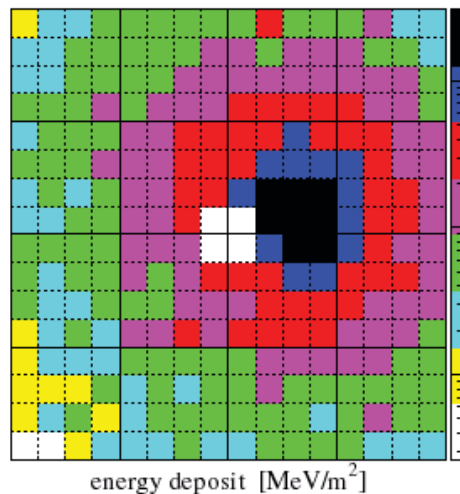
T. Antoni et al, Nucl. Instr. & Meth. A 513 (2004) 490

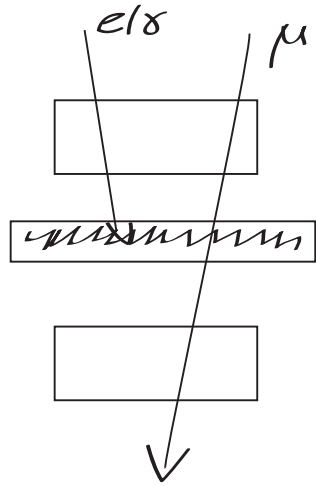
# Event reconstruction in the scintillator array

## electromagnetic component

e/ $\gamma$ -Detectors, Run 1, Event 71089, 96-03-05 22:07:48.956078

shower core	$\Delta r = 2.5 - 5.5 \text{ m}$
shower direction	$\Delta \theta = 0.5^\circ - 1.2^\circ$
shower size	$\Delta N_e/N_e = 6 - 12 \%$





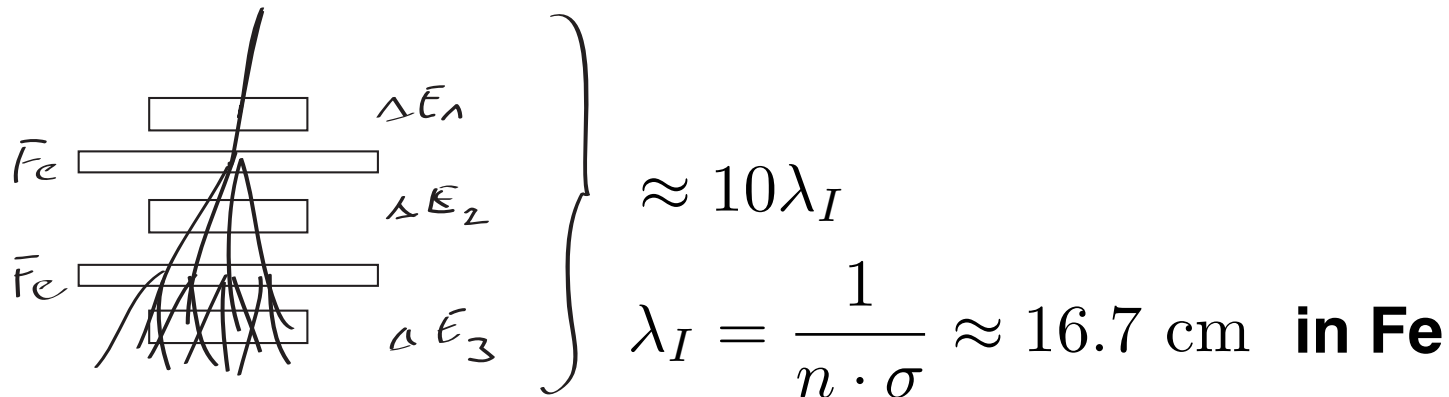
$e/\gamma$  detector

Fe+Pb absorber

$\mu$  detector

--> count the number of electrons and muons

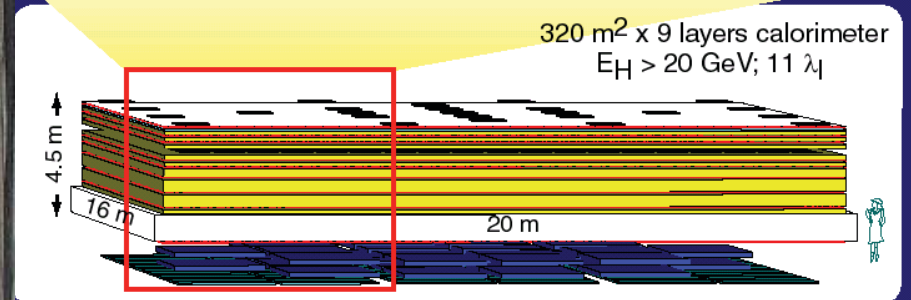
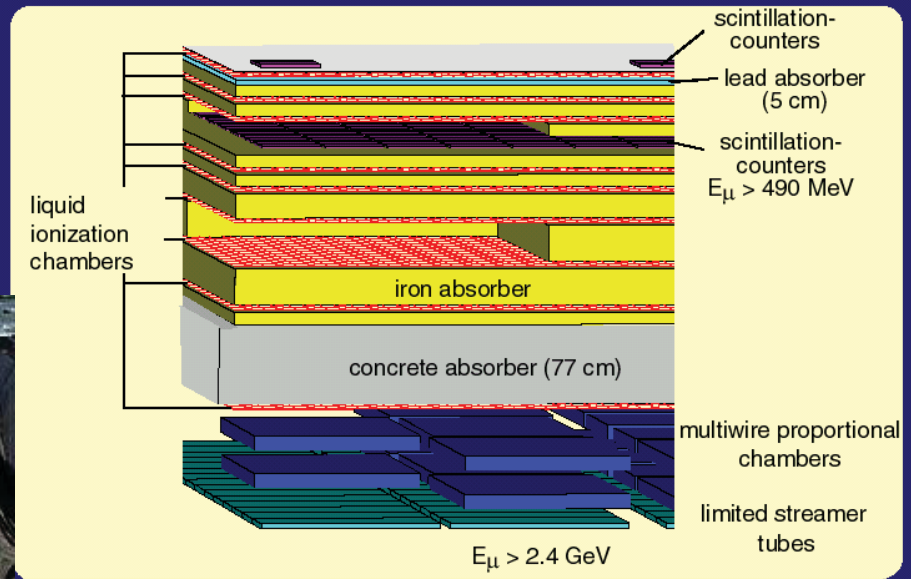
## measurement of hadrons hadron calorimeter



# KASCADE Hadron Calorimeter

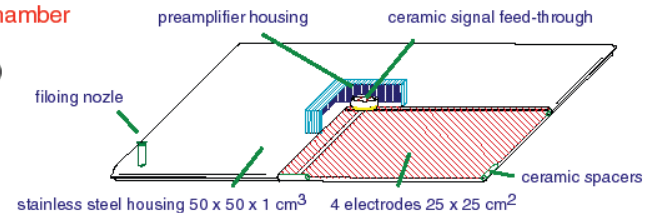


## KASCADE Hadron-Calorimeter



### Liquid ionization chamber

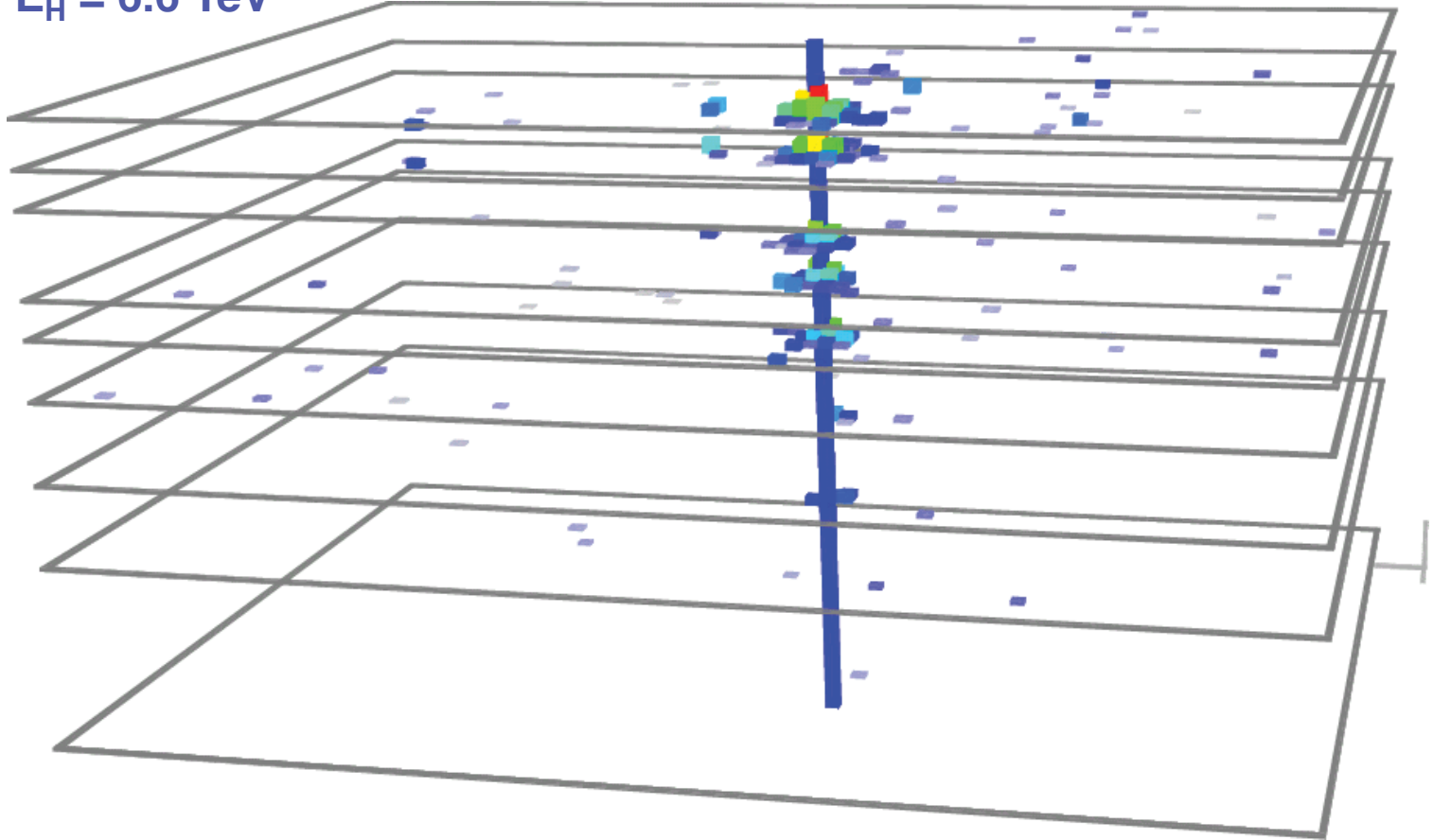
Tetramethylsilane (TMS)  
 Tetramethylpentane (TMP)



# Reconstruction of hadrons

Unaccompanied hadron

$E_H = 6.6 \text{ TeV}$



spatial resolution:

$\Delta_x \sim 10 - 12 \text{ cm}$

angular resolution:

$\Delta_\theta \sim 1^\circ - 3^\circ$

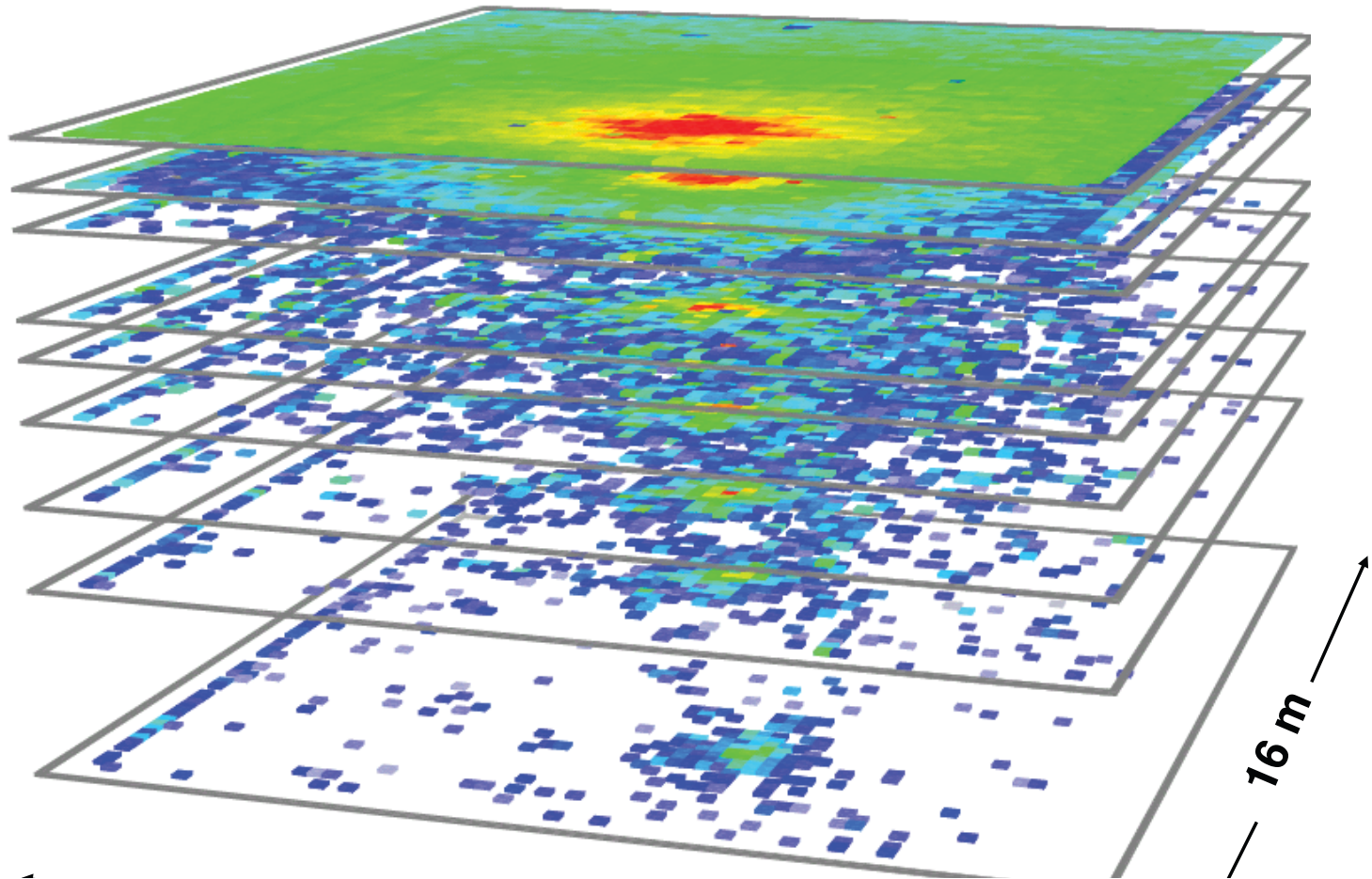
energy resolution:

$$\frac{\sigma(E)}{E} [\%] \approx \frac{250}{\sqrt{E/\text{GeV}}}$$

# Hadronic shower core

$E_0 \sim 6 \text{ PeV}$

Number of reconstructed hadrons  $N_h = 143$

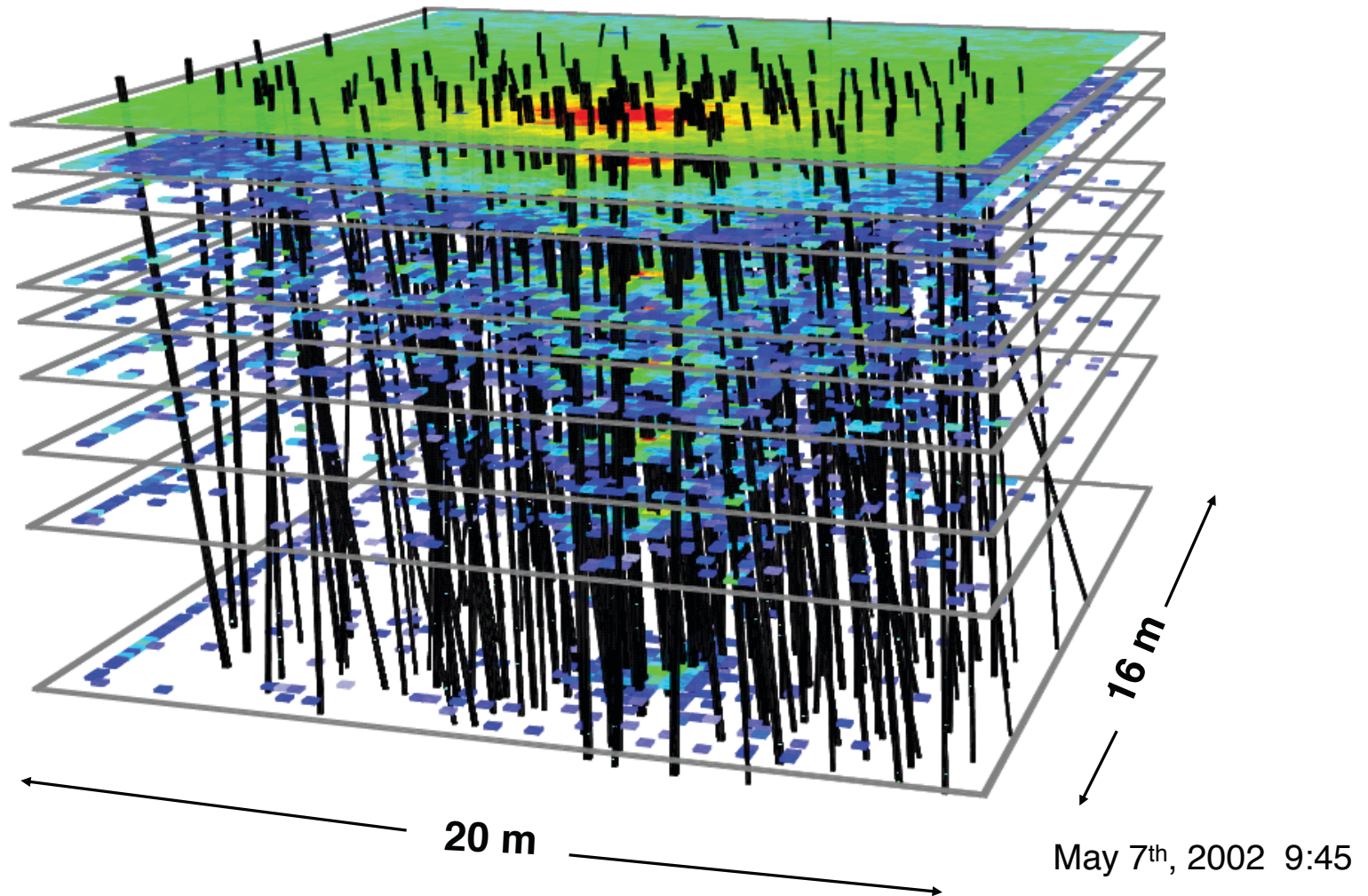




# Hadronic shower core

$E_0 \sim 6 \text{ PeV}$

Number of reconstructed hadrons  $N_h = 143$



to determine the properties of the primary particle,  
we measure the number of

- electrons  $N_e$
- muons  $N_\mu$
- hadrons  $N_h$

--> measure the lateral density distribution  $\rho_{e,\mu,h}(r)$

$$N_{e,\mu,h} = \int_0^\infty 2\pi r \rho_{e,\mu,h}(r) dr$$

we need a suitable parametrization

$$\rho(r) \propto \left(\frac{r}{r_M}\right)^{s-2} \left(1 + \frac{r}{r_M}\right)^{s-4.5}$$

$r_M$ : Molière radius = 0.25  $X_0$

in air ~80 m for electrons (analytically derived)

~400 m for muons (empirical)

~15 m for hadrons (empirical)

# Electron, muon, and hadron lateral distributions measured in air showers by the KASCADE experiment

KASCADE Collaboration

. Perhaps trivial, although experimentally very important, is the fact that in measurements, the shower particles are always sampled over a limited range of core distances  $r_1 < r < r_2$  only – in most cases, with an area coverage in this range not much exceeding 1% – while showers are often referred to in terms of integrated numbers of particles,

$$N = \int_{r_1}^{r_2} 2\pi r \rho(r) dr. \tag{1}$$

The total particle numbers,  $N$ , for different kinds and energy ranges of shower particles are obtained by choosing  $r_1 = 0$  and  $r_2 = \infty$  and are traditionally used both as measures for the primary energy in an individual experiment as well as a means for comparison of different experiments.

Greisen also noted that Eq. (2), except for the last factor, is a close approximation to the analytical calculations for electromagnetic showers performed by Kamata and Nishimura [3] if a shower *age parameter* of  $s = 1.25$  is assumed. Greisen’s approximation to the Nishimura–Kamata functions for  $0.5 < s < 1.5$  is referred to as the NKG function:

$$\rho_{\text{NKG}}(r, s, N_e) = \frac{N_e}{r_M^2} \frac{\Gamma(4.5 - s)}{2\pi\Gamma(s)\Gamma(4.5 - 2s)} \times \left(\frac{r}{r_M}\right)^{s-2} \left(1 + \frac{r}{r_M}\right)^{s-4.5}. \tag{3}$$

This function, often used to describe the charged particle lateral distribution, will in the following be applied individually to electron, muon, and hadron distributions by choosing appropriate sets of parameters  $(s, r_M)$

Deviations of the experimental LDF from the NKG function (Eq. (3)) have been discussed in the literature frequently (Section 1) and are subject to more detailed studies presented below. It turns out that the NKG function can describe the KASCADE electron LDF over the core distance range 10–200 m surprisingly well, but the best agreement is achieved with parameters far away from the conventional assumption of  $r_M \approx 80$  m. When fitting  $N_e$ ,  $r_M$ , and  $s$  simultaneously, the measured LDFs can be reproduced at the 1% level for  $r_M \approx 20$ –30 m and  $s \approx 1.6$ –1.8.

*T. Antoni et al. / Astroparticle Physics 14 (2001) 245–260*

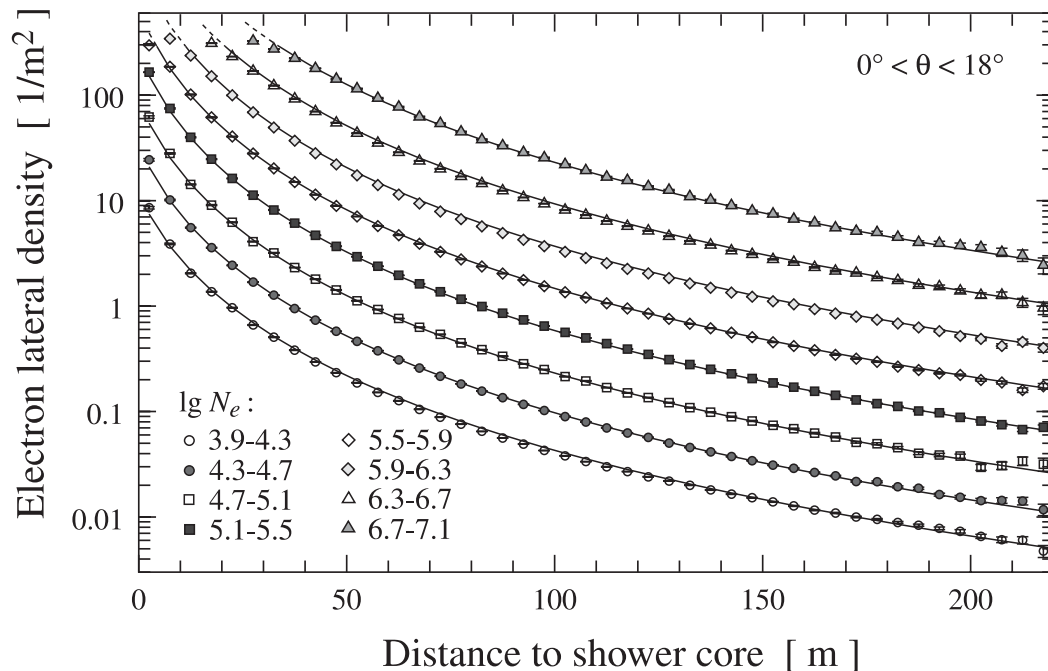


Fig. 2. Lateral distributions of electrons above a 5 MeV kinetic energy for zenith angles below  $18^\circ$ . The lines show NKG functions of fixed age parameter  $s = 1.65$  but varying scale radius  $r_e$  (see the text).

Table 1  
KASCADE detector components used in this analysis<sup>a</sup>

Detector	Channels	Separation (m)	Total area (m <sup>2</sup> )	Threshold $E_{\text{kin}}$	Particle
Array $e/\gamma$	252	13	490	5 MeV	e
Array $\mu$	192	13	622	230 MeV $\times$ sec $\theta$	$\mu$
Trigger	456	–	208	490 MeV $\times$ sec $\theta$	$\mu$
MWPCs	26 080	–	129	2.4 GeV $\times$ sec $\theta$	$\mu$
Calorimeter	38 368	–	304	50 GeV	Hadrons

<sup>a</sup> Detection thresholds refer to the particle energies above the absorber material of the detectors.

## 5. Muon lateral distributions

The KASCADE experiment measures lateral distributions of muons for three different energy thresholds (Table 1). In the following, we group the showers in bins of truncated muon numbers  $N_{\mu}^{\text{tr}}$ . Punch-through and efficiency corrections are applied as described in Section 3. Ranges of core distances for the different muon energy thresholds are limited by the uncertainties in the punch-through corrections at small core distances and by the geometry of the KASCADE detector array. Since  $N_e/N_{\mu}^{\text{tr}}$  rises with the shower size, the impact of punch-through corrections becomes more severe at higher energies and the minimum core distances have to be increased correspondingly. For showers with cores inside KASCADE, the upper limit is about 220 m for array detectors and 100 m for central detector components.

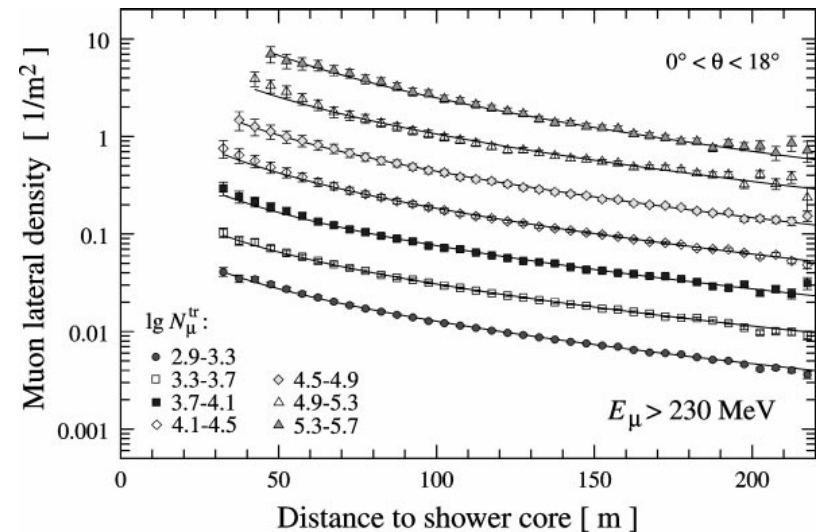


Fig. 7. Lateral distribution of muons above 230 MeV kinetic energy measured with the array detectors. The lines indicate NKG functions fitted to the data. Error bars are of statistical nature including an uncertainty of 10% on the punch-through correction applied.

## NKG functions with $r_{\mu} = 420$ m

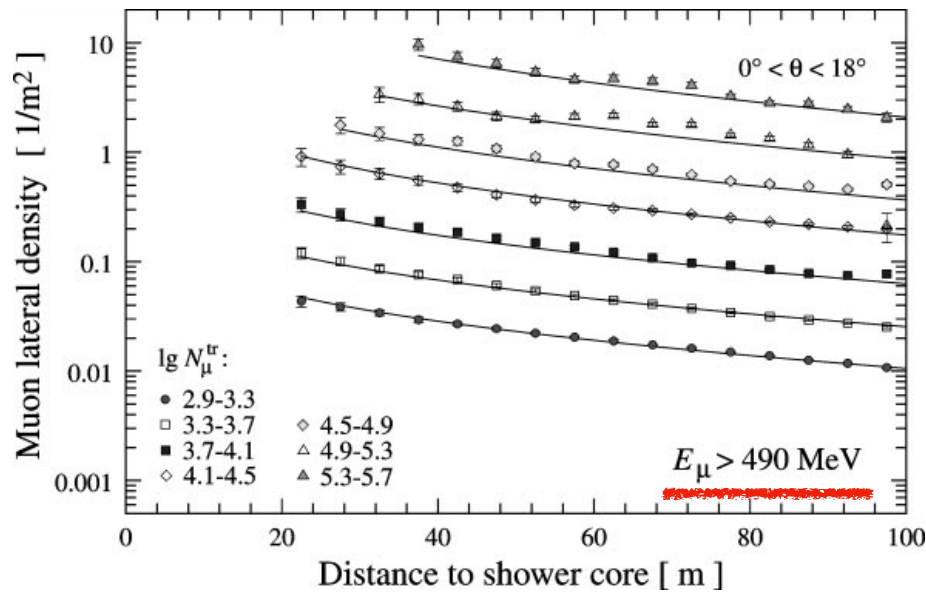


Fig. 8. Lateral distribution of muons above 490 MeV kinetic energy as measured with the trigger plane. Data are binned according to  $N_{\mu}^{\text{tr}}$  as measured by the array and the lines represent NKG fits to the data.

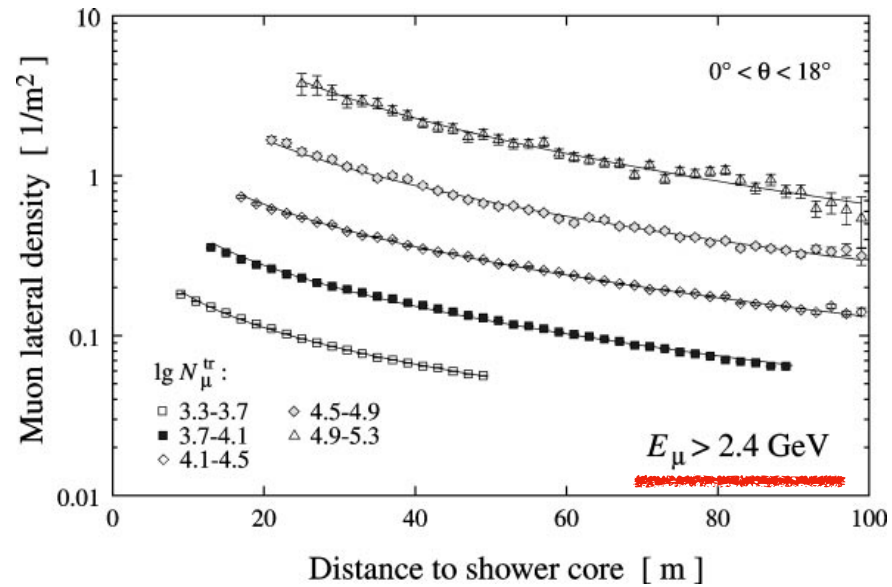


Fig. 9. Lateral distribution of muons above 2.4 GeV kinetic energy. Error bars are statistical only. Data are binned according to  $N_{\mu}^{\text{tr}}$  as measured by the array and the lines represent NKG fits to the data.

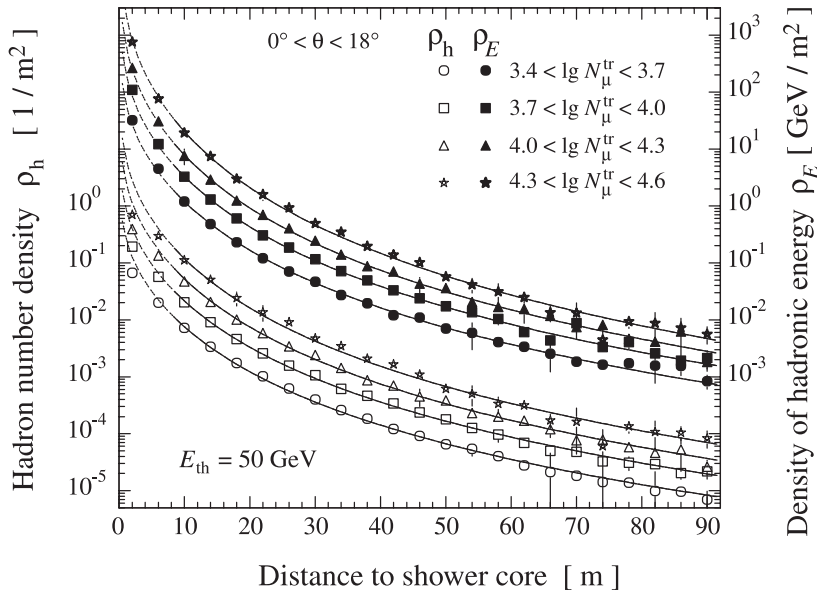


Fig. 12. Density of hadron number (left scale, open symbols) and of hadronic energy (right scale, filled symbols) versus the core distance for showers of truncated muon numbers as indicated. Threshold energy for hadrons is 50 GeV. The curves represent fits of the NKG formula to the data at  $r \geq 8$  m with a radius fixed to  $r_h = 10$  m.

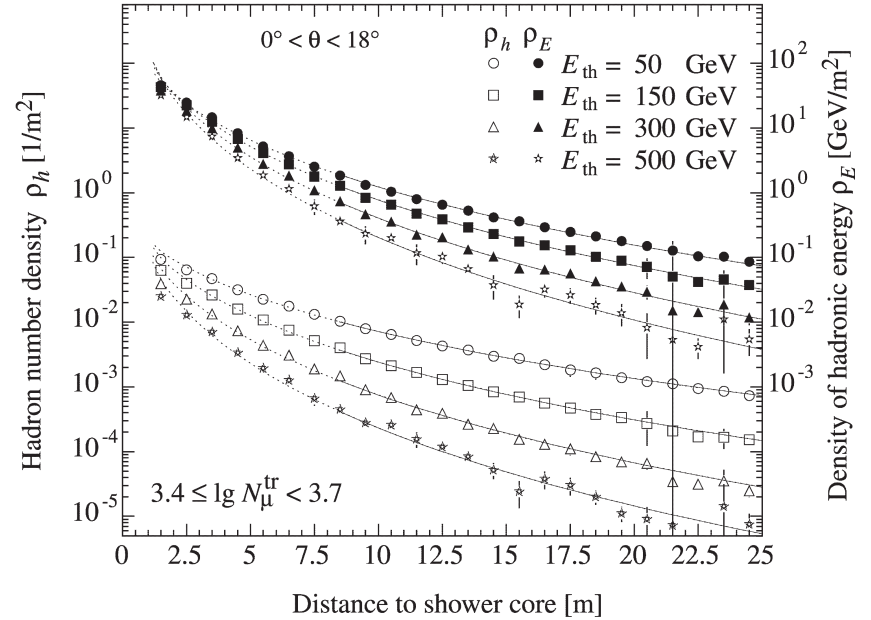


Fig. 14. Density of hadron number (left scale, open symbols) and of hadronic energy (right scale, filled symbols) versus shower core distance for various thresholds of hadron energy. The curves represent fits of the data to the NKG function as in Fig. 12.

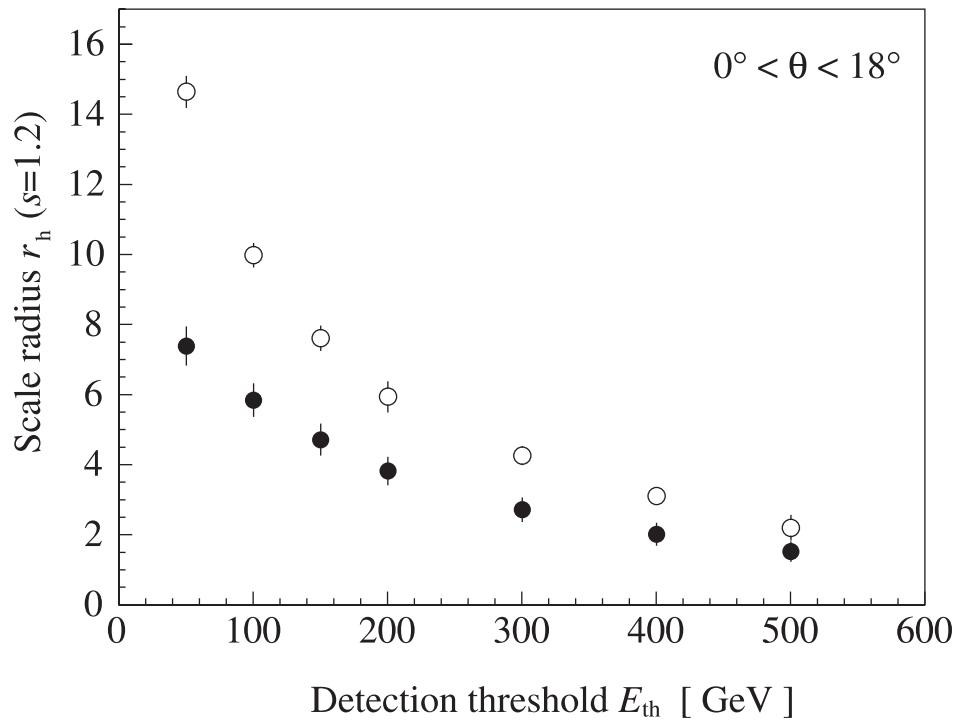


Fig. 15. Hadronic scale radius  $r_h$  as a function of detection threshold  $E_{th}$  for a fixed shower age of  $s = 1.2$  and  $3.4 < \lg N_{\mu}^{tr} < 3.7$ . Filled symbols are for hadronic energy density and open symbols for hadronic particle density.

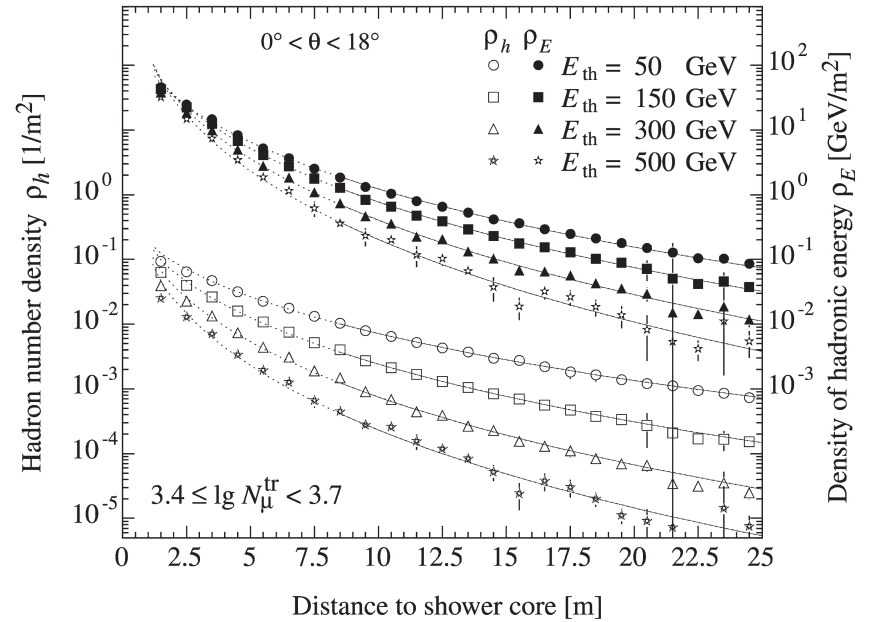


Fig. 14. Density of hadron number (left scale, open symbols) and of hadronic energy (right scale, filled symbols) versus shower core distance for various thresholds of hadron energy. The curves represent fits of the data to the NKG function as in Fig. 12.



The core position of the shower is found by fitting the signal  $S(r)$  of the detector stations with a lateral distribution function (LDF) tailored for the response of the particular array. The parameters of the lateral distribution can be determined from data directly using vertical showers. Because the NKG function (Eq. 15.35) was developed to describe EM showers only, various modified versions have been developed to describe the charged particles in more complex hadronic showers. An example of a modified NKG form is

$$\rho_e(m^{-2}) = f(x) = C_1(s) x^{(s-2)} (1+x)^{(s-4.5)} (1+C_2 x^d), \quad (16.29)$$

where  $x = r/r_1$ , and the normalization constant is<sup>3</sup>

$$C_1(s) = \frac{N_e}{2\pi r_1^2} [B(s, 4.5 - 2s) + C_2 B(s + d, 4.5 - d - 2s)]^{-1}.$$

The scale for the NKG lateral distribution function is the Molière unit, which at the depth of the Akeno array of 920 g/cm<sup>2</sup>, for example, is  $r_1 \approx 85$  m. For showers of size  $N_e \approx 10^6$  at sea level, Greisen [538] uses  $s = 1.25$ ,  $d = 1$  and  $C_2 = 0.088$ . Nagano et al. [539] at 920 g/cm<sup>2</sup> use  $d = 1.3$  and  $C_2 = 0.2$  with  $s$  fitted for each shower. The modified NKG form used for the surface array of Auger is

$$S(r) = \tilde{C} \left(\frac{r}{r_s}\right)^{-\beta} \left(1 + \frac{r}{r_s}\right)^{-\beta}. \quad (16.30)$$

A list of often-used LDF parametrizations can be found in [540]. In general, one of the fitting parameters accommodates fluctuations in shape while another characterizes the density at a particular distance,  $r_s$ .

For arrays that can separate muons from  $e^\pm$  it is useful to have a separate lateral distribution function tailored for the somewhat flatter muon distribution. A standard form is due to Greisen [538]:

$$\rho_\mu(\text{m}^{-2}) = \frac{\Gamma(2.5)}{2\pi\Gamma(1.25)\Gamma(1.25)} \left(\frac{1}{320}\right)^{1.25} N_\mu r^{-0.75} \left(1 + \frac{r}{320}\right)^{-2.5}. \quad (16.31)$$

The distribution is normalized so that  $N_\mu$  is the total number of muons in the shower at the surface.

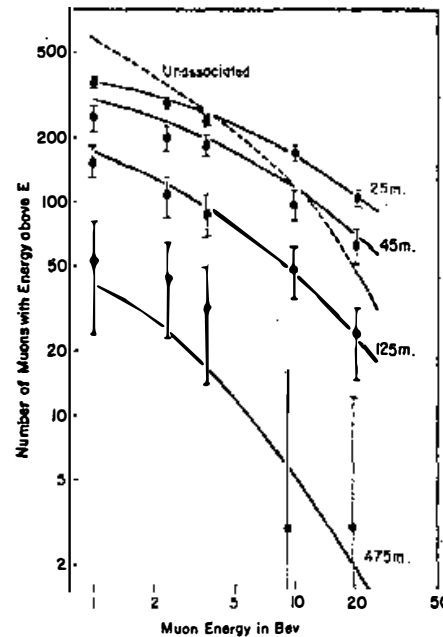
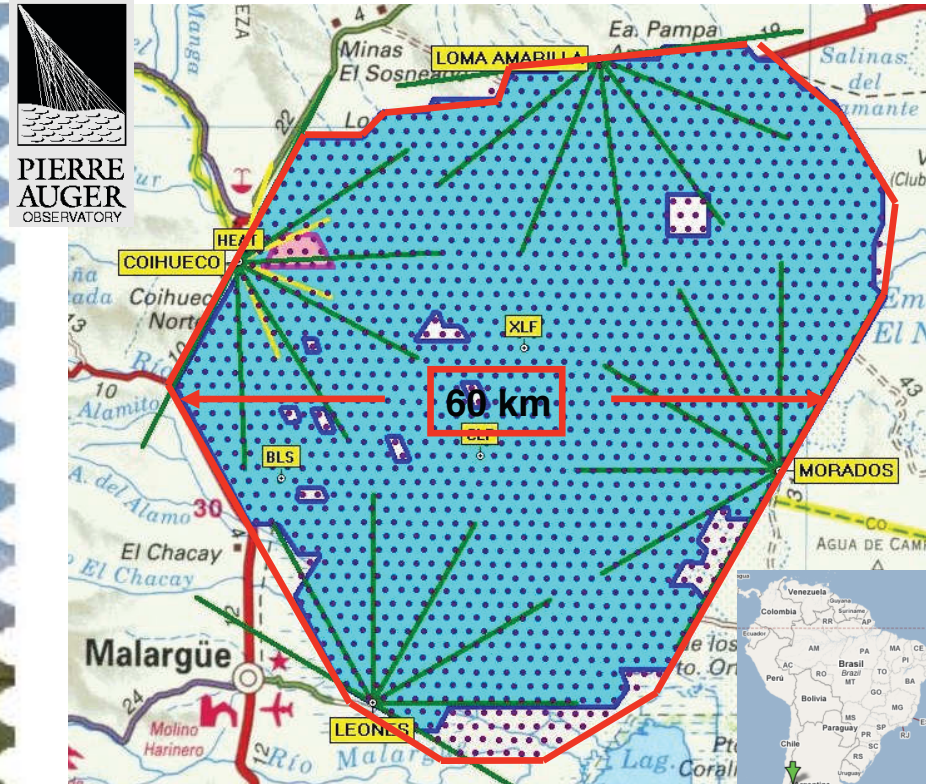


FIG. 4. Mu-meson energy spectra at various distances from extensive air-shower axes, measured by Bennett (8). The data cover a wide range of shower sizes,  $10^3 < N < 10^7$ , and each solid curve refers to a considerable interval of radius, with mean values indicated. The dashed curve refers to mesons that were not coincident with detected showers. The solid curves were computed with the following expression for the density of mesons (per sq. m.) with energy above  $E$  (Bev) at distance  $r$ (m.) from the axis of showers containing  $N$  charged particles:

# The Pierre Auger Observatory



3000 km<sup>2</sup>

1661 water Cherenkov detectors

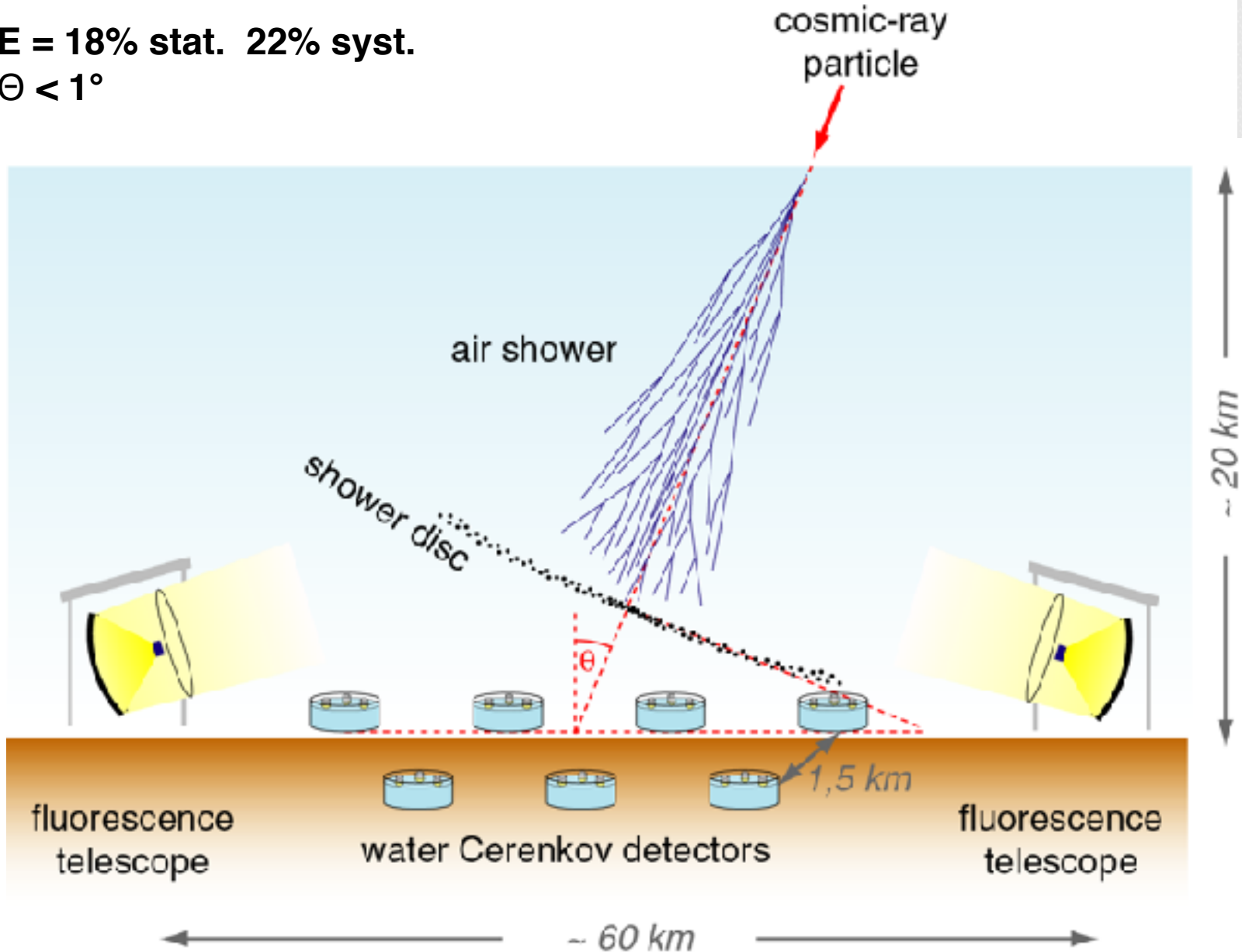
24+3 Fluorescence telescopes

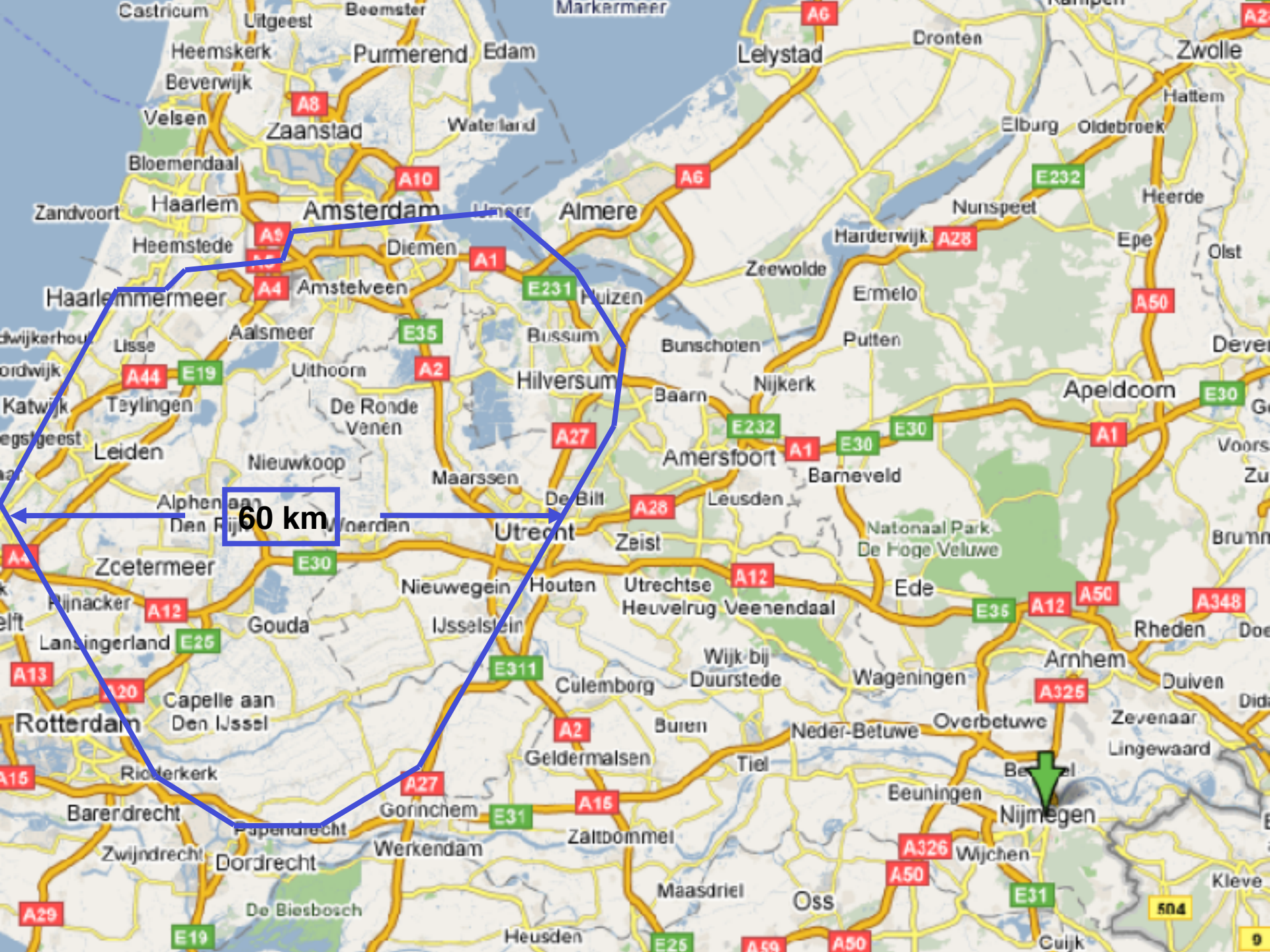
~150 radio antennas

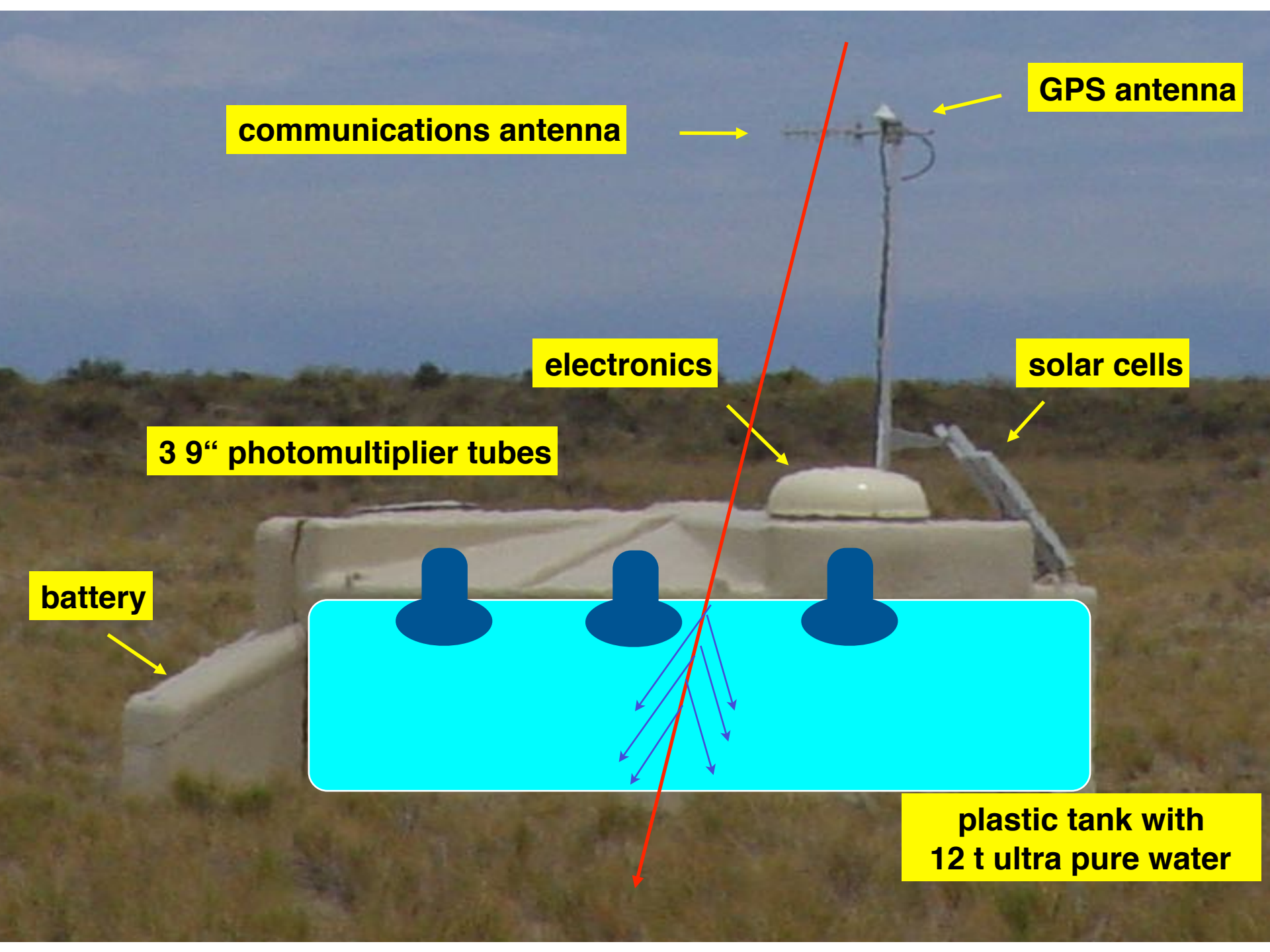
# The detection principle

$\Delta E = 18\%$  stat. 22% syst.

$\Delta\theta < 1^\circ$







**communications antenna**

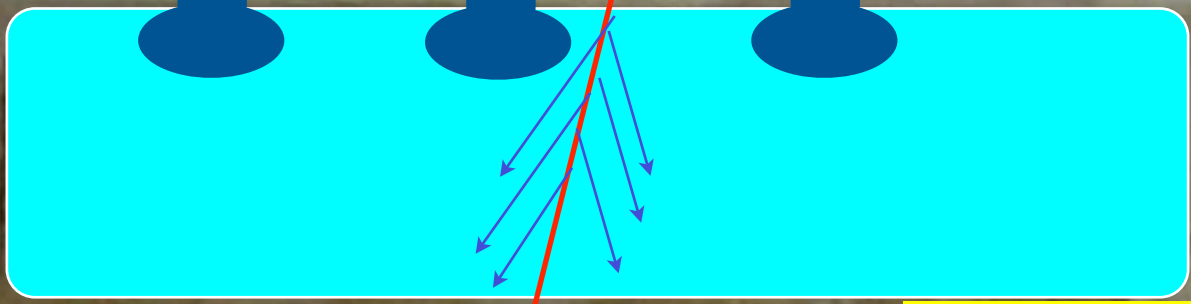
**GPS antenna**

**electronics**

**solar cells**

**3 9" photomultiplier tubes**

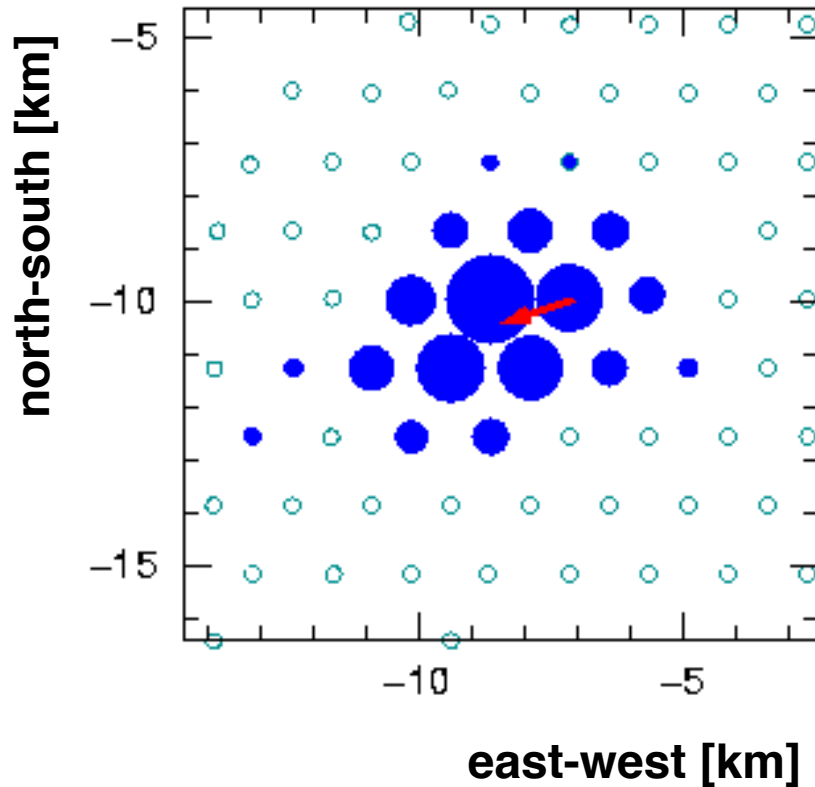
**battery**



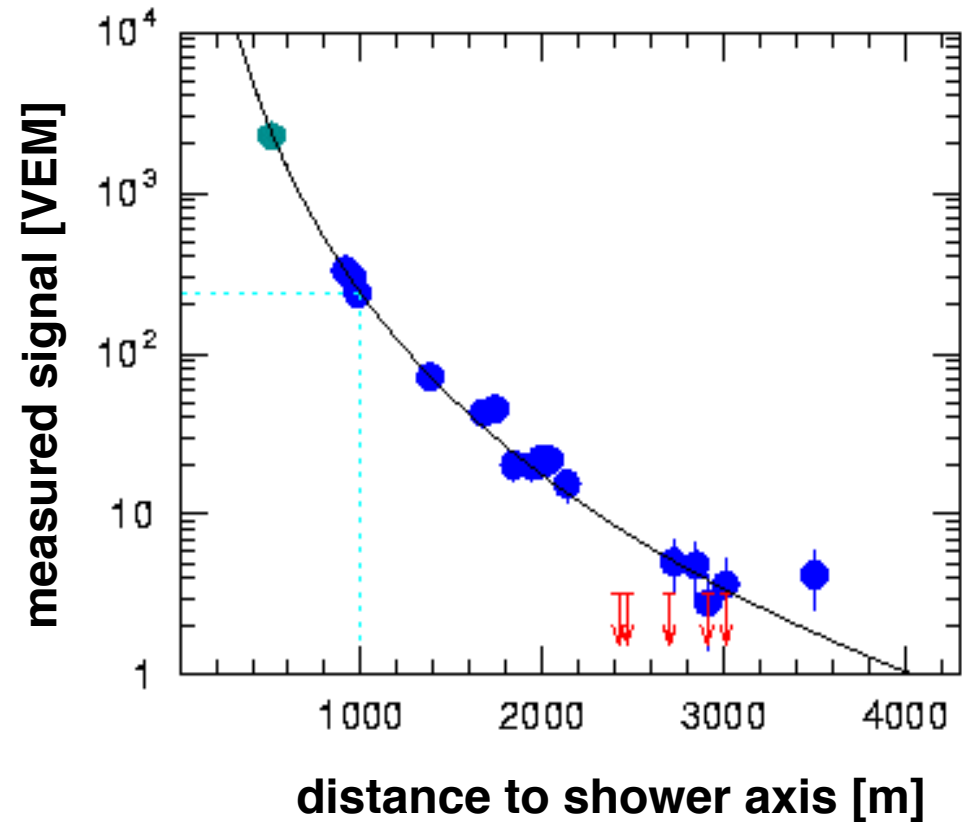
**plastic tank with  
12 t ultra pure water**

# Air shower registered with water Cherenkov detectors

ID 762238



ID 762238



To reconstruct the energy of a shower, a measured ground parameter is defined. Examples are the number of detected particles at ground calculated by integrating the lateral distribution or a signal density at a specific lateral distance. The latter is illustrated in Figure 16.5 for the Auger surface detector array. The measured signals of the detector stations of one particular event are reconstructed with the LDF of Eq. 16.30 with different values of  $\beta$ . A fix point is found at a core distance of about  $r_{\text{opt}} = 1100 \text{ m}$  [541]. The signal (i.e. particle density) obtained for this distance is independent of the details of the LDF used for reconstruction and, hence, can be used as a robust estimator for determining the shower energy through comparison with Monte Carlo reference showers or cross-calibration with other calorimetric energy measurements. The optimum distance depends mainly on the spacing of the detectors and is not related to shower-to-shower fluctuations.



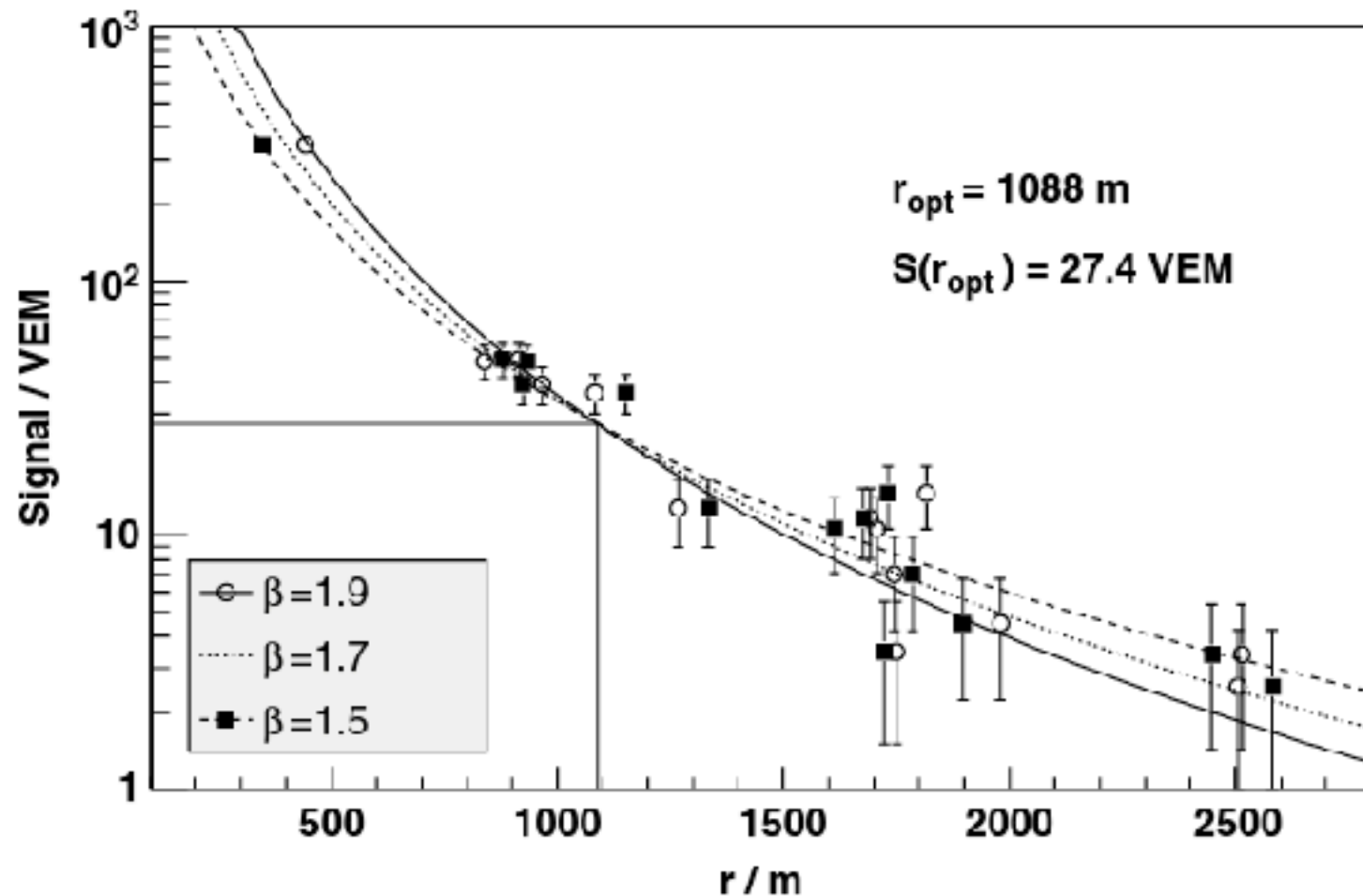


Figure 16.5 Example of the determination of the optimum distance for measuring the particle density of an air shower in the Pierre Auger Observatory (see text). The detector signal is expressed in units of the signal expected for vertical muons (vertical equivalent muons, VEM). From [541].

# particle type from air-shower measurements

The most promising surface-detector approach is the separate measurement of the number of electrons and muons. The corresponding predictions for air showers simulated with the hadronic interaction models EPOS [146, 161] and Sibyll [158] (interactions with  $E > 80 \text{ GeV}$ ) and FLUKA [193] (interactions with  $E \leq 80 \text{ GeV}$ ) are shown in Figure 16.6. The simulation results confirm the predictions of the superposition model (16.19) with  $N_\mu \propto E_0^\alpha$ ,  $\alpha \approx 0.90$  and a relative difference in the muon number between iron and proton showers of  $\sim 40\%$ . The difference in the number of electrons at each energy is mainly related to the shallower depth of shower maximum of iron showers relative to proton showers.

With the energy transferred to the EM shower component being closely related to (and anti-correlated with) the number of muons at ground, one can devise an almost model-independent estimator for the primary energy

$$E_0 = E_{\text{em}} + E_{\text{had}} \approx \bar{E}_e N_e^{(\text{max})} + E_{\text{dec}} N_\mu, \quad (16.32)$$

where  $\bar{E}_e > E_e$  is a typical energy scale one has to assign to electrons to compensate for the non-detected photons. In practical applications, the energy is parametrized as  $\ln E = a \ln N_e + b \ln N_\mu + c$ , with  $a, b, c$  being parameters determined from simulations. A similar expression can be written for  $\ln A$  to find the primary mass; see [540, 544]. Depending on the distance of the observation level to the depth of the typical shower maximum, fluctuations in the particle numbers can be large and need to be accounted for in energy and composition reconstruction.

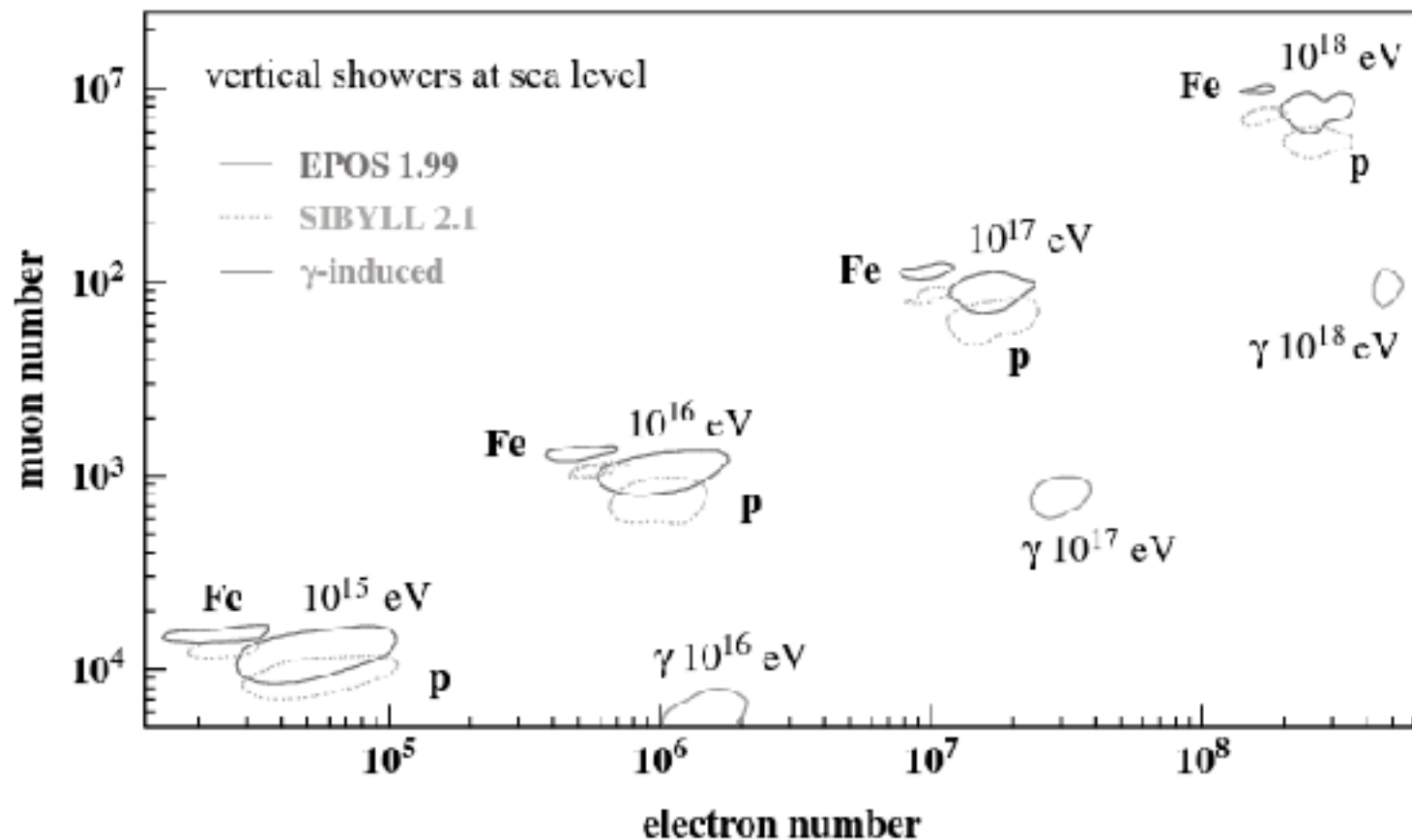
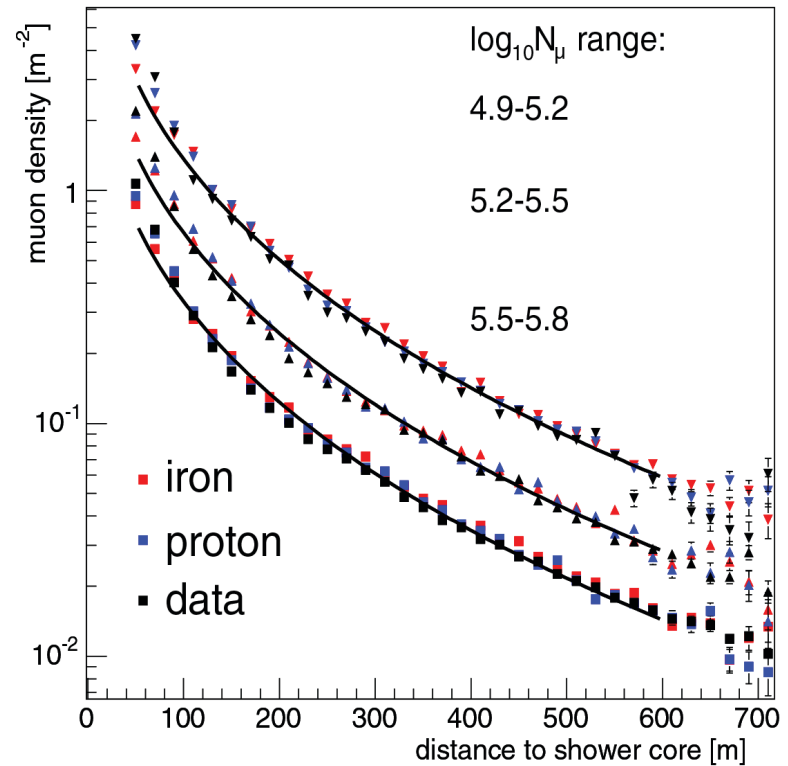
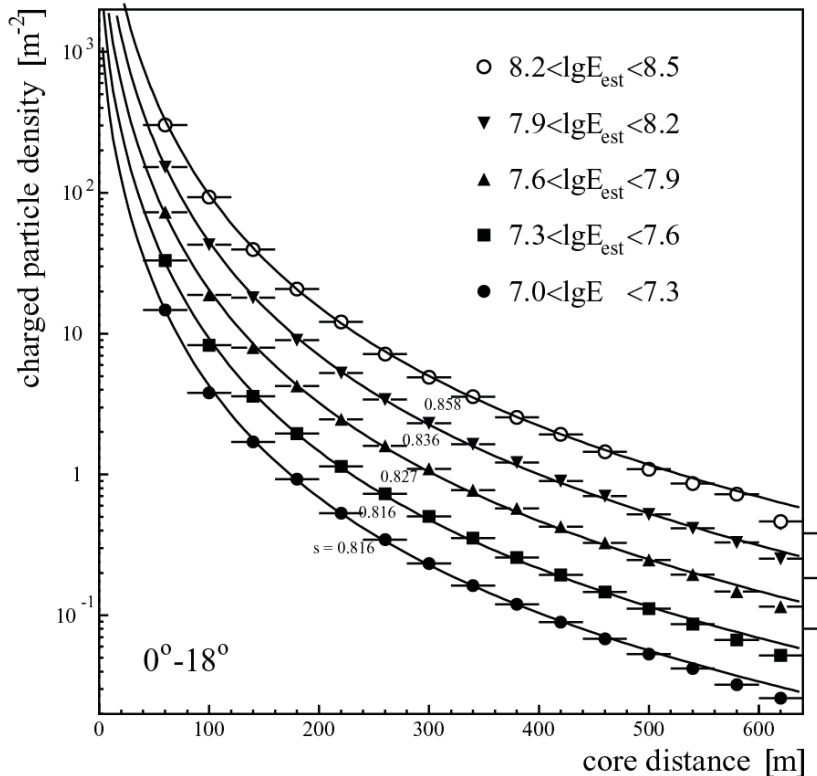


Figure 16.6 Predicted correlation between the number of muons and electrons of vertical showers at sea level. The simulations were done with CORSIKA [640] using the same cutoff energies for the secondary particles as in Figure 16.1. The curves encircle approximately the one-sigma range of the fluctuations. From [33], © 2011 by Annual Reviews ([www.annualreviews.org](http://www.annualreviews.org)), reproduced with permission.

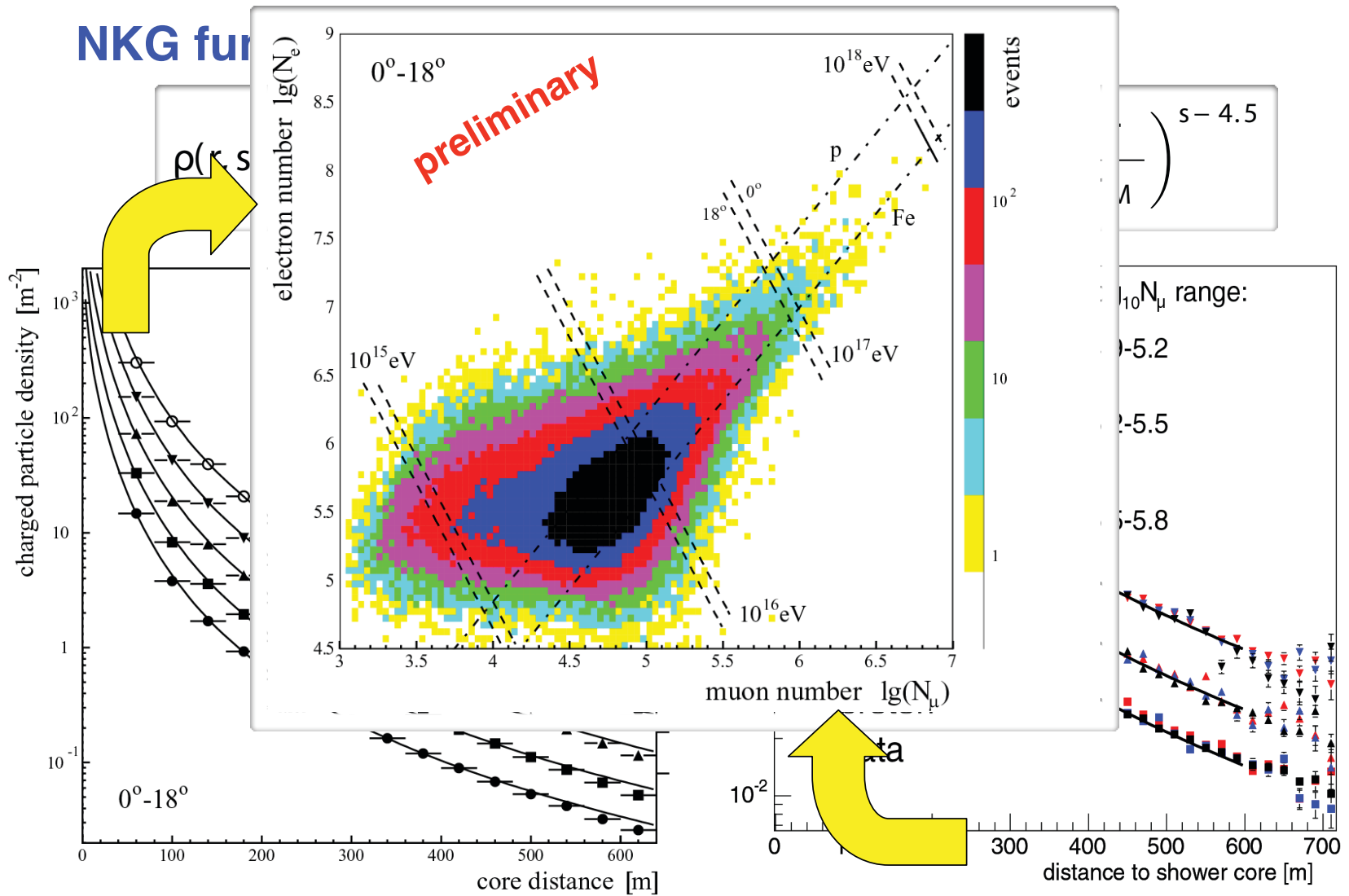
# KASCADE-Grande – Lateral distributions

## NKG function

$$\rho(r, s, N_e) = \frac{N_e}{r_M^2} \frac{\Gamma(4.5 - s)}{2\pi\Gamma(s)\Gamma(4.5 - 2s)} \left(\frac{r}{r_M}\right)^{s-2} \left(1 + \frac{r}{r_M}\right)^{s-4.5}$$



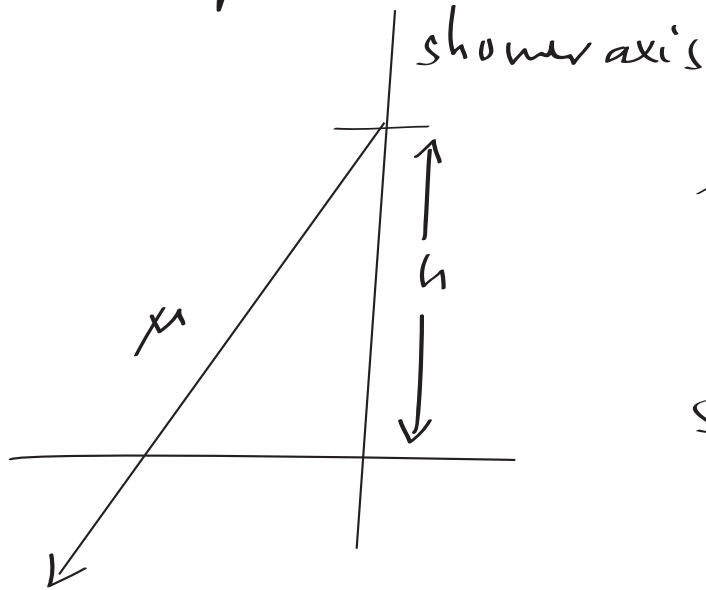
# KASCADE-Grande – Lateral distributions



to determine the mass of the primary particle

•  $\frac{e}{\mu}$  ratio or  $\frac{h}{\mu}$  ratio

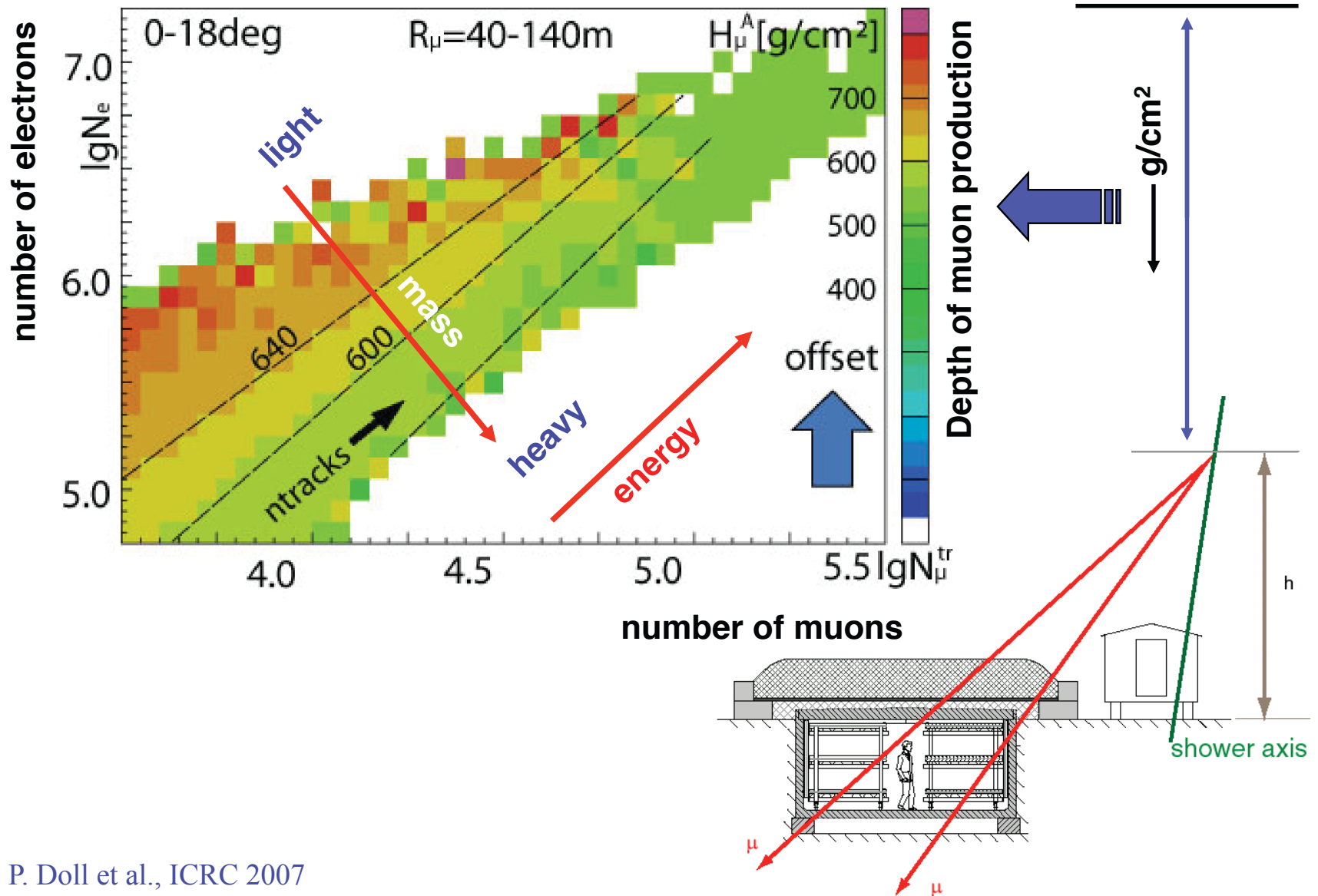
• maximum production height or  $X_{max}$  for  $e/\mu$  comp.



large  $h \rightarrow$  heavy primary particle

small  $h \rightarrow$  light particle

# Muon production height – KASCADE muon tracking detector



P. Doll et al., ICRC 2007

## 16.8 Atmospheric Cherenkov light detectors

The large number of Cherenkov photons emitted by charged particles traversing a medium with refractive index  $n > 1$  can be used for efficient detection of air showers in a wide range of energies. The atmospheric Cherenkov technique is useful both for  $\gamma$ -ray astronomy and for study of cosmic ray air showers. Imaging atmospheric Cherenkov telescopes (IACTs) can detect showers with thresholds down to 30 GeV [566, 567]. Their reach at high energy, however, is limited to  $\sim 100$  TeV by their relatively small effective area. Non-imaging Cherenkov detectors can be set up similar to an array of particle detectors, offering the possibility to instrument very large areas at ground and reach very high energy [568, 569]. Typically only the Cherenkov light of the abundant secondary particles in an air shower is detected, but also the direct Cherenkov light of the primary particle can be measured [570, 571]. From the point of view of cosmic ray physics, the advantage of a Cherenkov array is that it can reconstruct depth of shower maximum at energies lower than those accessible to fluorescence telescopes.

- non-imaging detector
- Imaging Atmospheric Cherenkov Telescope (IACT)



The phenomenology of atmospheric Cherenkov emission is largely determined by how the emission angle and energy threshold depend on the index of refraction as a function of altitude. It is convenient to express the threshold of particle energy  $E$  for Cherenkov light emission in terms of the Lorentz  $\gamma$ -factor

$$\gamma \geq \frac{n(h)}{\sqrt{n(h)^2 - 1}}, \quad (16.33)$$

with  $E = \gamma m$  and  $m$  the particle mass. The altitude dependence of the refractive index  $n(h)$  is a function of the local air density and satisfies approximately

$$n(h) = 1 + 0.000283 \frac{\rho_{\text{air}}(h)}{\rho_{\text{air}}(0)}, \quad (16.34)$$

where  $\rho_{\text{air}}$  is the density of air. The energy threshold for electrons and the Cherenkov angle  $\theta_{\text{Ch}}$  in air,  $\cos \theta_{\text{Ch}} = 1/(\beta n(h))$ , are given in Table A.2 as a function of altitude. Typical values at  $h = 10\text{km}$  are  $\theta_{\text{Ch}} = 0.8^\circ$  (12 mrad) and a threshold of  $\gamma = 72$ , corresponding to  $E = 37\text{MeV}$  for electrons and  $E = 7.6\text{GeV}$  for muons.

The Cherenkov light cone of a particle at 10 km height has a radius of about 120 m at ground. This means that most of the light is expected within a circle of this radius. Due to multiple Coulomb scattering the shower particles do not move parallel to the shower axis. The angular distribution follows in first approximation an exponential

$$\frac{dN_\gamma}{d\theta} = \frac{1}{\theta_0} e^{-\theta/\theta_0}, \quad \theta_0 = 0.83 \left( \frac{E_{\text{th}}}{\text{MeV}} \right)^{-0.67}, \quad (16.35)$$

with  $E_{\text{th}}$  being the Cherenkov energy threshold [527, 532]. Typical values of  $\theta_0$  are in the range  $4 \dots 6^\circ$ . The interplay of the altitude-dependent Cherenkov angle and the emission height leads to a characteristic lateral distribution of photons at the ground, as illustrated in Figure 16.7. The absorption and scattering of Cherenkov light in the atmosphere limits the detectable wavelength range to about 300 – 450 nm, where the upper limit follows from the  $\lambda^{-2}$  suppression of large wavelengths. One possible parametrization of the lateral distribution of the Cherenkov light has the form [574]

in addition to particles at ground (electrons, muons, hadrons) in air showers **electromagnetic radiation** is produced:

- **Cherenkov radiation**

charged particles moving with a speed  
emitt Cherenkov radiation

$$v > v_{air} = \frac{c}{n_{air}}$$

air close to ground

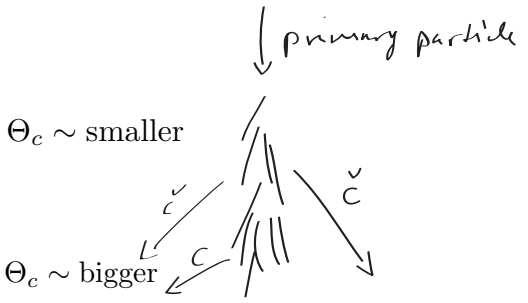
$$n \approx 1.000283 \quad \cos \Theta_c = \frac{c}{n \cdot v} \approx \frac{1}{n}$$

$\Rightarrow \Theta_c < 1.4^\circ$  in air

**Cherenkov threshold**

$$E_e = 22 \text{ MeV} \quad E_\mu = 4.4 \text{ GeV}$$

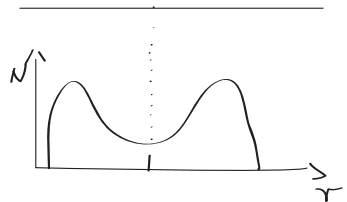
--> most Cherenkov light is emitted from relativistic electrons



focusing effect due to increase of  $n$  with increasing depth -->  $n(h)$

$$\Rightarrow \Theta_c = f(h)$$

ring with ~100 m diameter on the ground



e.g. 1 TeV gamma ray produces  $\sim 3 \cdot 10^6$  Cherenkov photons 300-500 nm

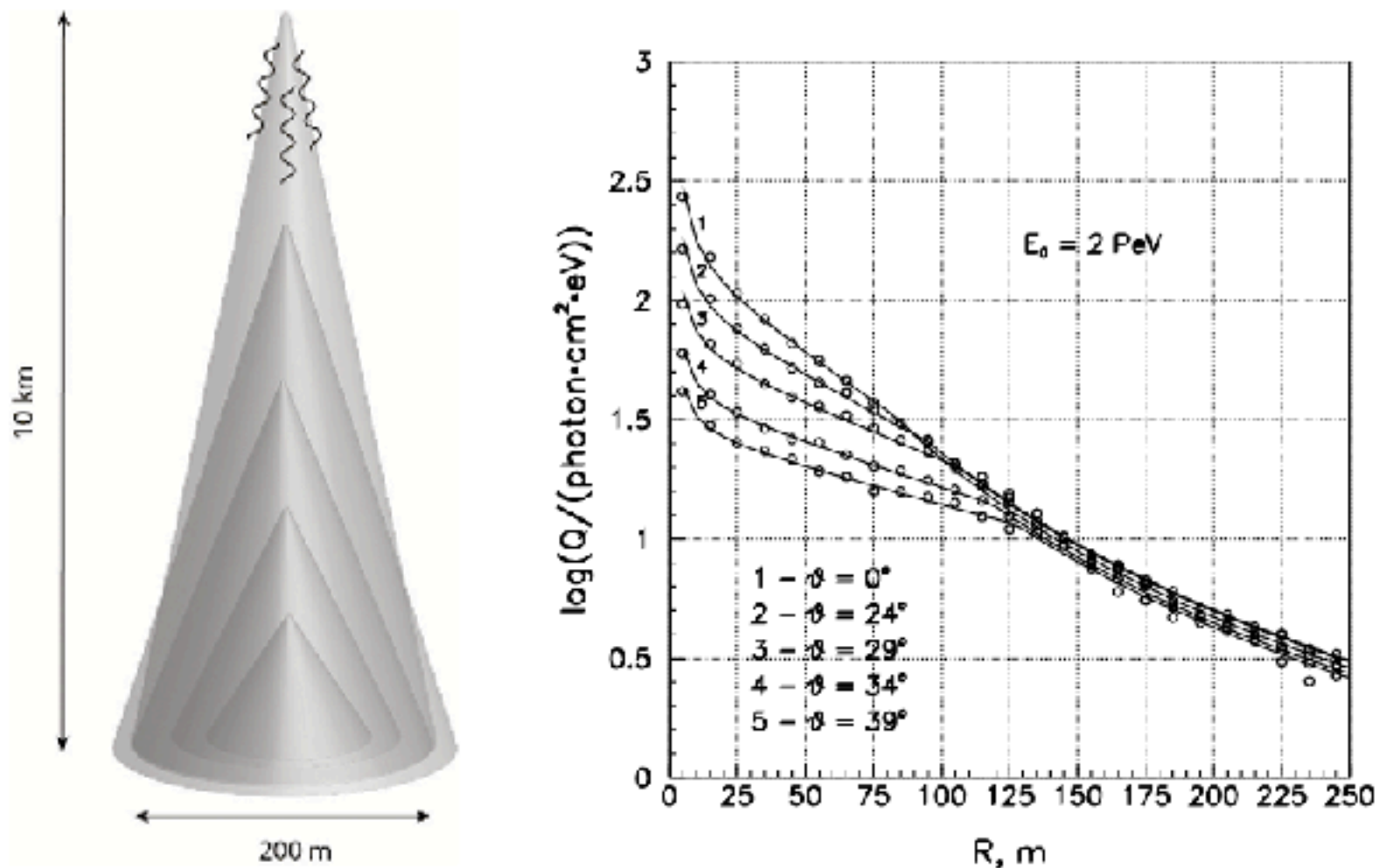


Figure 16.7 Left: Illustration of the relation between production height and Cherenkov opening angle for producing the observed Cherenkov light distribution at ground. Right: Simulated lateral distributions of Cherenkov light produced by proton-induced showers of different zenith angle [572]. The simulations were done for a height of 2000 m above sea level.

One possible parametrization of the lateral distribution of the Cherenkov light has the form [574]

$$C(r) = \begin{cases} C_{120} \cdot \exp(a[120 \text{ m} - r]); & 30 \text{ m} < r \leq 120 \text{ m} \\ C_{120} \cdot (r/120 \text{ m})^{-b}; & 120 \text{ m} < r \leq 350 \text{ m} \end{cases}, \quad (16.36)$$

with the parameters  $C_{120}$ ,  $a$  and  $b$ .

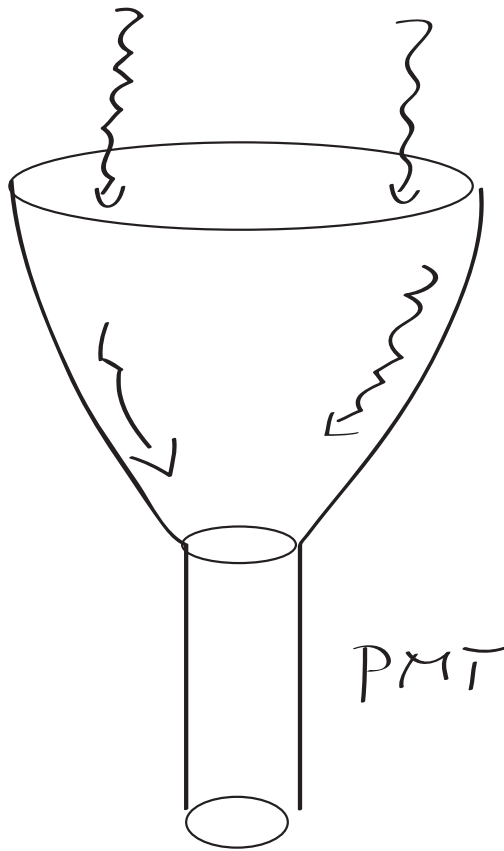
Clear, moonless nights are required for taking data with air Cherenkov detectors, resulting in an effective duty cycle of 10 – 15%. Also continuous monitoring of the atmospheric conditions including the density profile of the atmosphere is necessary [575].

Arrays of photodetectors are used in non-imaging Cherenkov experiments to sample the lateral distribution of light in dark and clear nights. After reconstructing the core position, the measured parameter  $C_{120}$  and the slope are linked to the properties of the primary particle. Simulations show that the density of photons at 120 m from the core is almost directly proportional to the energy of the shower and that the slope is related to the depth of shower maximum [573]. Examples of surface arrays applying this non-imaging technique of shower detection via Cherenkov light are AIROBICC [576], EAS-TOP [577], BLANCA [574], Tunka [568] and Yakutsk [569]. The latter two are currently in operation, with Tunka being extended from an array of originally 25 stations to 133.

# Measurement of Čerenkov light

two techniques

1) non-imaging detectors      open  $\checkmark$  detectors



$\phi \sim 0,5m$

Winston cone  
to collect light

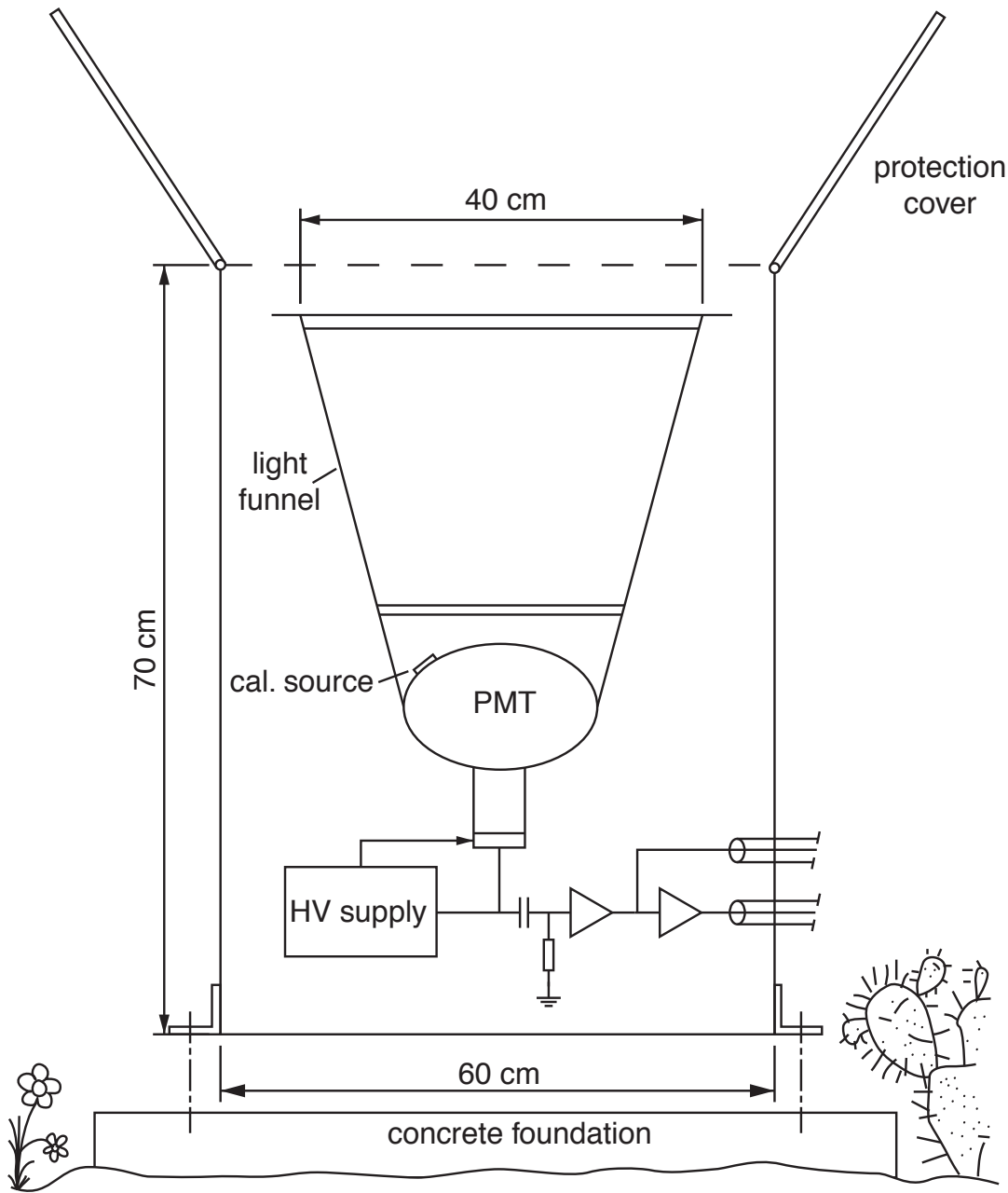


Fig. 4.19. Integrating Cherenkov cone of an AIRO-BICC station and auxiliaries. Directly above the PMT a glass filter restricted the incoming light to wavelengths smaller than 500 nm and a plexiglass cover protected against dew, white frost and dust [29]

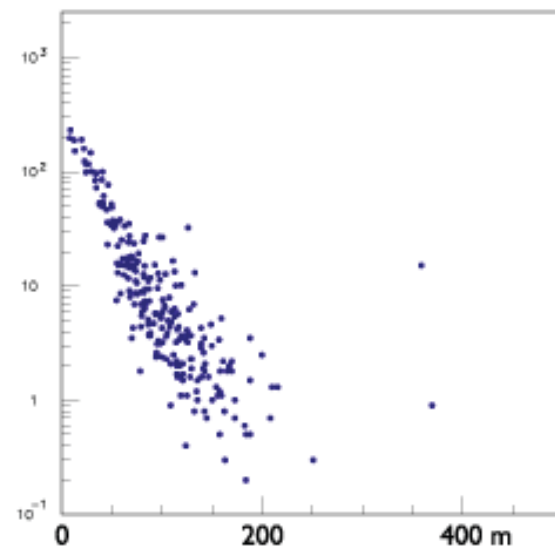
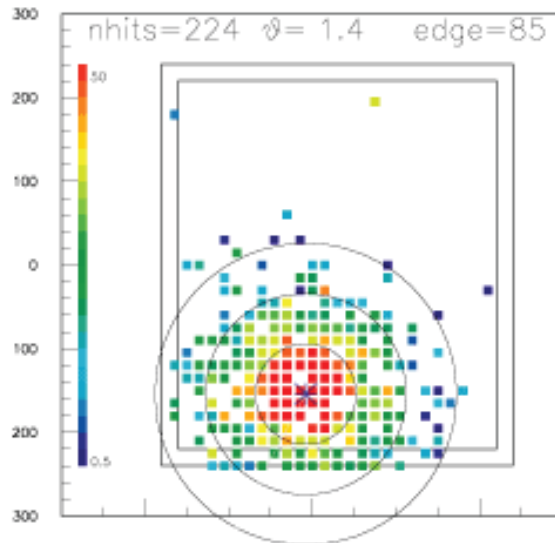




# Single event in CASA-BLANCA

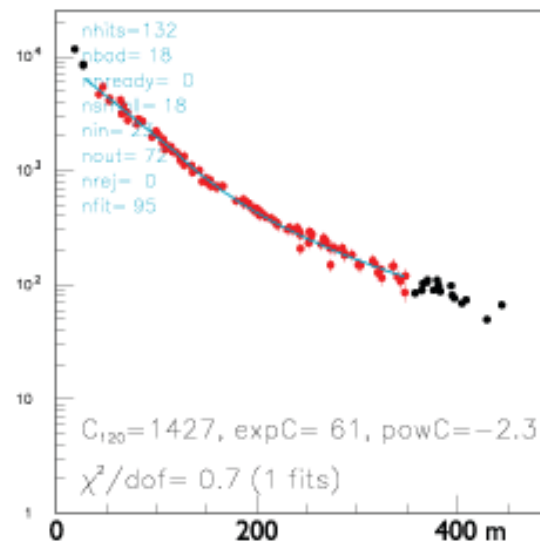
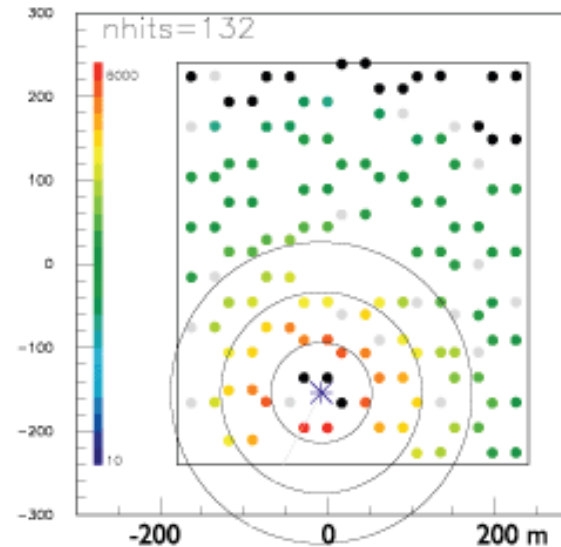
## CASA

measures particles,  
determines shower core



## BLANCA

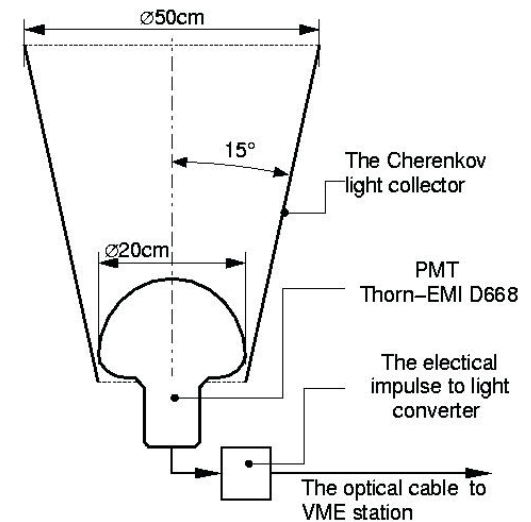
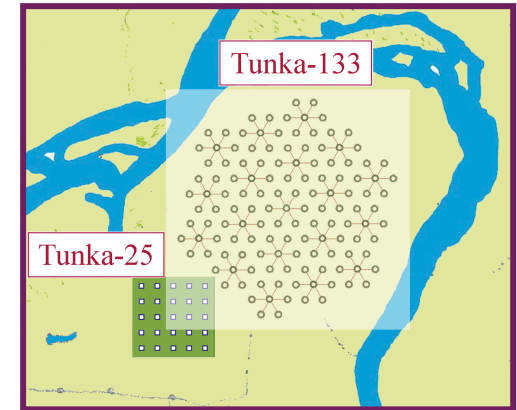
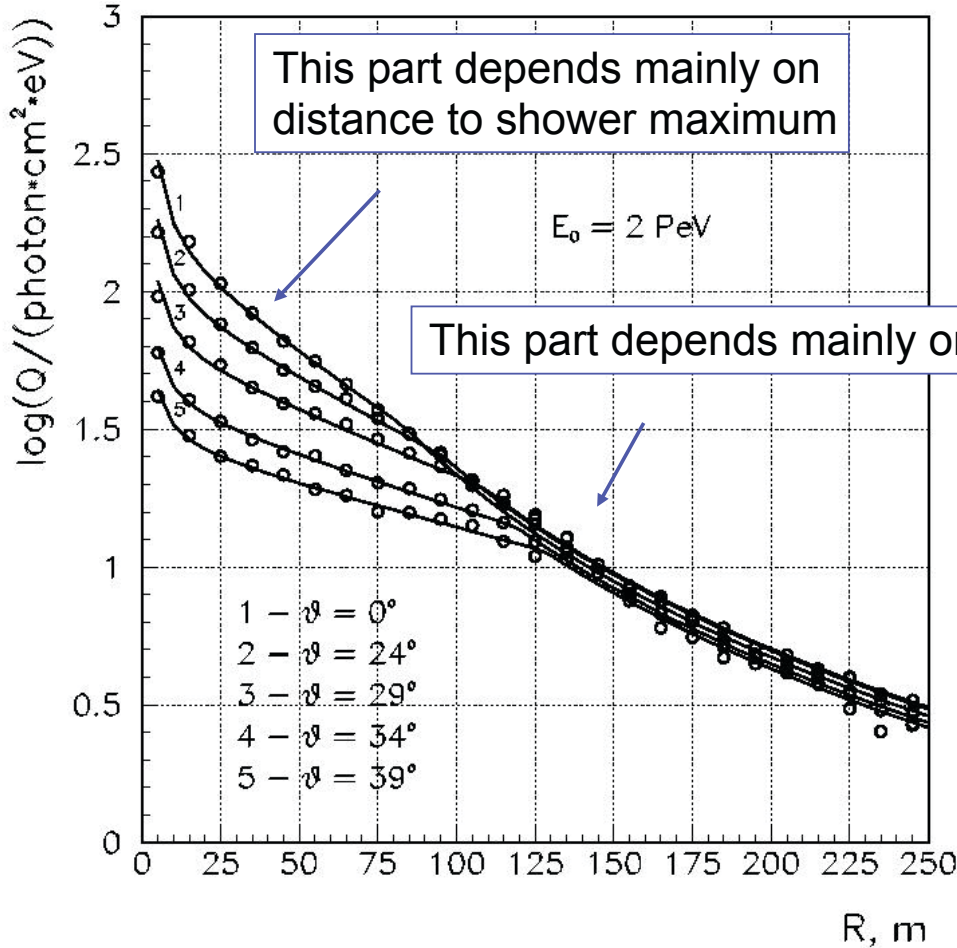
measures Cherenkov light,  
determines lateral C-distribution



# Tunka Experiment

## Lateral distribution of Cerenkov light

$$C(r) = \begin{cases} C_{120} \times \exp(s[120\text{ m} - r]) & 30\text{ m} < r < 120\text{ m} \\ C_{120} \times (r/120\text{ m})^{-\hat{a}} & 120\text{ m} < r < 350\text{ m} \end{cases}$$





# Imaging Atmospheric Cherenkov Telescope (IACT)

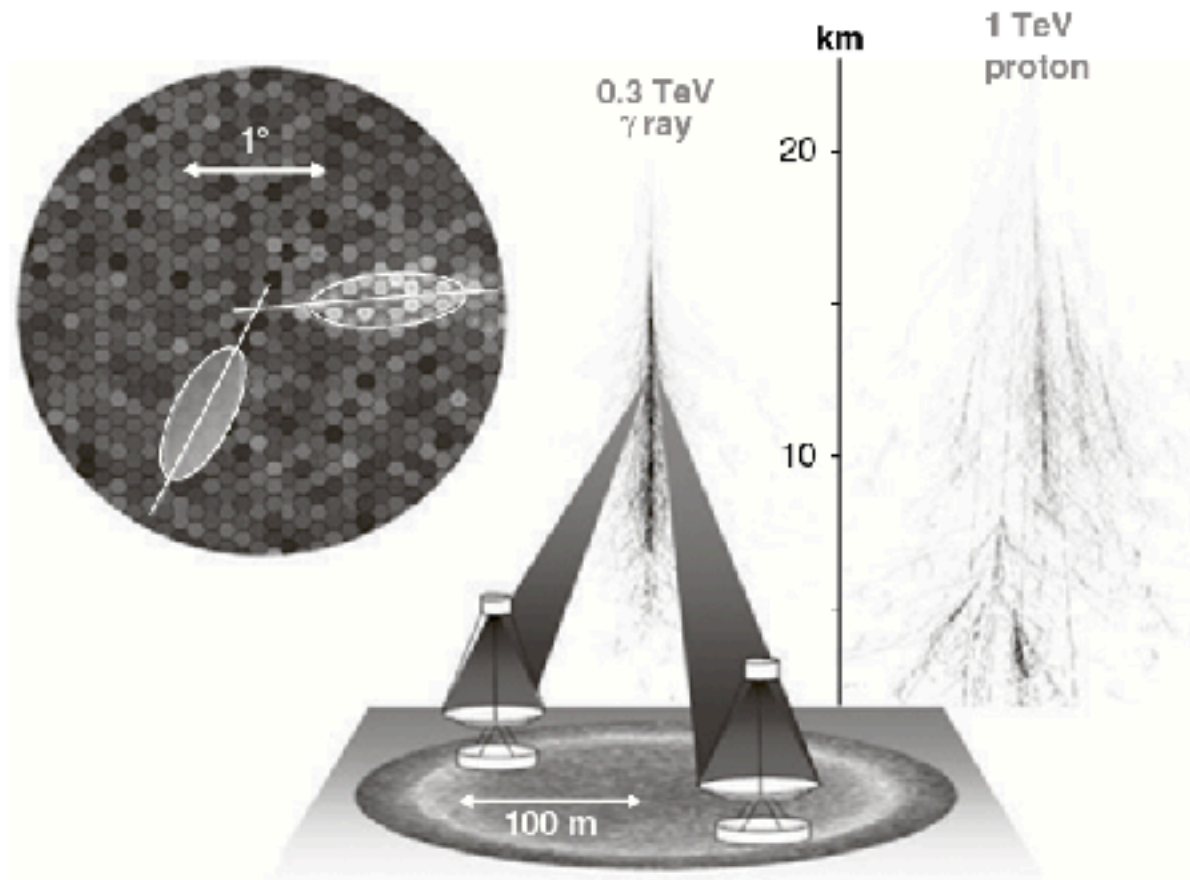
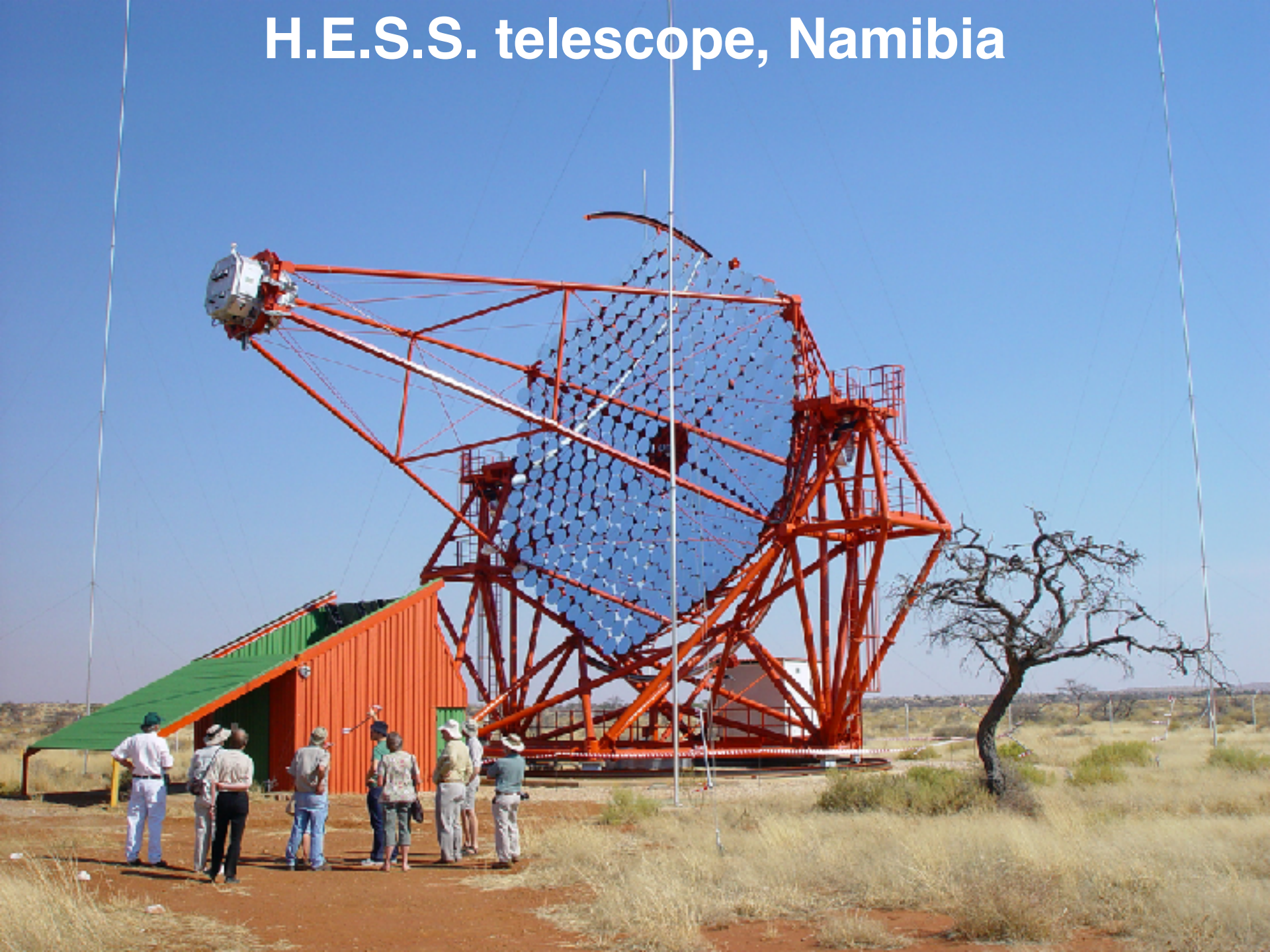


Figure 16.8 Illustration of the stereo-detection principle of imaging atmospheric Cherenkov telescopes [567]. The superimposed camera images are shown on the left-hand side. The intersection of the shower axes in this combined image corresponds to the arrival direction of the shower. From [567], © 2009 by Annual Reviews ([www.annualreviews.org](http://www.annualreviews.org)), reproduced with permission.

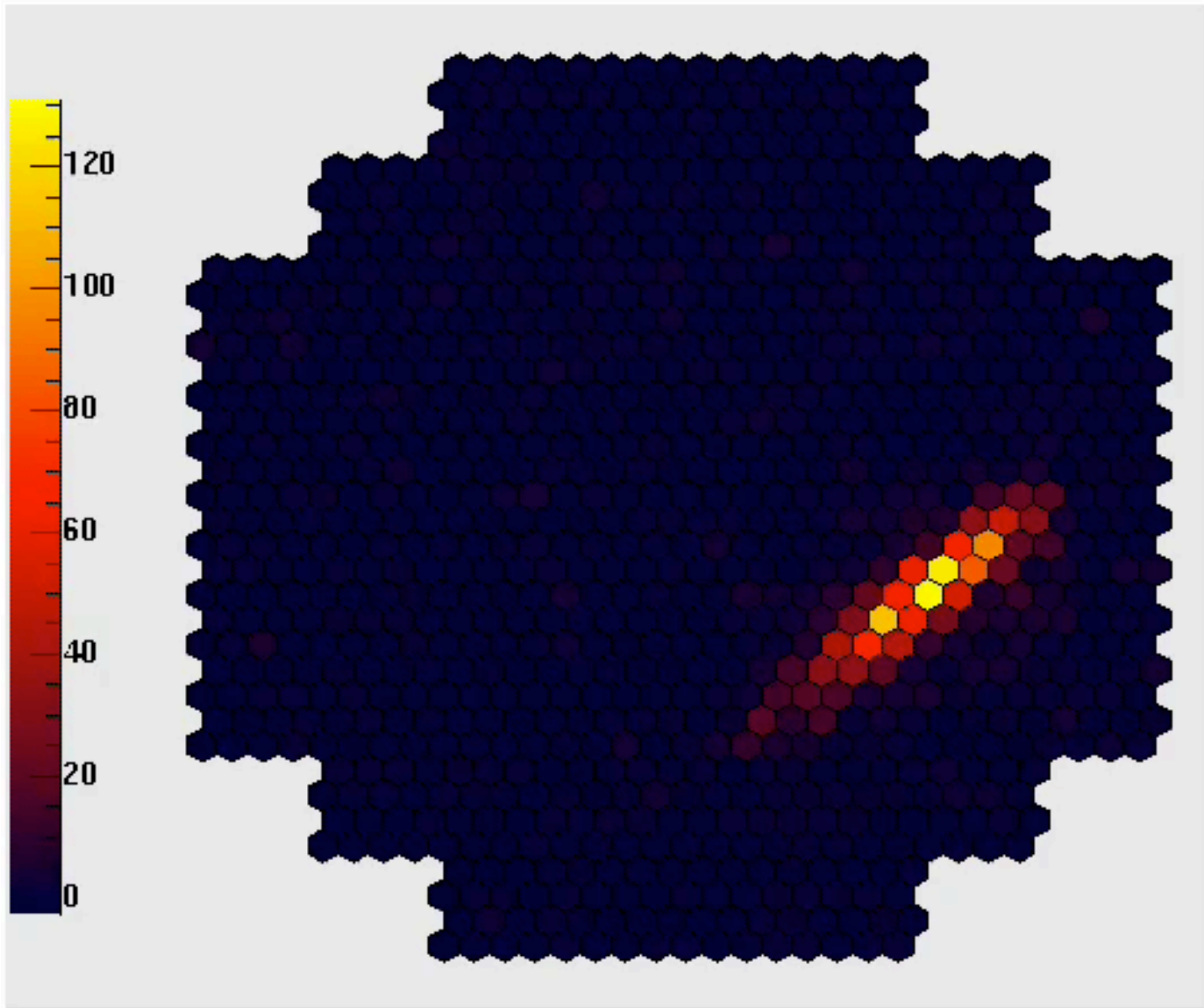
# H.E.S.S. telescope, Namibia



# H.E.S.S. telescope, Namibia



# H.E.S.S. telescope, Namibia



## - Fluorescence light

charged particles moving through the atmosphere excite nitrogen molecules

--> fluorescence light (300-450 nm)

isotropic radiation of fluorescence light



Cherenkov telescope

fluorescence telescope

--> air showers can be observed from aside

~0.5% of  $dE/dx$  ( $\sim 2.2 \text{ MeV}/(\text{g cm}^2)$ )

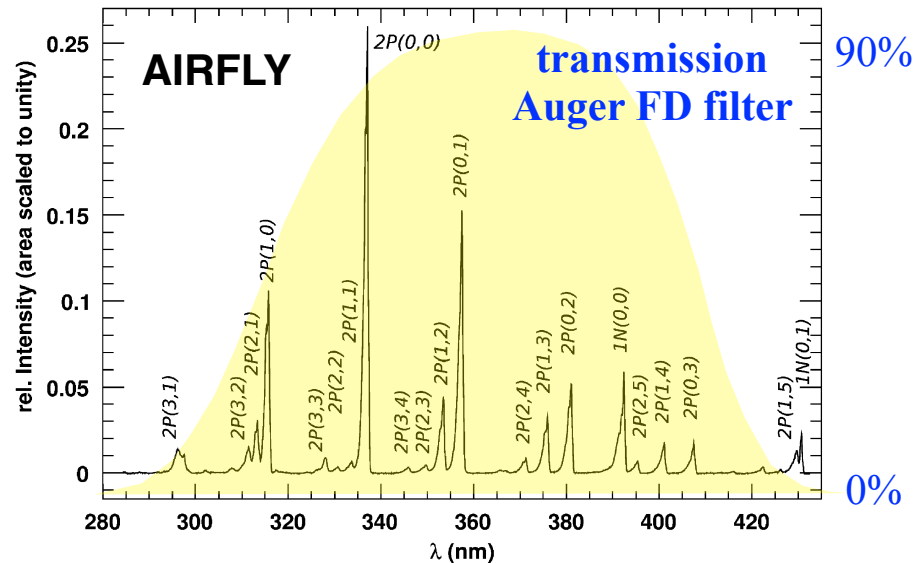


Fig. 2. Air-fluorescence spectrum excited by 3 MeV electrons at 800 hPa as measured by the AIRFLY Collaboration [32].



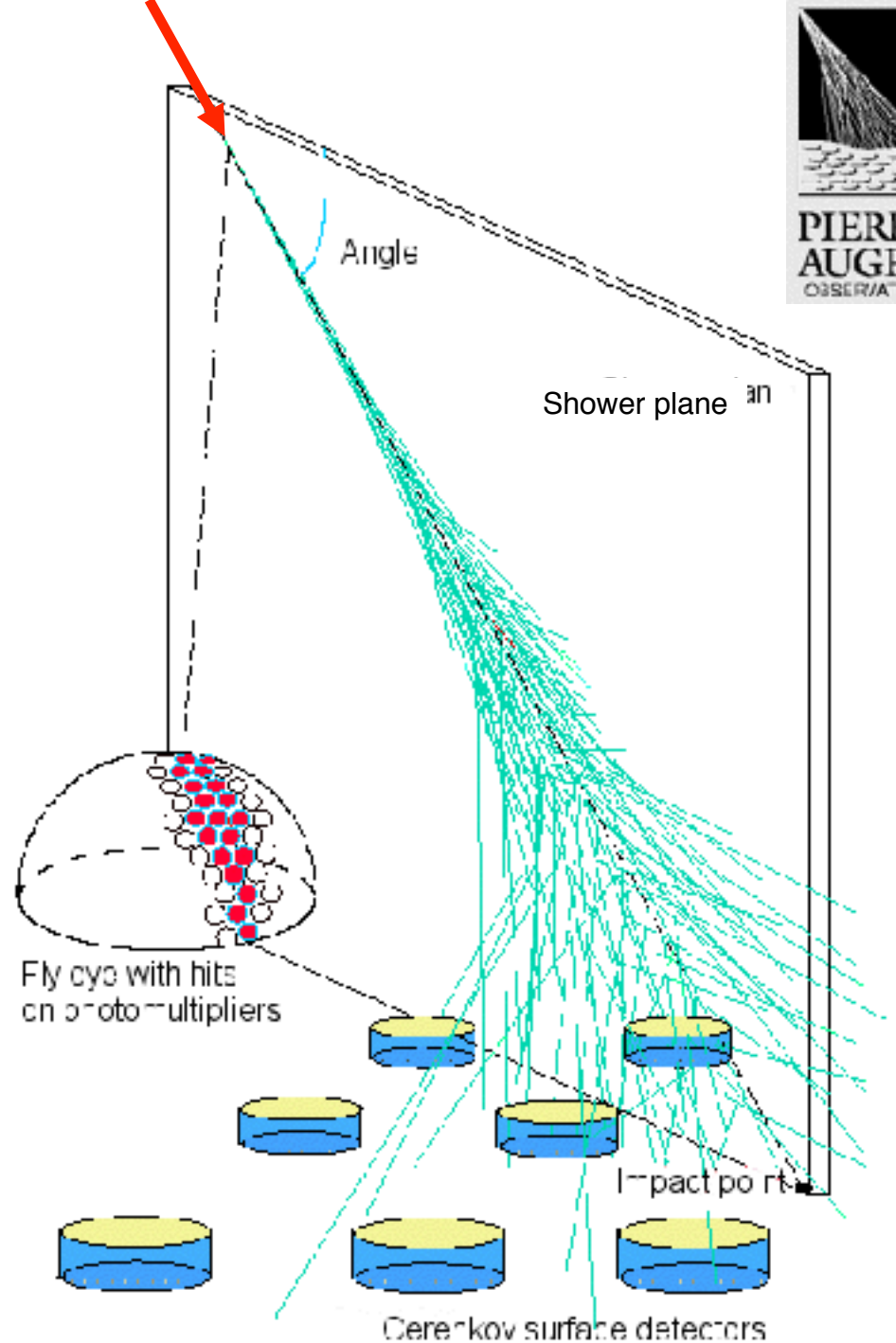


# Hybrid Detector

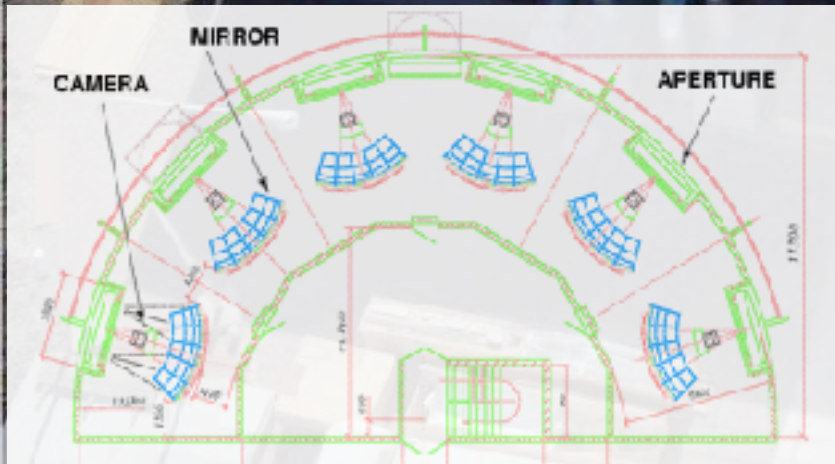


Fluorescence telescope

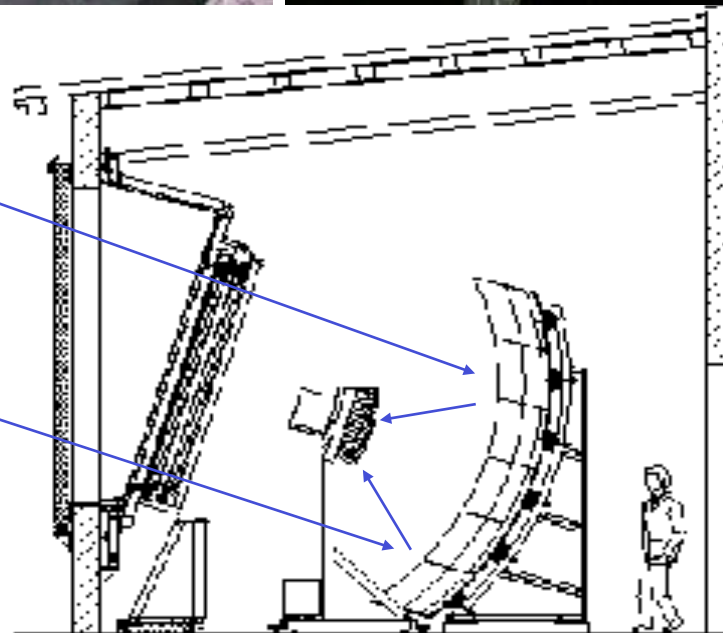
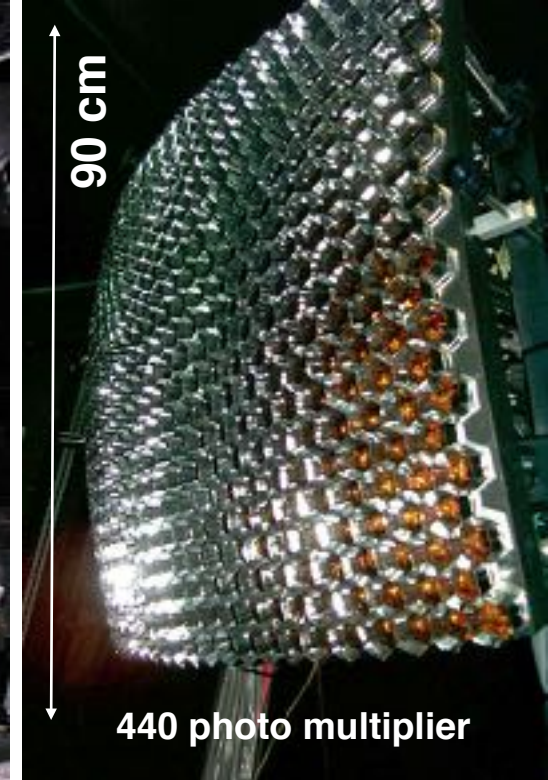
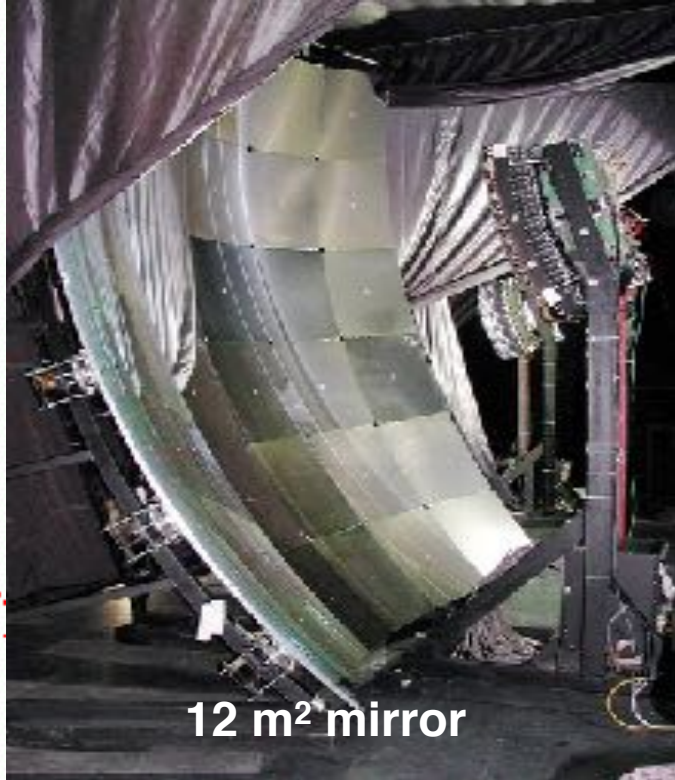
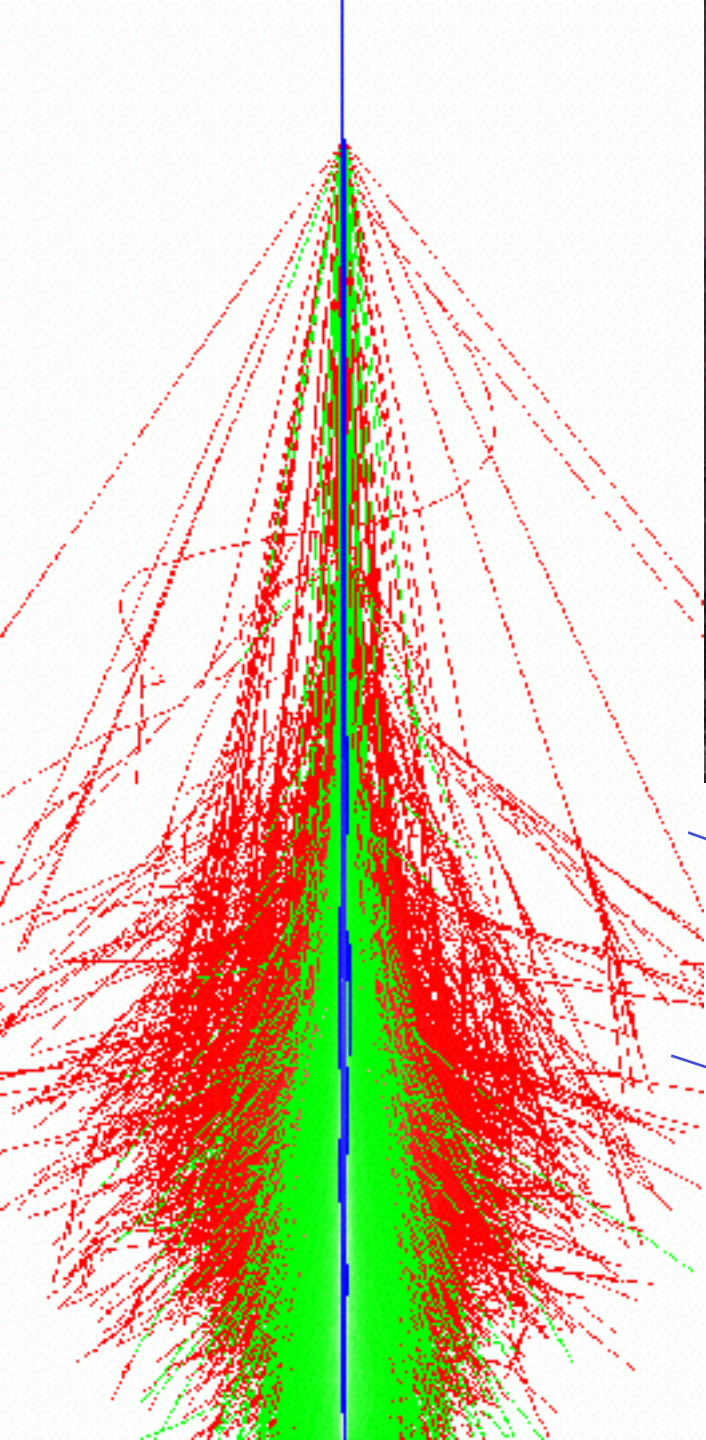
Surface array  
(water Cherenkov detectors)

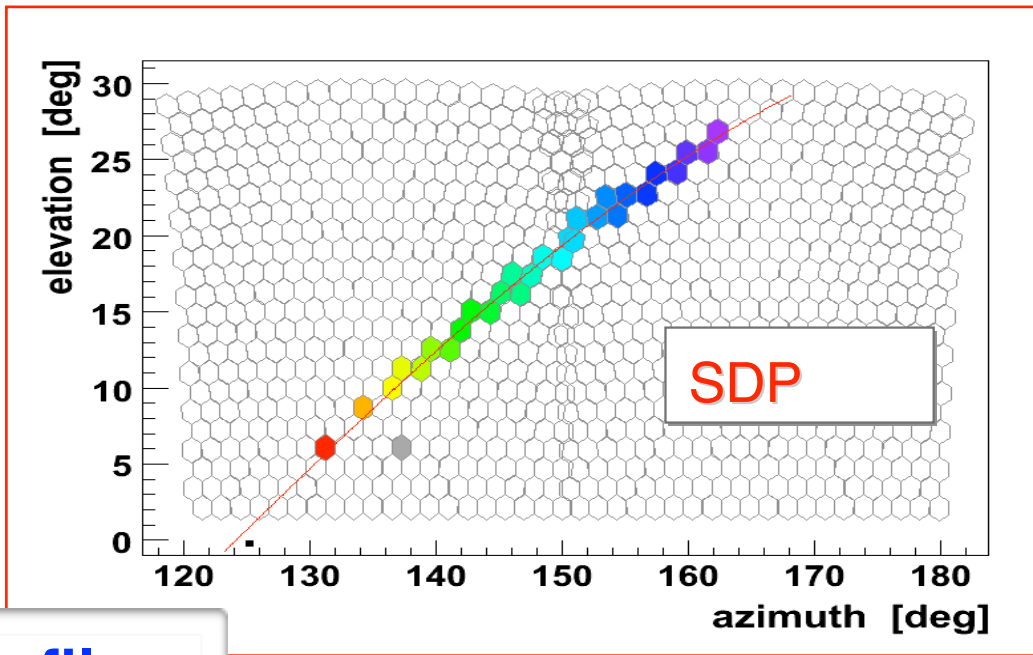
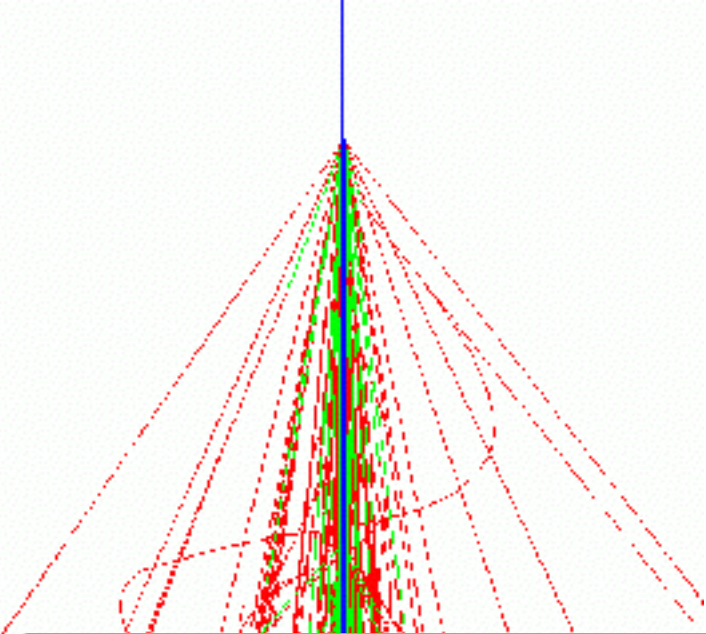


# FD eye view

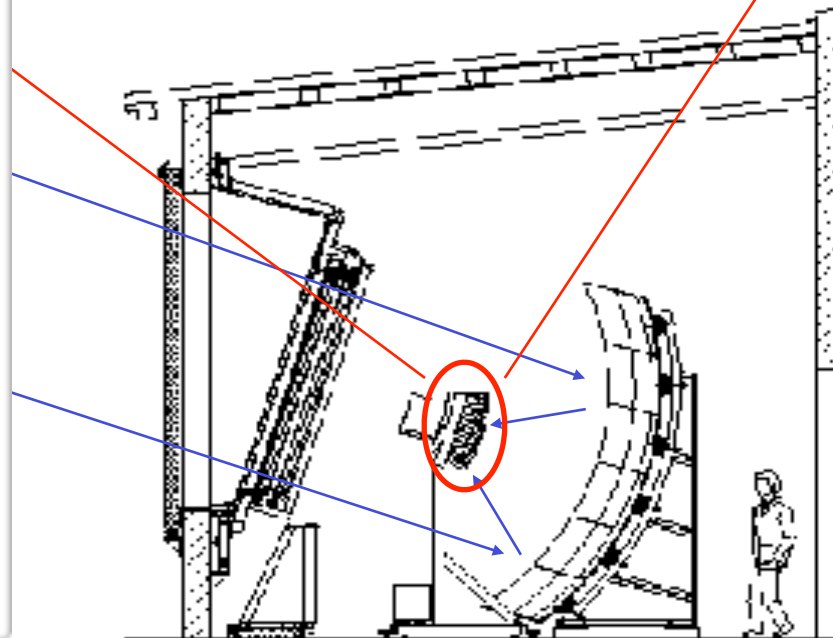
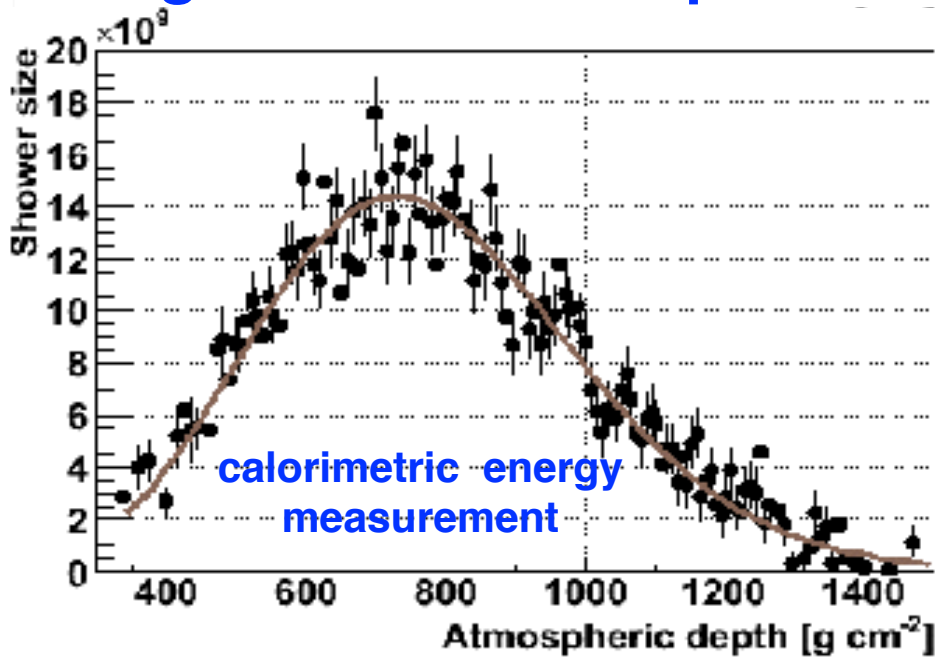


Four sites  
Six telescopes  
viewing  $30^\circ \times 30^\circ$  each





## longitudinal shower profile



The function proposed by Gaisser and Hillas [598] gives a good phenomenological description of individual as well as averaged longitudinal shower profiles

$$N(X) = N_{\max} \left( \frac{X - X_1}{X_{\max} - X_1} \right)^{(X_{\max} - X_1)/\Lambda} \exp \left( -\frac{X - X_{\max}}{\Lambda} \right). \quad (16.38)$$

It is often used to extrapolate the measured shower profiles to depth ranges outside the field of view of the telescopes and to fit for  $N_{\max}$  and  $X_{\max}$ . In doing so,  $X_1$  and  $\Lambda = 55 - 65 \text{ g/cm}^2$  are parameters of the fit. In particular,  $X_1$  can be negative. The same function with  $X_1 > 0$  interpreted as the point of first interaction is sometimes used as a toy model to illustrate fluctuations in air shower, as discussed in Appendix [A.8](#).

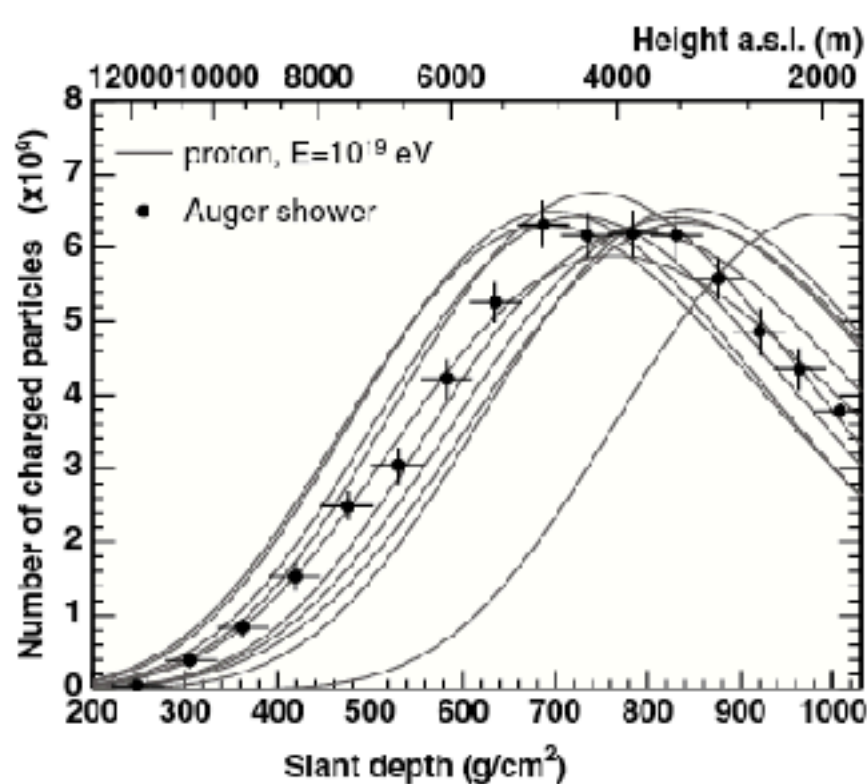
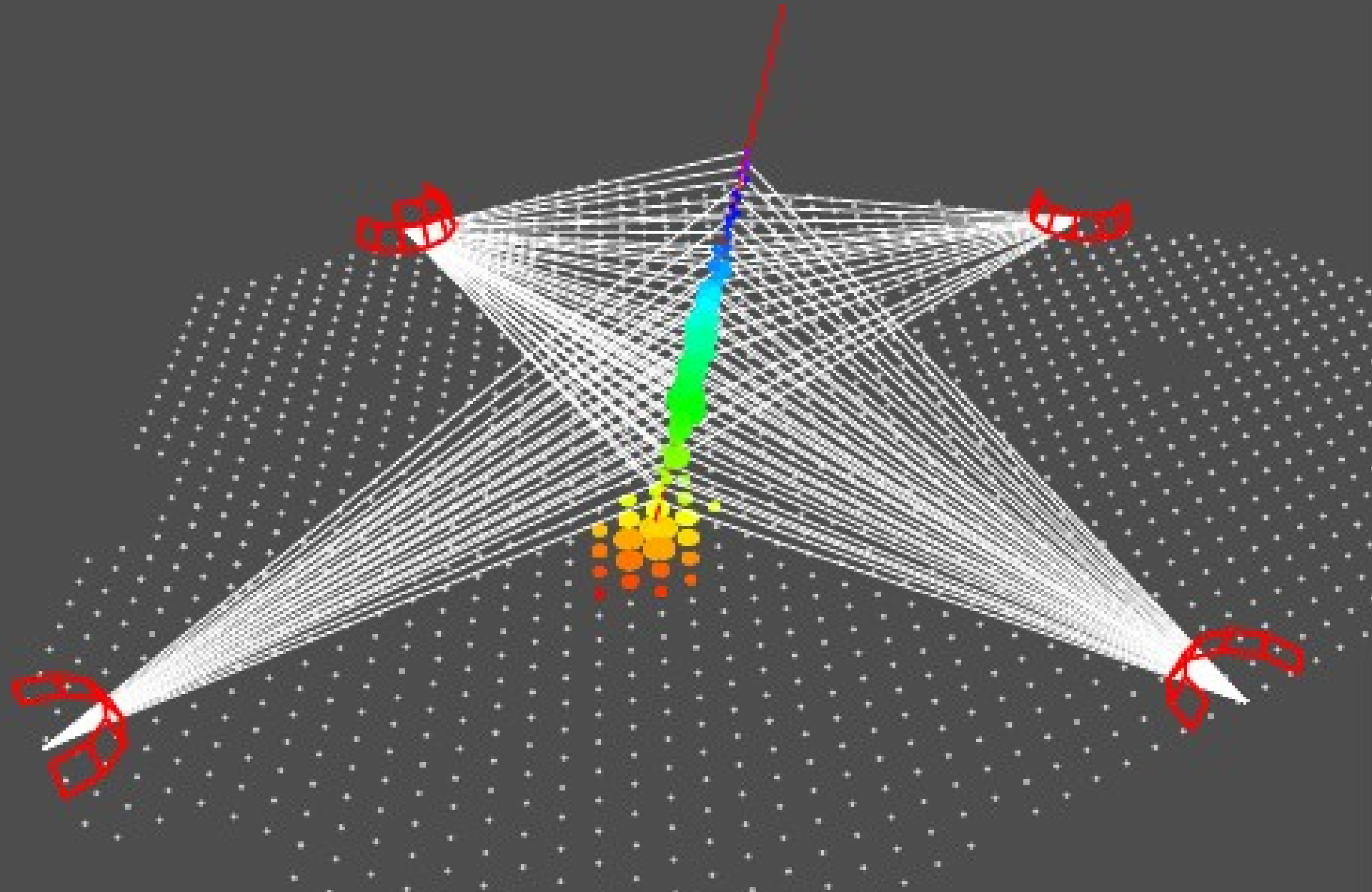


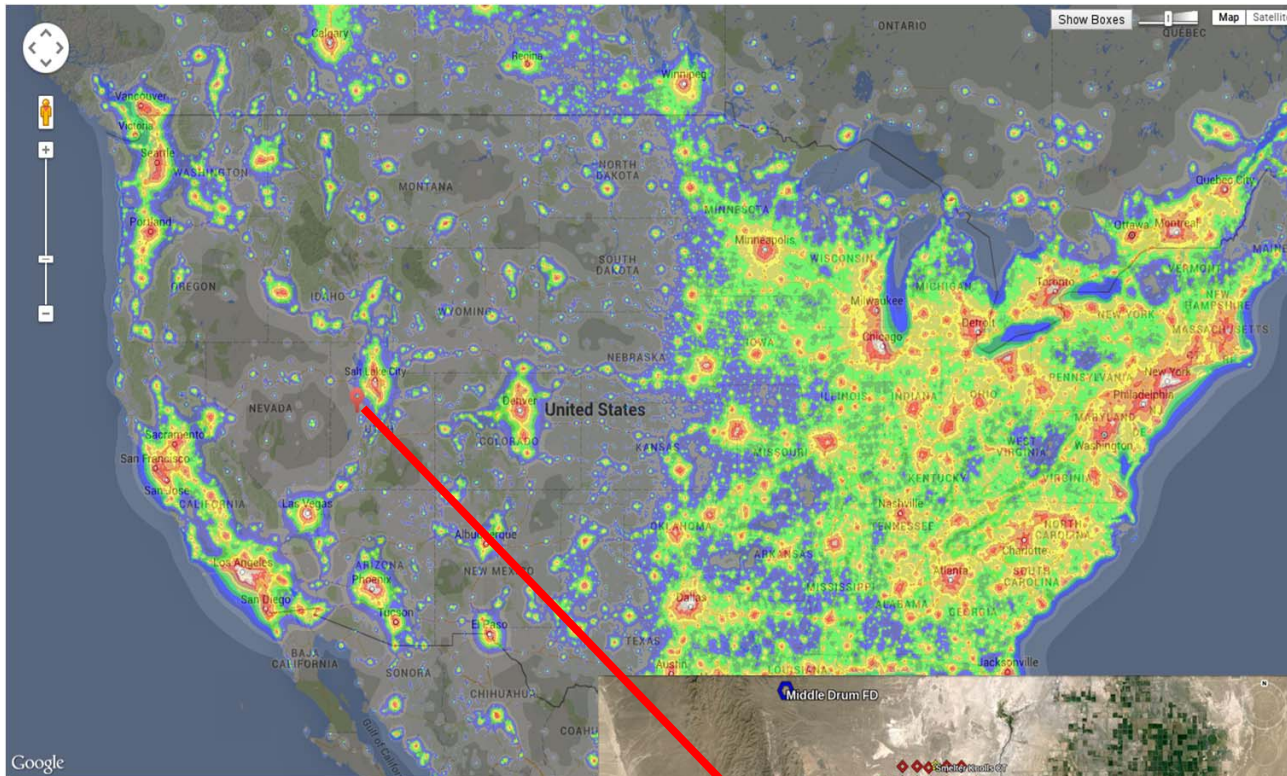
Figure 16.9 Profile of one shower measured with the Pierre Auger Observatory [594]. The reconstructed energy of this shower is about  $10^{19}$  eV. The data are shown together with 10 simulated proton (left) and 10 iron showers (right) to demonstrate the composition sensitivity of the depth of shower maximum. The showers were simulated with the Sibyll interaction model [158, 520] and the CONEX air shower package [514].

# A Hybrid Event

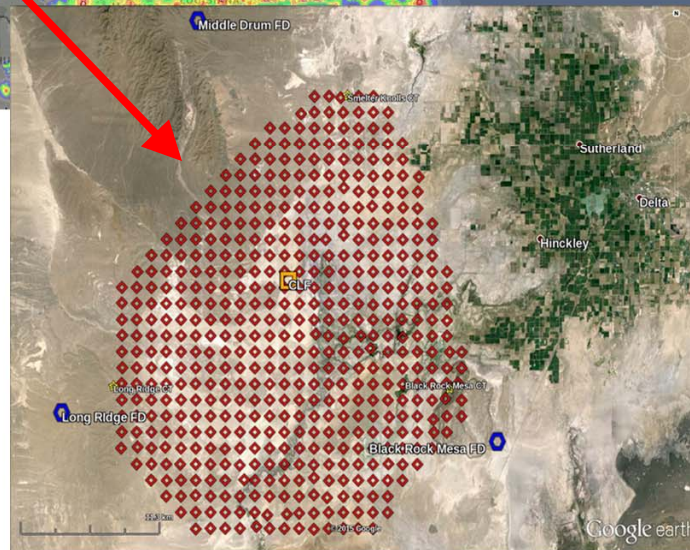




# Telescope Array Observatory



U.S. Light Pollution Map



Largest cosmic ray observatory in the Northern hemisphere.

$\sim 700 \text{ km}^2 \rightarrow \lesssim$  land area of New York City.

Millard County, Utah

$39.30^\circ \text{ N}$

$112.91^\circ \text{ W}$

1550 m ASL

$\sim 800 \text{ g/cm}^2$  vertical depth

The High Energy component of Telescope Array – 38 fluorescence telescopes (9728 PMTs) at 3 telescope stations overlooking an array of 507 scintillator surface detectors (SD) - operational as of 2008.

# TA Fluorescence Detectors

Middle Drum

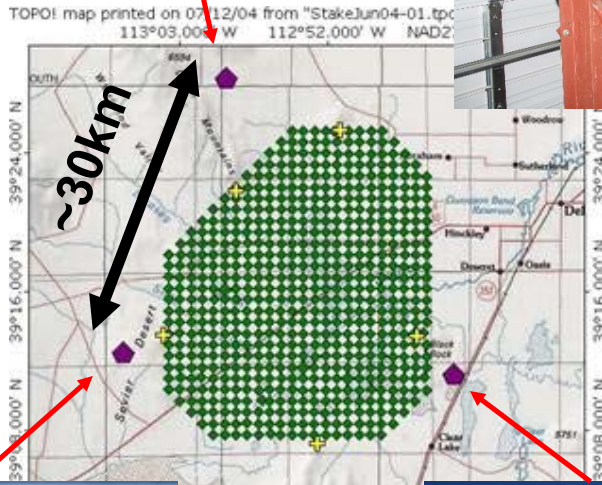


14 telescopes @ station  
256 PMTs/camera



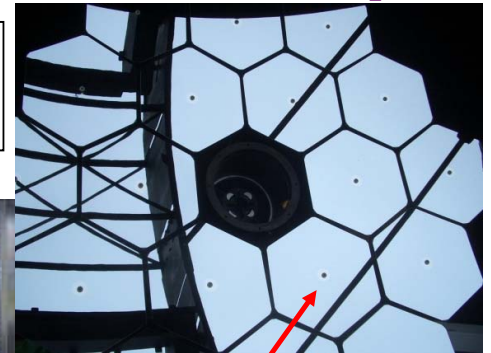
5.2 m<sup>2</sup>

Reutilized from HiRes-I



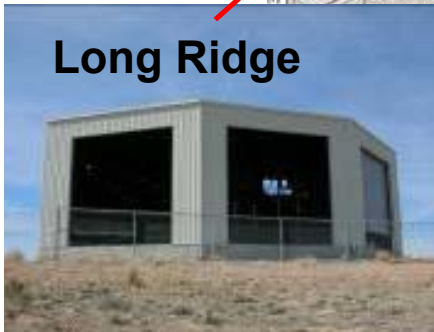
12 telescopes/station  
256 PMTs/camera

New Telescopes



6.8 m<sup>2</sup>

Long Ridge



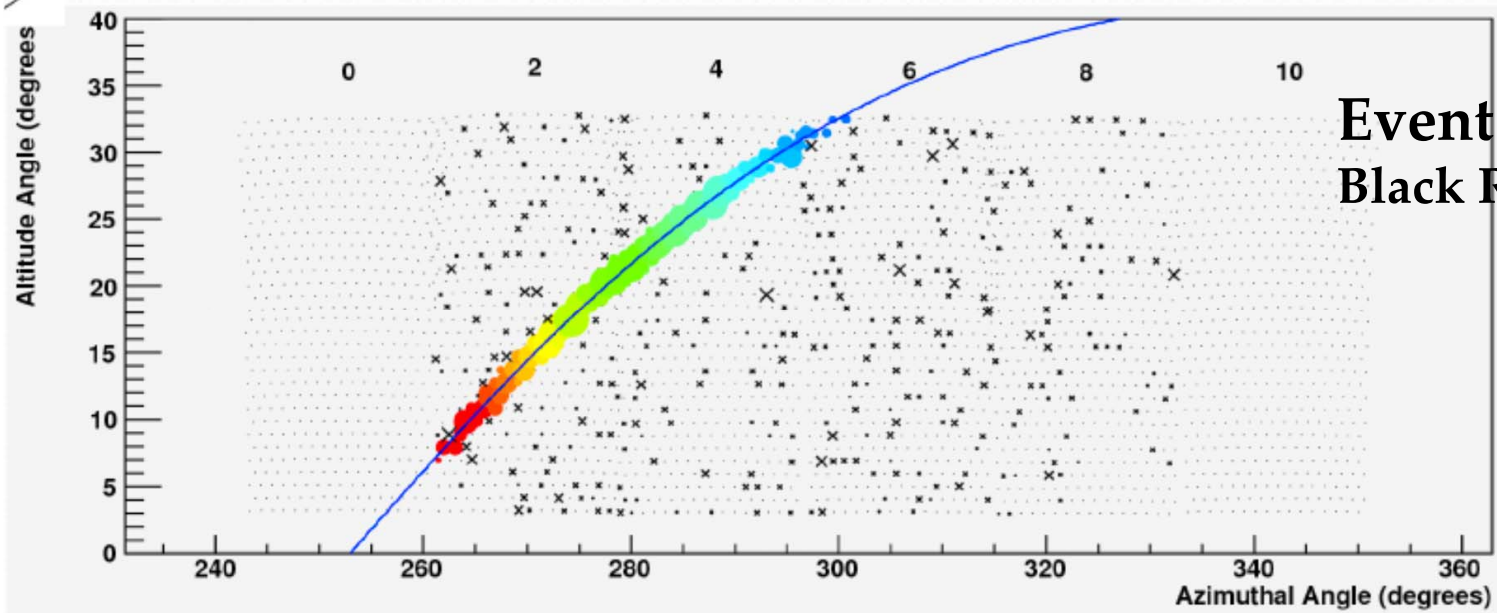
Black Rock Mesa



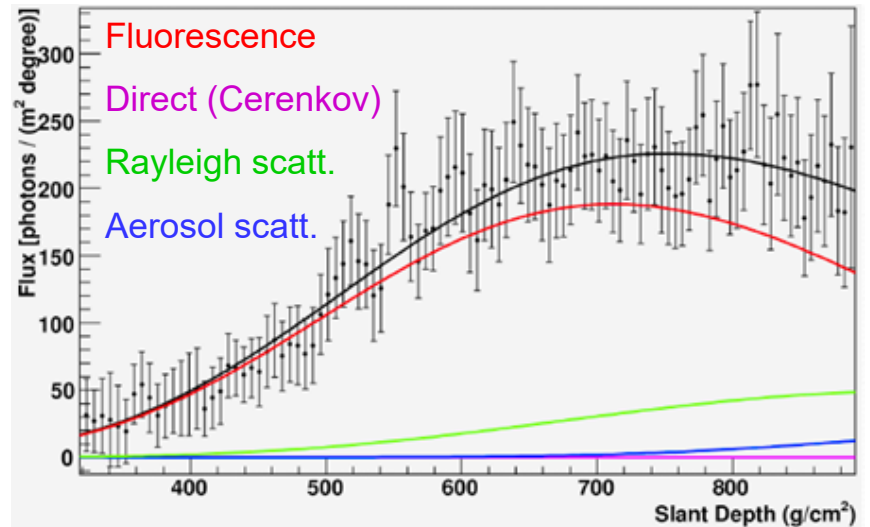
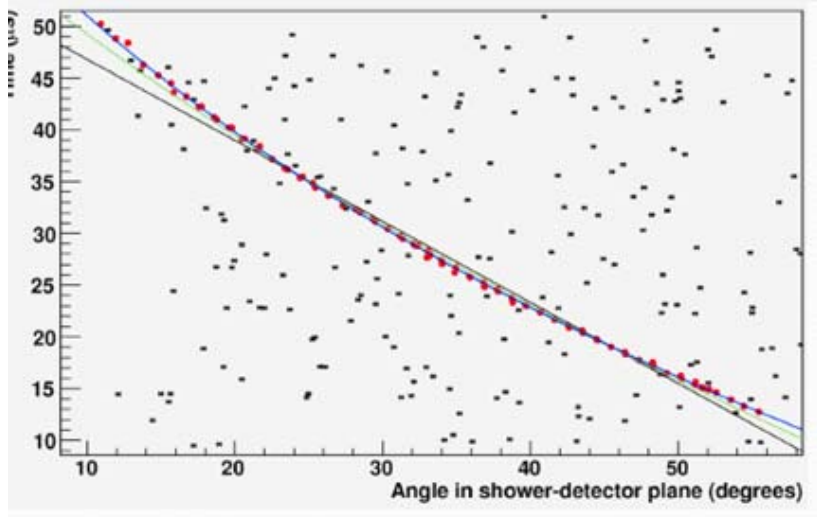
~1 m<sup>2</sup>

Busan, S.Korea

# Typical Fluorescence Event



Event Display  
Black Rock Mesa



Monocular timing fit (time vs angle)

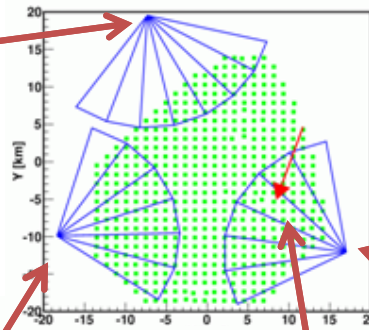
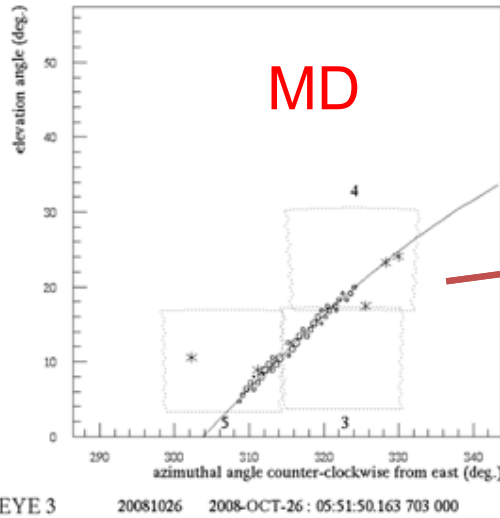
Reconstructed Shower Profile

## Scintillator Detectors on a 1.2 km square grid



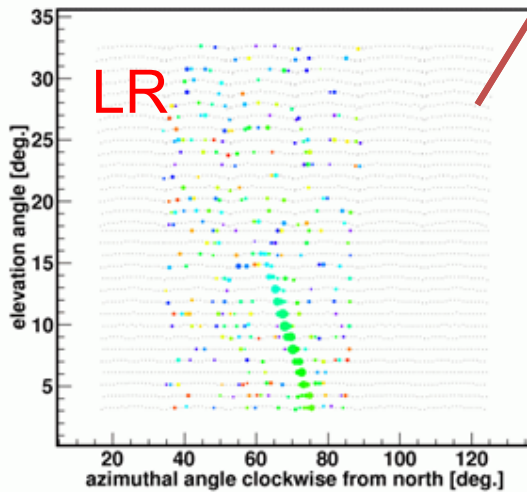
- Power: Solar/Battery
- Readout: Radio
- Self-calibrated:  
 $\mu$  background
- Operational: 3/2008

# Example Event

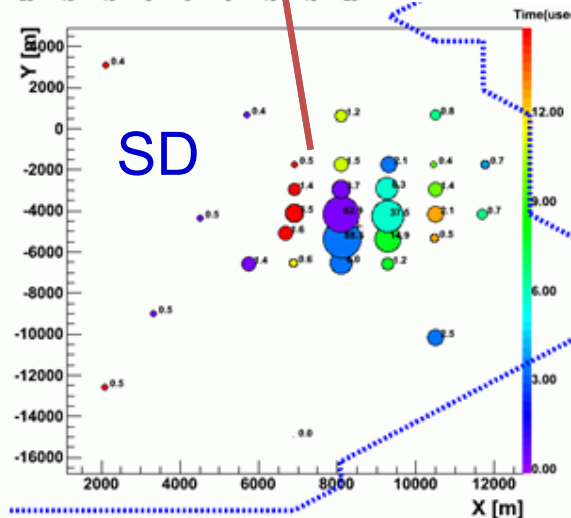


	$\theta$ [°]	$\phi$ [°]	x[km]	y[km]
MD mono	51.43	73.76	7.83	-3.10
BR mono	51.50	77.09	7.67	-4.14
Stereo BR&LR	50.21	71.30	8.55	-4.88

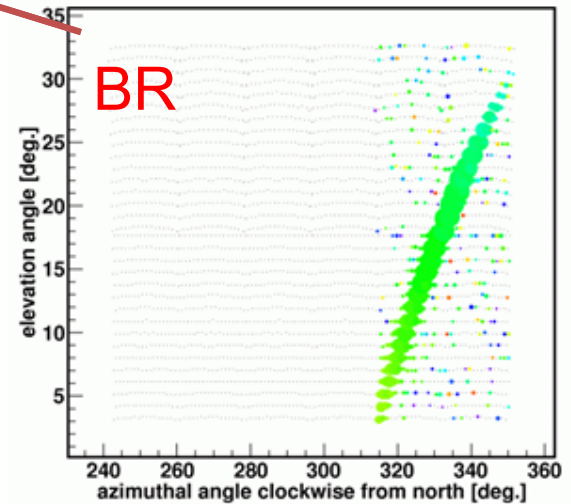
Event from 2008-10-26



17 July 2017



J.N. Matthews



35<sup>th</sup> ICRC, Busan, S.Korea