

Astroparticle Physics

2020/21

Lectures:

- [1. Historical introduction, basic properties of cosmic rays](#)
- [2. Hadronic interactions and accelerator data](#)
- [3. Cascade equations](#)
- [4. Electromagnetic cascades](#)
- [5. Extensive air showers](#)
6. Detectors for extensive air showers
7. High energy cosmic rays and the knee in the energy spectrum of cosmic rays
8. Radio detection of extensive air showers
9. Acceleration, astrophysical accelerators and beam dumps
10. Extragalactic propagation of cosmic rays
11. Ultra high energy cosmic rays
12. Astrophysical gamma rays and neutrinos
13. Neutrino astronomy
14. Gamma-ray astronomy

<http://particle.astro.ru.nl/goto.html?astropart2021>

lecture 7

High energy cosmic rays and the knee in the energy spectrum

Gaisser chapter 17

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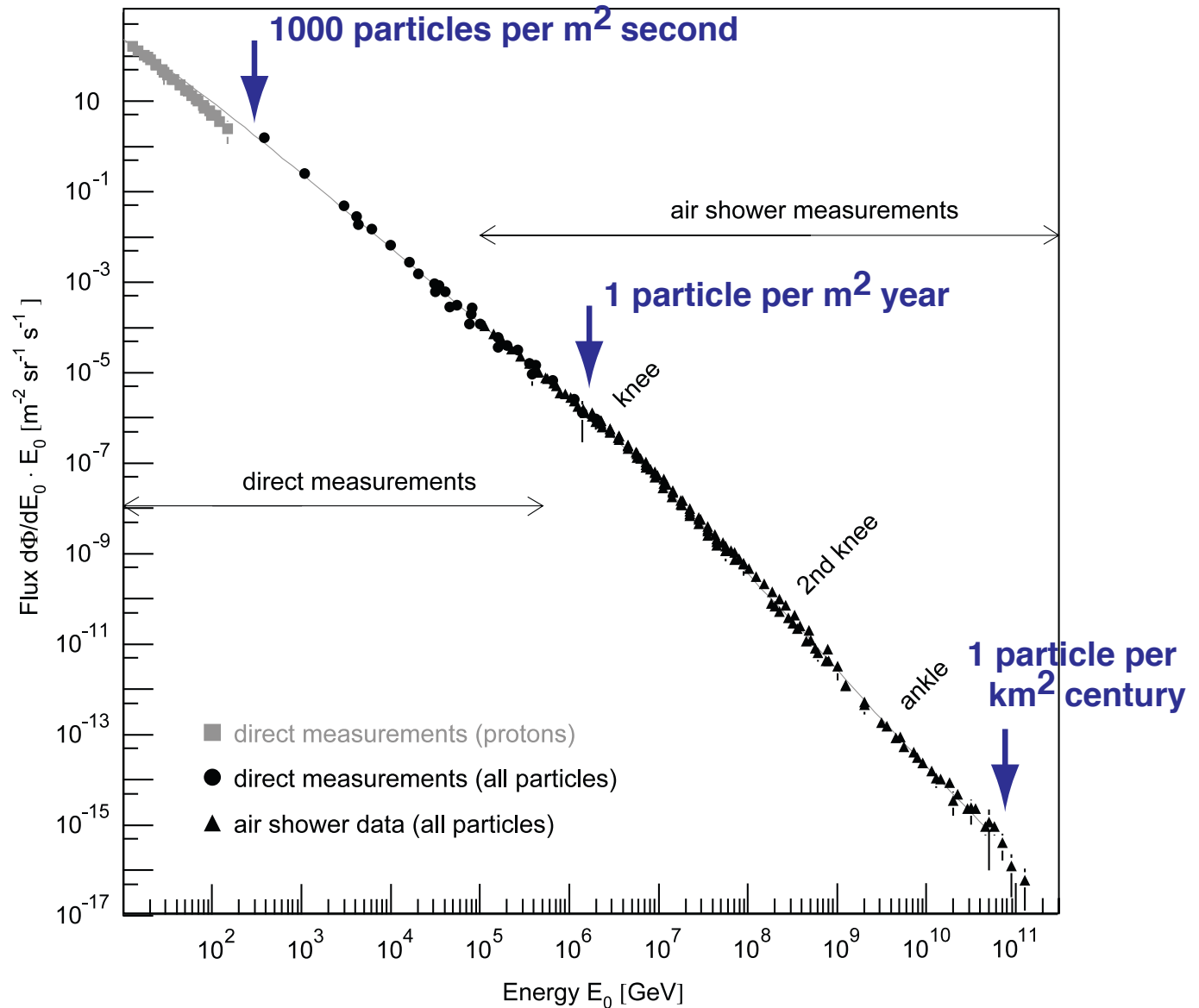


Fig. 1. All-particle energy spectrum of cosmic rays as measured directly with detectors above the atmosphere and with air shower detectors. At low energies, the flux of primary protons is shown.

Energy spectrum of cosmic rays

energy density

$$\rho_E = \frac{4\pi}{c} \int \frac{E}{\beta} \frac{dN}{dE} dE \approx 1 \frac{\text{eV}}{\text{cm}^3}$$

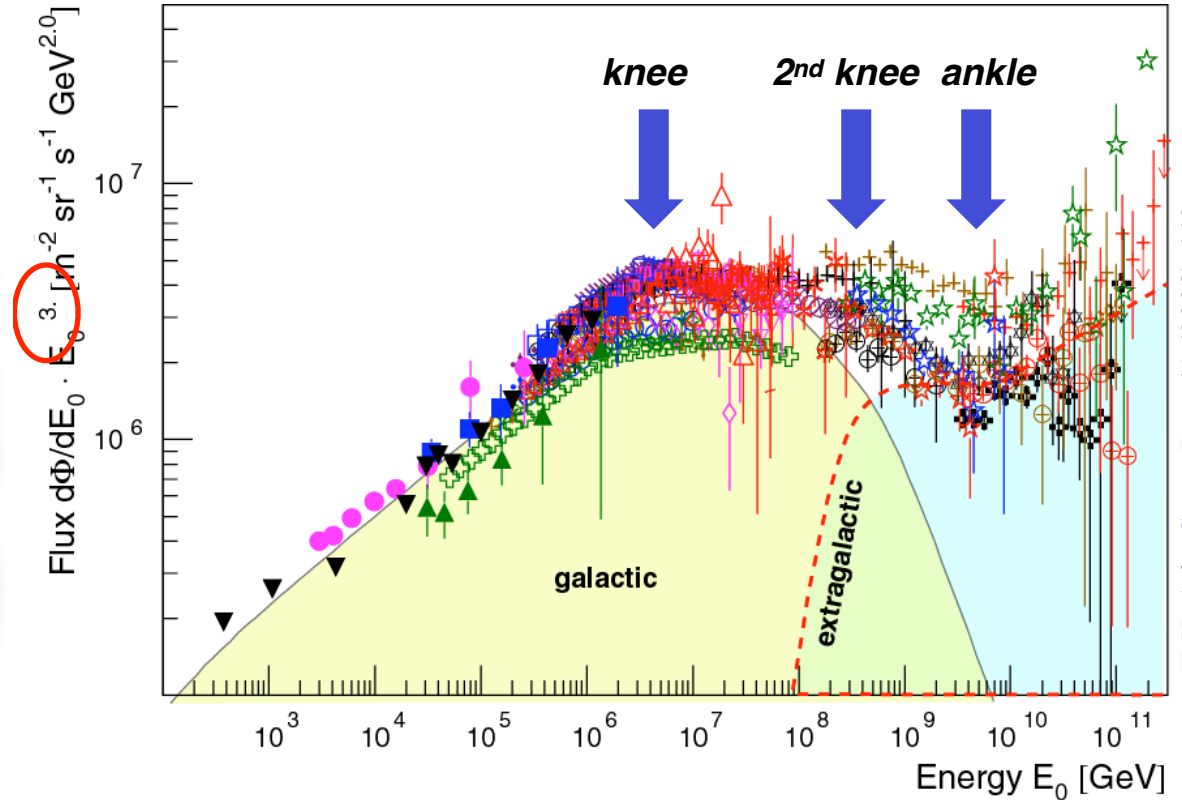
for comparison:

$$\rho_B = \frac{B^2}{2\mu_0} \approx 0.25 \text{ eV/cm}^3$$

$$\rho_{SL} \approx 0.3 \text{ eV/cm}^3$$

$$\rho_{IR} \approx 0.4 \text{ eV/cm}^3$$

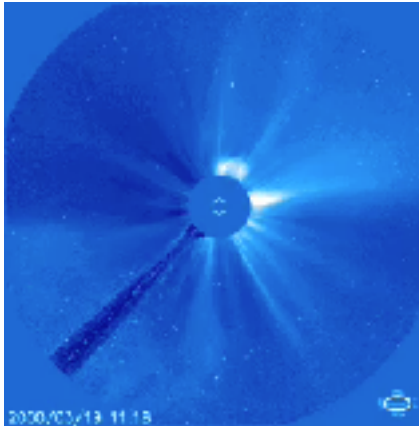
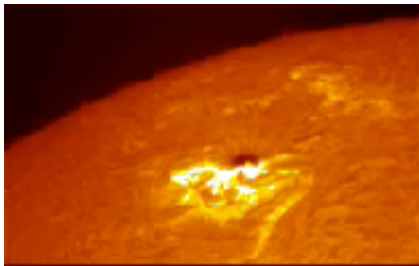
$$\rho_{3K} \approx 0.25 \text{ eV/cm}^3$$



JRH, Adv. Space Res. 41 (2008) 442

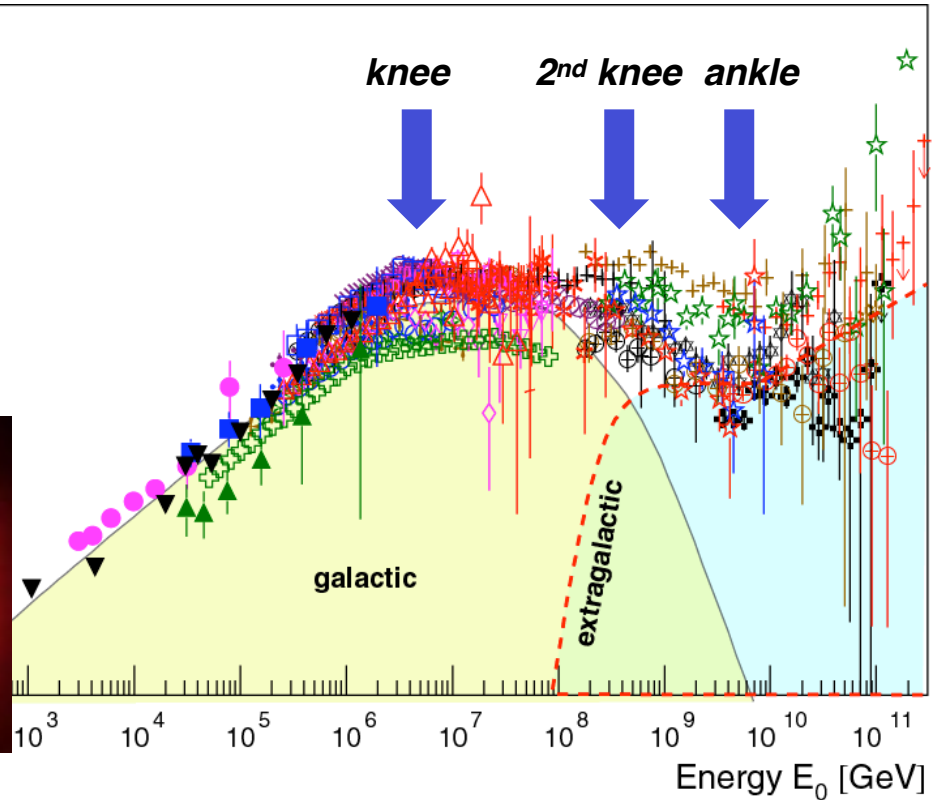
Energy spectrum of cosmic rays

Solar flares



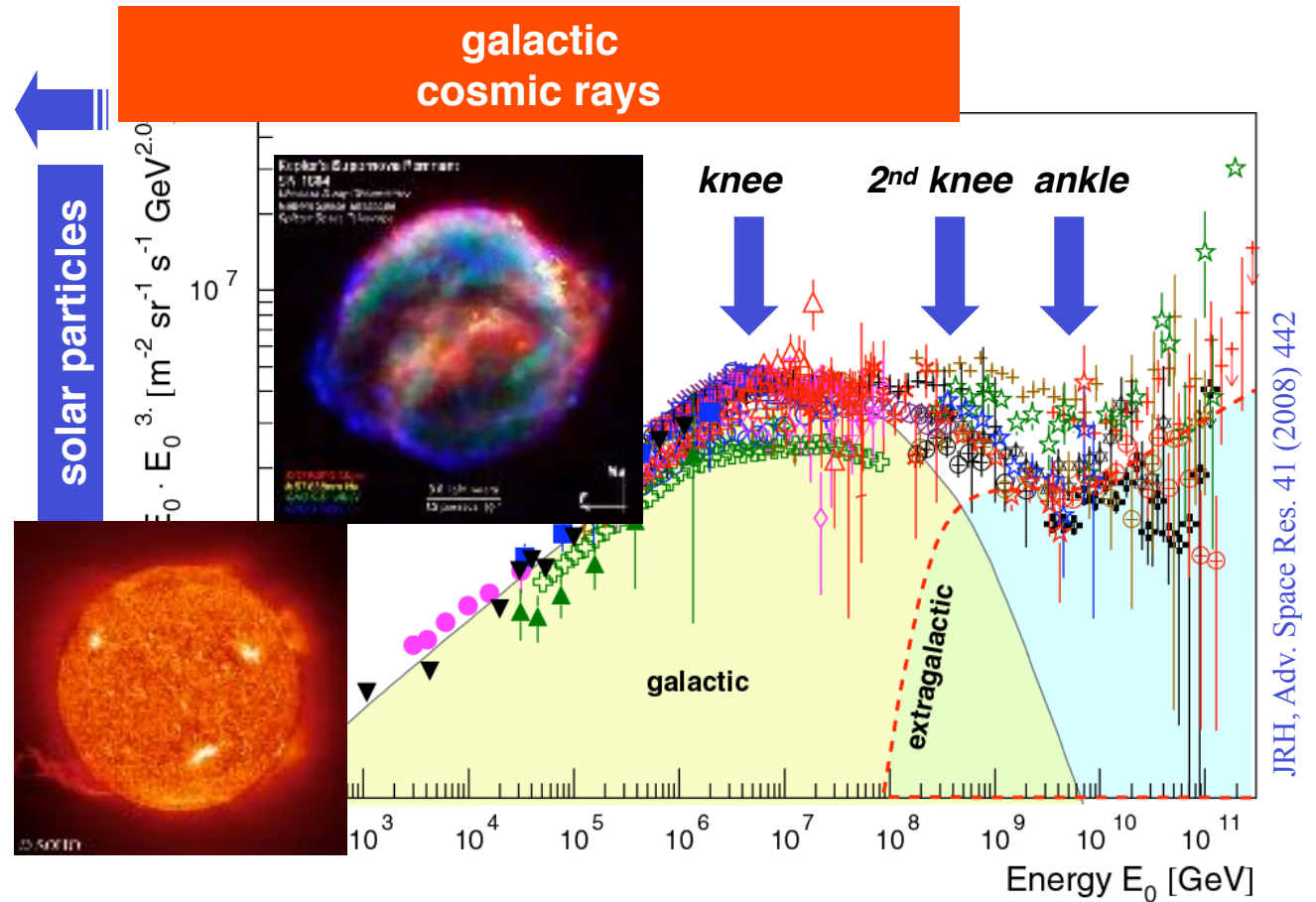
solar particles

$E_0 \cdot E_0^3 \cdot [\text{m}^{-2} \text{sr}^{-1} \text{s}^{-1} \text{GeV}^{2.0}]$



JRH, Adv. Space Res. 41 (2008) 442

Energy spectrum of cosmic rays



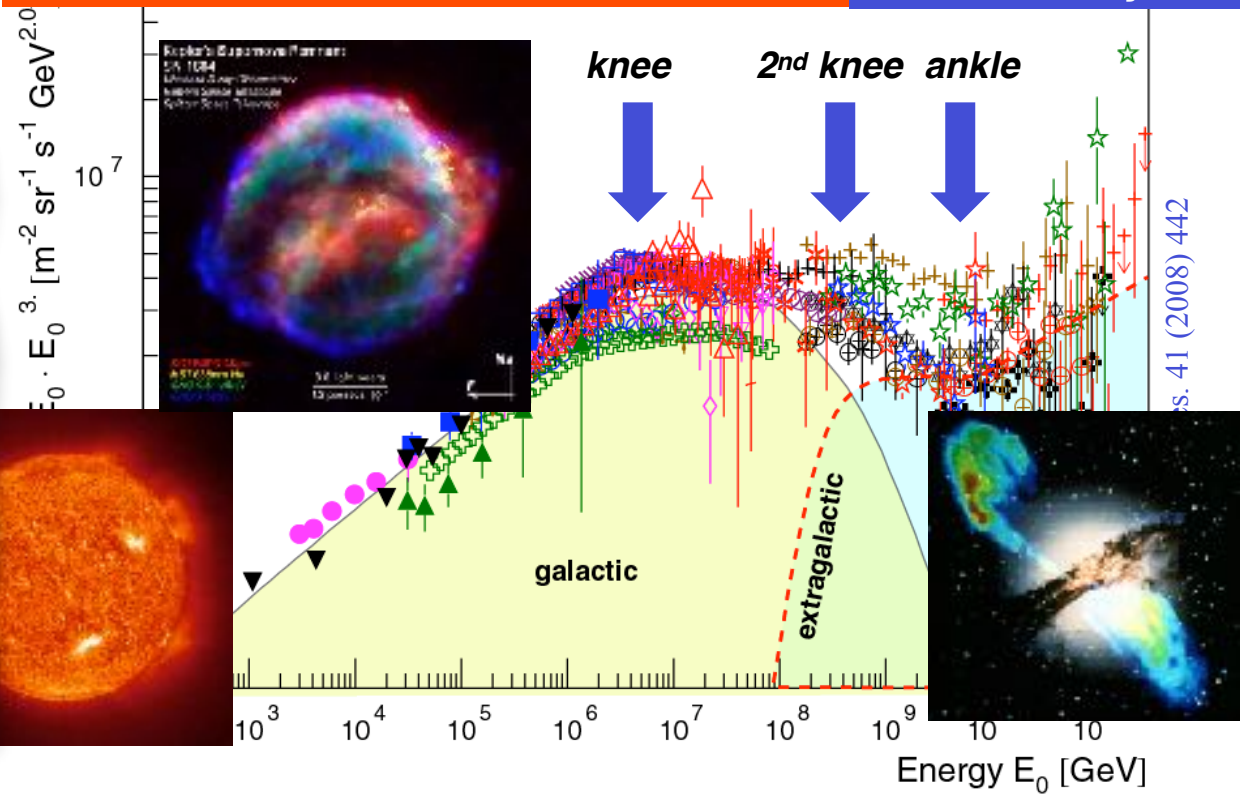
JRH, Adv. Space Res. 41 (2008) 442

Energy spectrum of cosmic rays

galactic cosmic rays extragalactic cosmic rays

Radius of particle in magnetic field

$$r = \frac{p}{ZeB} \quad r[\text{pc}] = 1.08 \cdot \frac{E [\text{PeV}]}{B [\mu\text{G}]}$$



$r =$ 0.04 pc 3.6 pc 360 pc 36 kpc

s. 41 (2008) 442

Extensive air showers – Mass

Simple Heitler model of (hadronic) showers

Primary mass:

- Average depth of shower maximum X_{\max}

$$X_{\max}^A \sim \ln \frac{E_0}{A}$$

$$X_{\max}^A = X_{\max}^p - X_0 \ln A$$

$$X_{\max}^{\text{Fe}} = X_{\max}^p - 150 \text{ g/cm}^2$$

- N_e - N_μ ratio

$$\frac{N_e}{N_\mu} \approx 35.1 \left(\frac{E_0}{A \text{ PeV}} \right)^{0.15} \quad \text{or}$$

$$\lg \left(\frac{N_e}{N_\mu} \right) = C - 0.065 \ln A$$

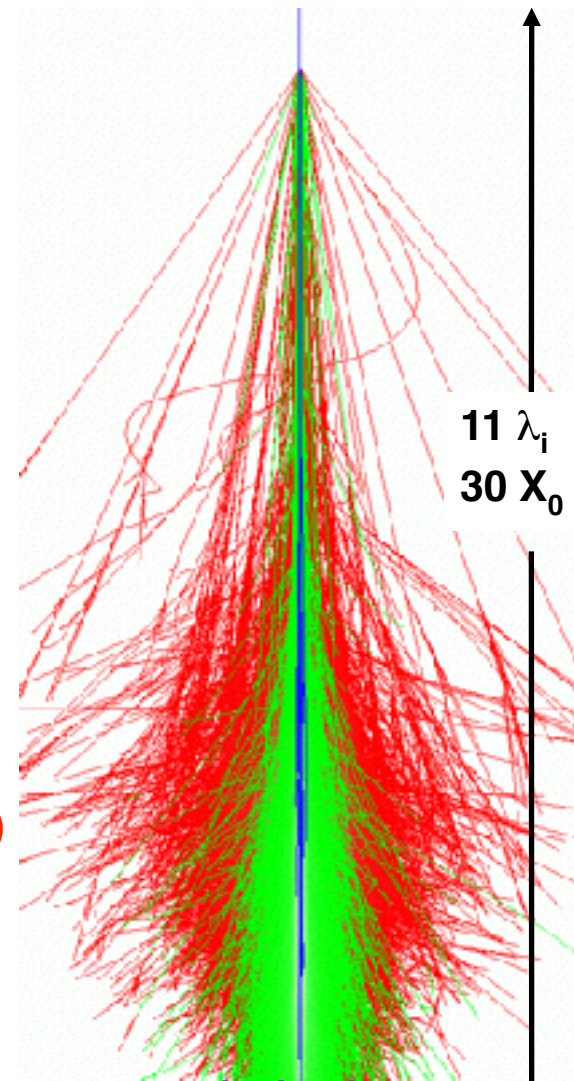
$$\Delta \ln A \sim 1$$

$$\rightarrow \Delta X_{\max} \sim 36 \text{ g/cm}^2$$

$$\rightarrow \Delta (N_e/N_\mu) \sim 16\%$$

$$\Delta \ln A \approx 0.8$$

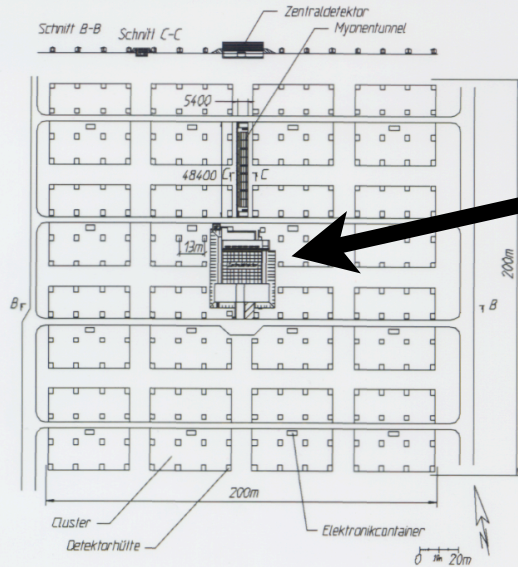
in „best“ experiments



electromagnetic muonic
shower components

The KASCADE Array

200 m



□ Array

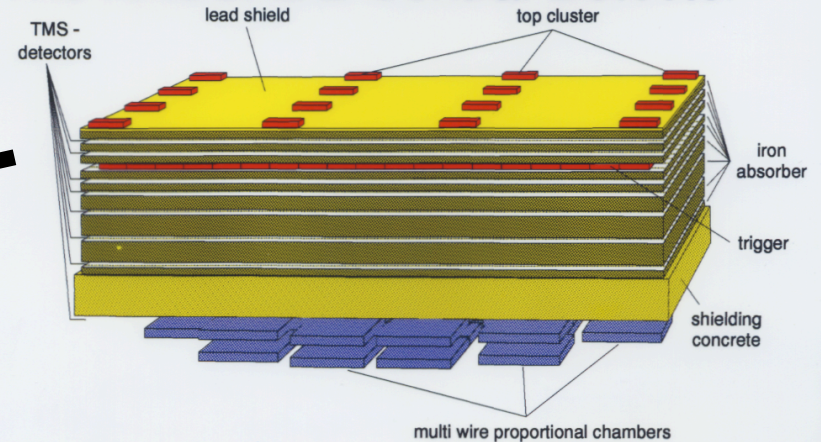
200 m

252 stations with

- e/γ detectors: 4 cm liquid scintillator ($\approx 3.1 \text{ m}^2$)
- muon detectors: 3 cm plastic scintillator ($\approx 3.2 \text{ m}^2$)
- shield 10 cm Pb + 4 cm Fe $\rightarrow 20 X_0$
- ⇒ $E_\mu > 300 \text{ MeV}$
- measurement of particle densities and arrival times

JRH 06/96

The KASCADE Central Detector



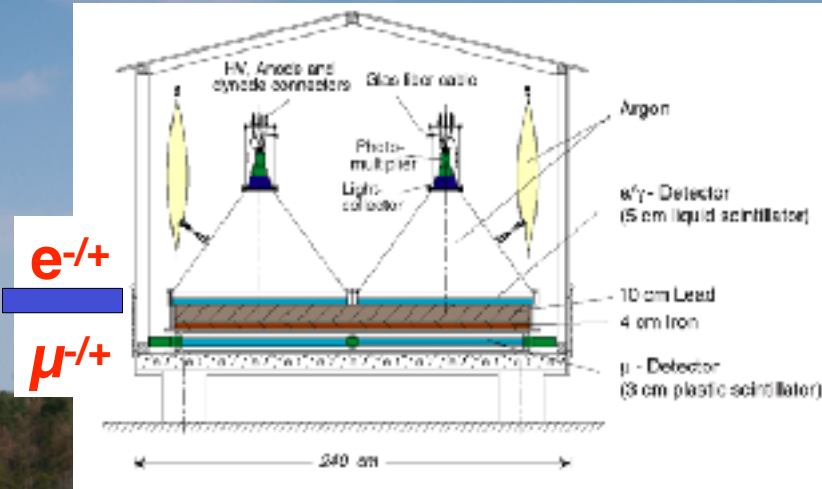
4 different detector systems:

- Top array $32 \times 0.45 \text{ m}^2$ scintillation counters (14 m^2)
- Trigger layer $456 \times 0.45 \text{ m}^2$ scintillation counters (206 m^2)
- 32 Multiwire proportional chambers ($2 \times 150 \text{ m}^2$) $E_\mu > 2 \text{ GeV}$
- Iron sampling calorimeter ($16 \times 20 \text{ m}^2$)
- 8 active layers $\rightarrow 10\,000$ liquid ionisation chambers
- 40 000 electronic channels
- separation of individual hadrons by a fine segmentation
- depth $\approx 11 \lambda_I \rightarrow$ hadrons up to 10 TeV 95% contained

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KARlsruhe Shower Core and Array DETector

Simultaneous measurement of electromagnetic, muonic, hadronic shower components

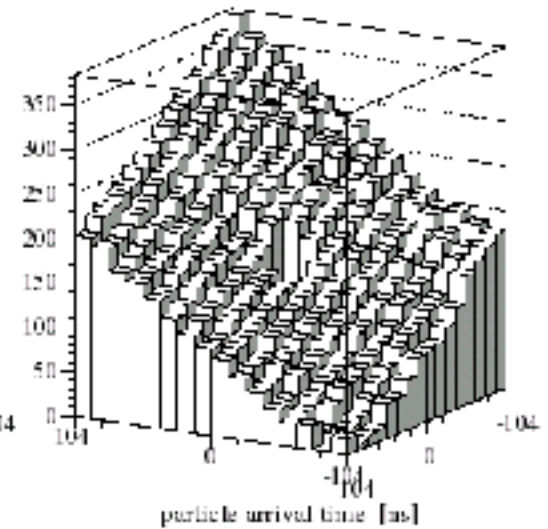
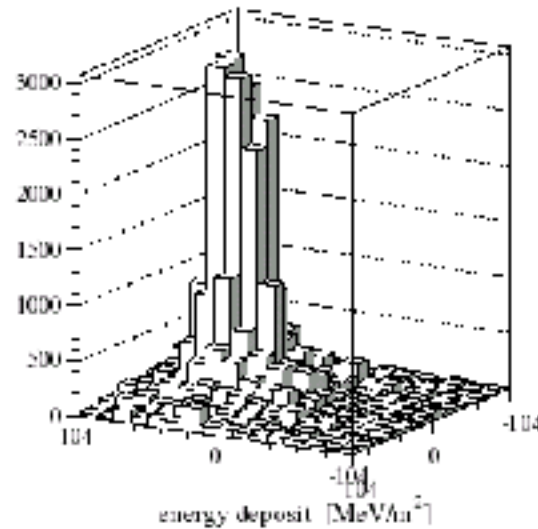
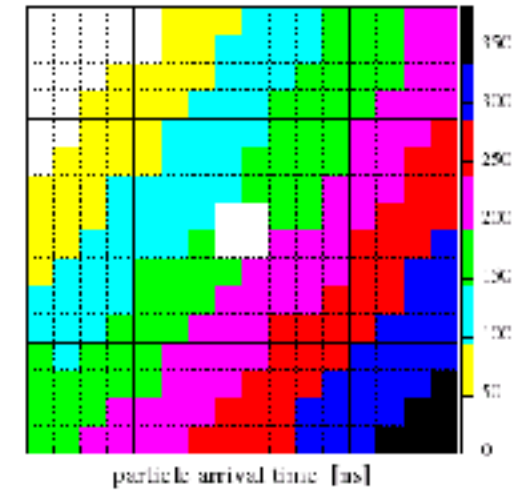
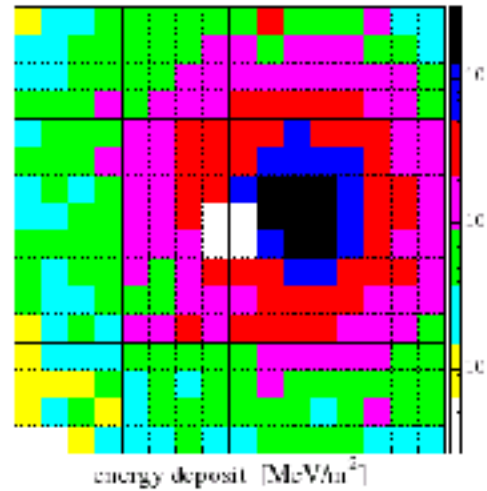


Event reconstruction in the scintillator array

electromagnetic component

shower core	$\Delta r = 2.5 - 5.5 \text{ m}$
shower direction	$\Delta\alpha = 0.5^\circ - 1.2^\circ$
shower size	$\Delta N_e/N_e = 6 - 12 \%$

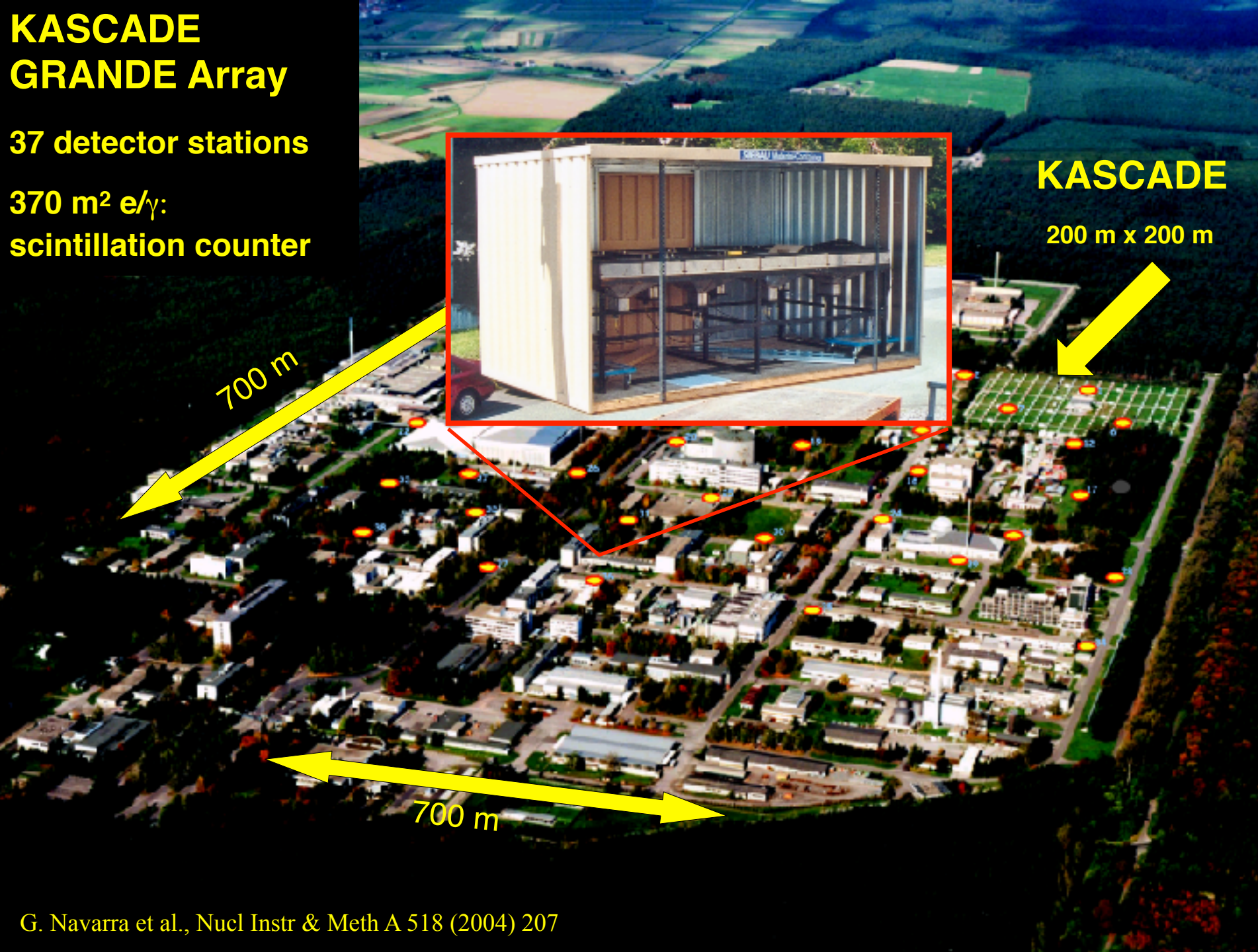
e/ γ -Detectors, Run 1, Event 71089, 96-03-05 22:07:48.956078



KASCADE GRANDE Array

37 detector stations

370 m² e/γ:
scintillation counter



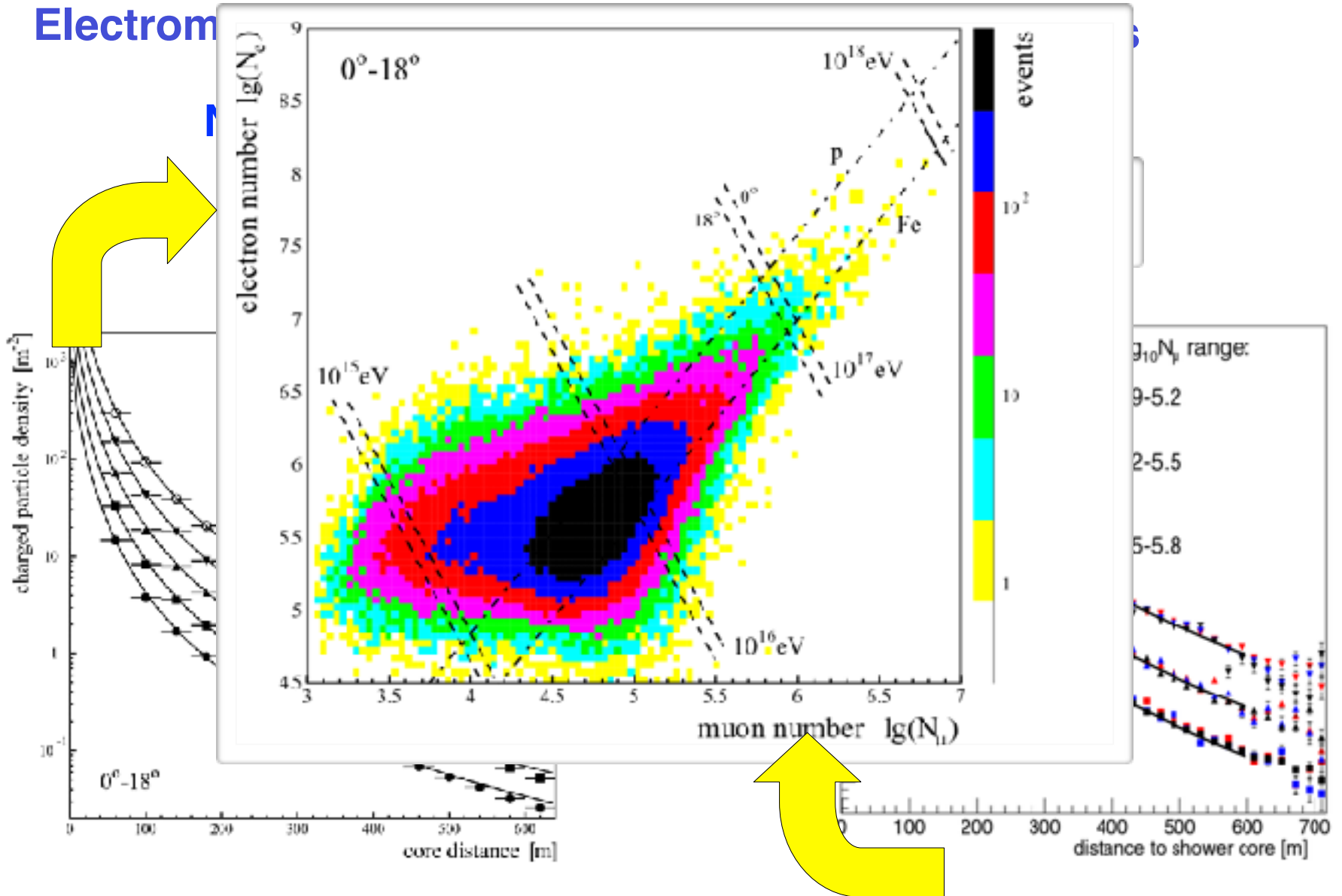
KASCADE

200 m x 200 m

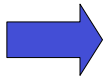
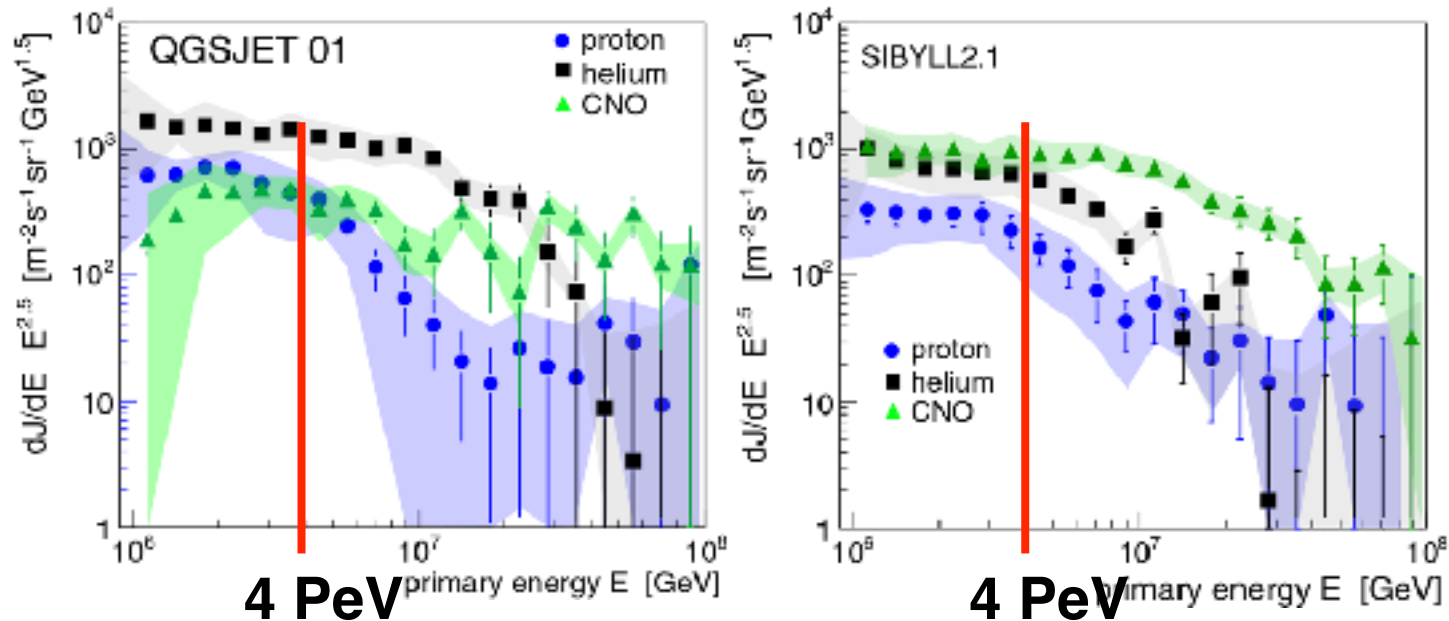
700 m

700 m

KASCADE-Grande – Lateral distributions



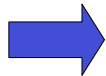
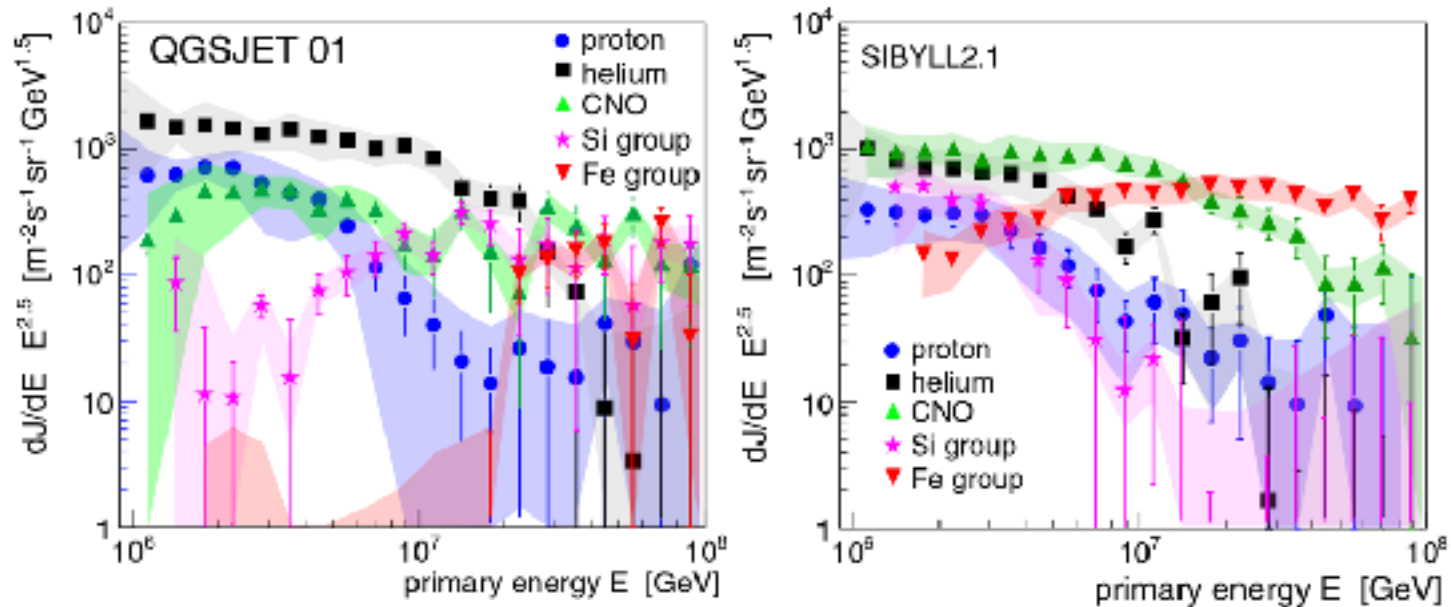
KASCADE: Energy spectra for elemental groups



Knee caused by cut-off for light elements

**Astrophysical interpretation limited
by description of interactions in the
atmosphere**

KASCADE: Energy spectra for elemental groups



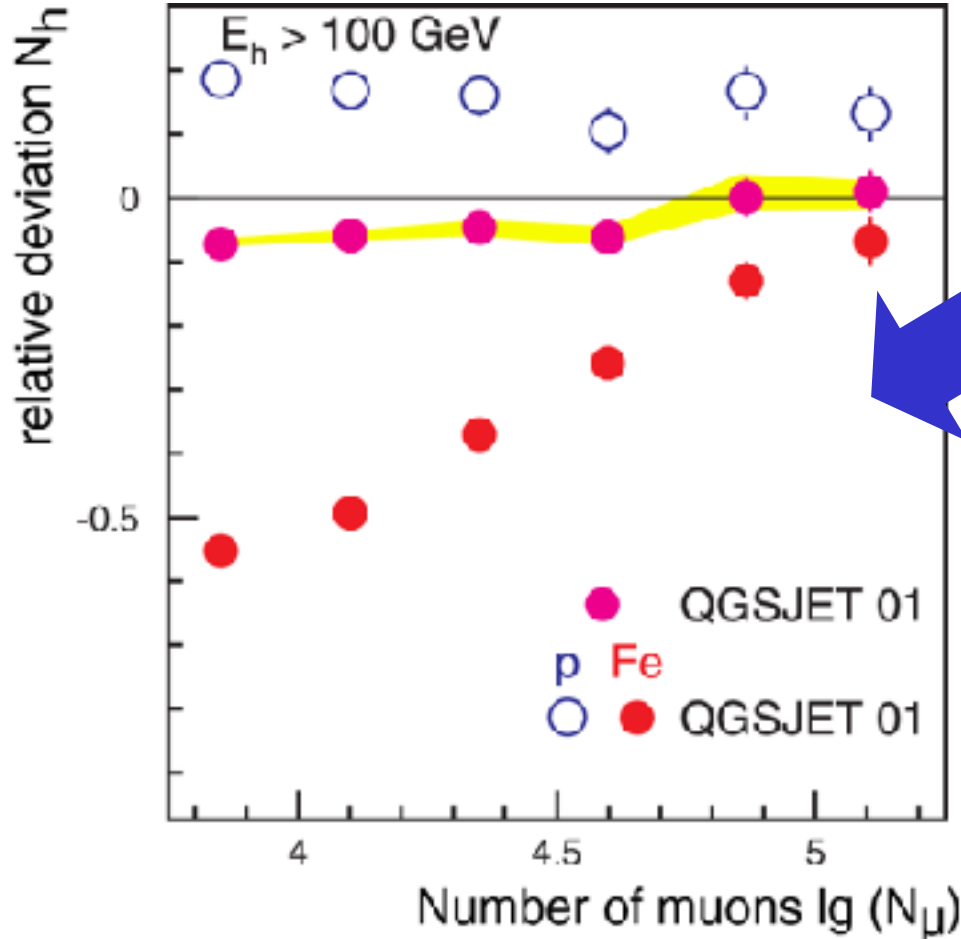
Knee caused by cut-off for light elements

**Astrophysical interpretation limited
by description of interactions in the
atmosphere**

Test of hadronic interaction models

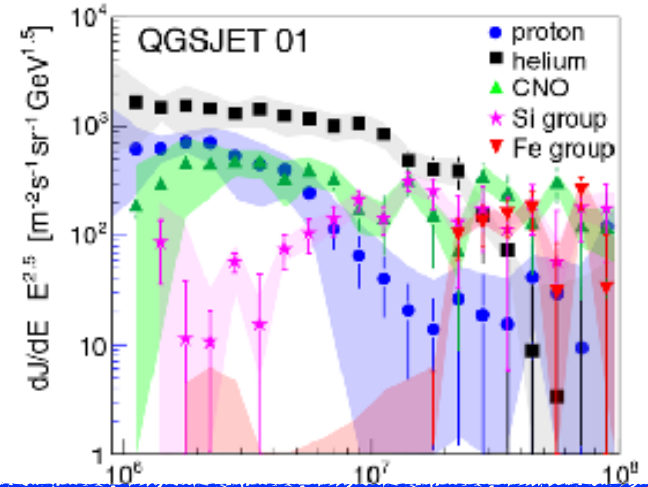
QGSJET 01

Number of hadrons vs. number of muons



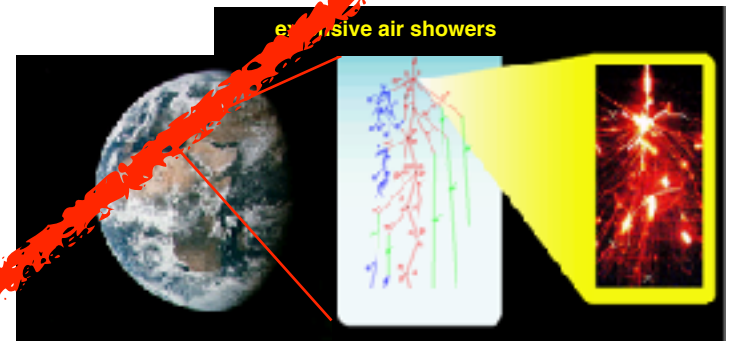
inconsistencies on 10% level

$N_e - N_\mu$ analysis



in literature:

ideas that knee is caused by new interactions in atmosphere
 —> energy is carried away by „invisible channels“



Electron, muon and hadron size spectra of EAS in the "knee" region

R. Glasstetter^a and J.R. Hörandel^a for the KASCADE Collaboration^a

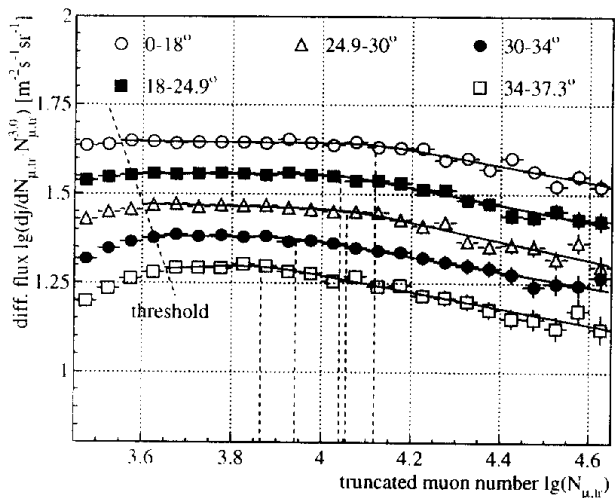
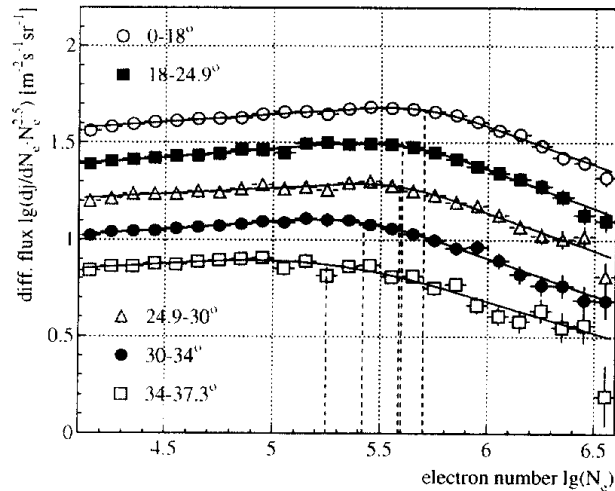


Figure 1. Electromagnetic (top) and muonic (bottom) shower size spectra for different zenith angle bins.

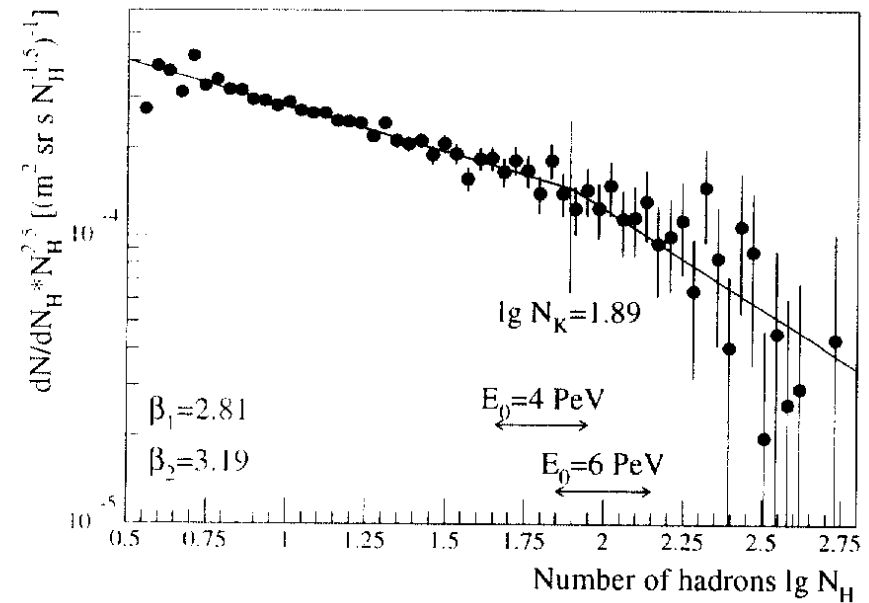


Figure 3. Hadronic shower size spectrum.

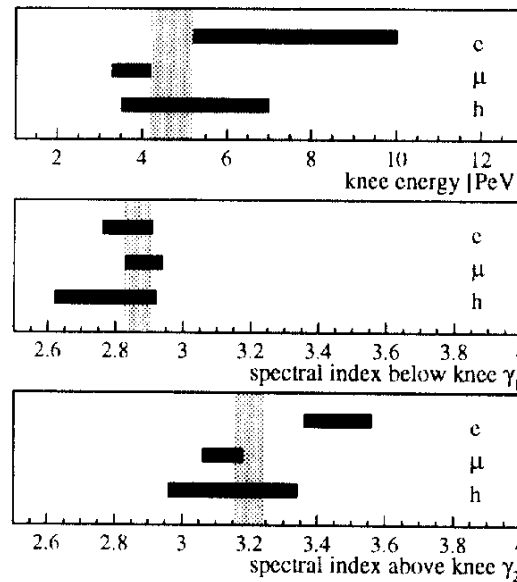
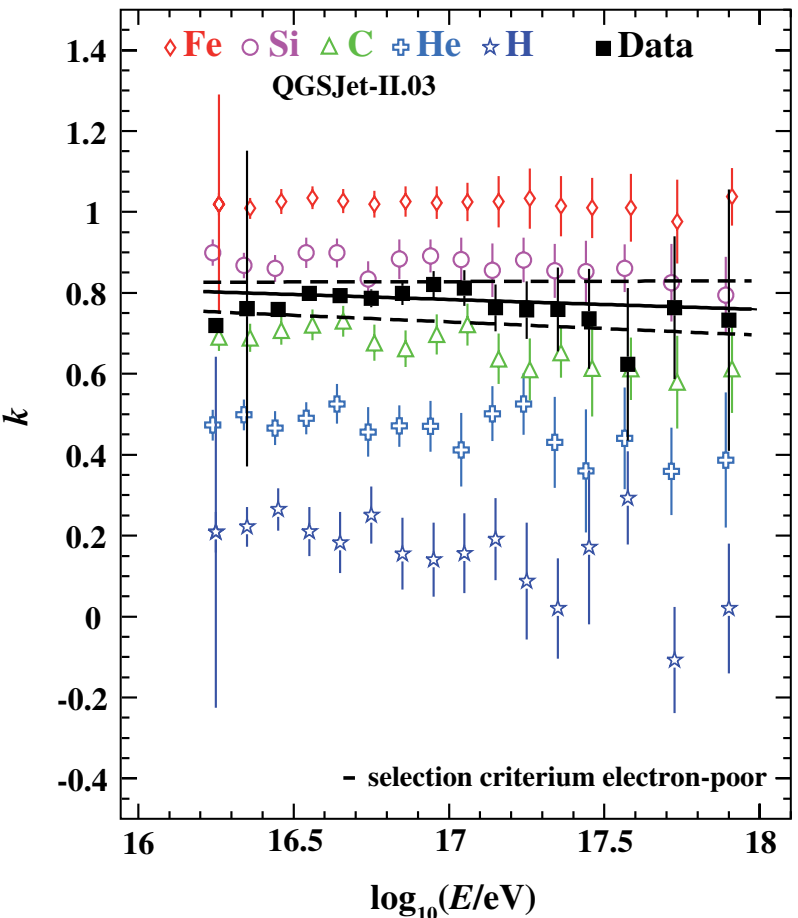


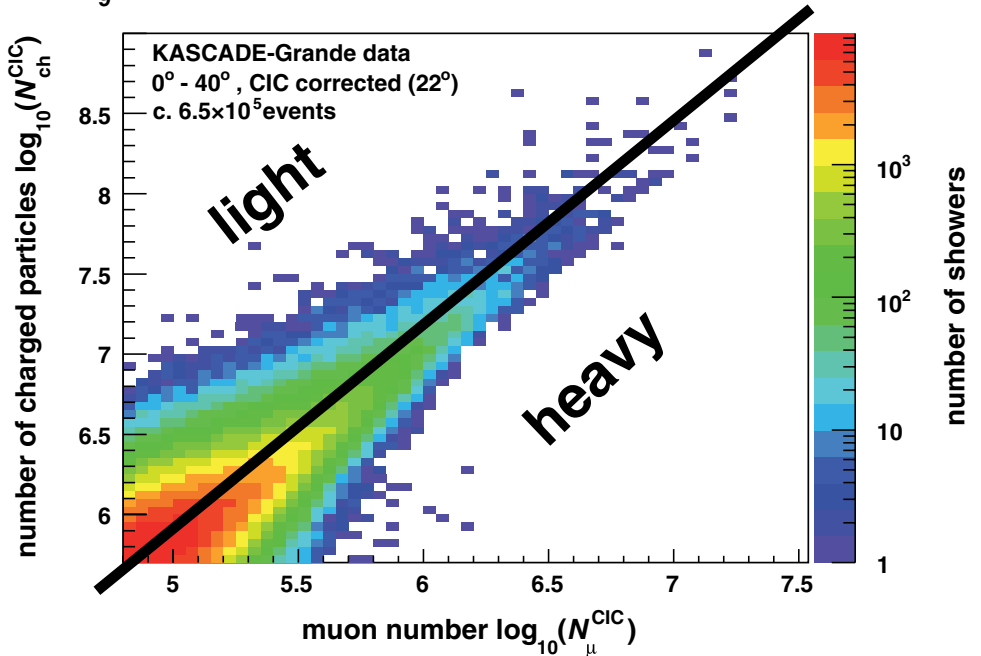
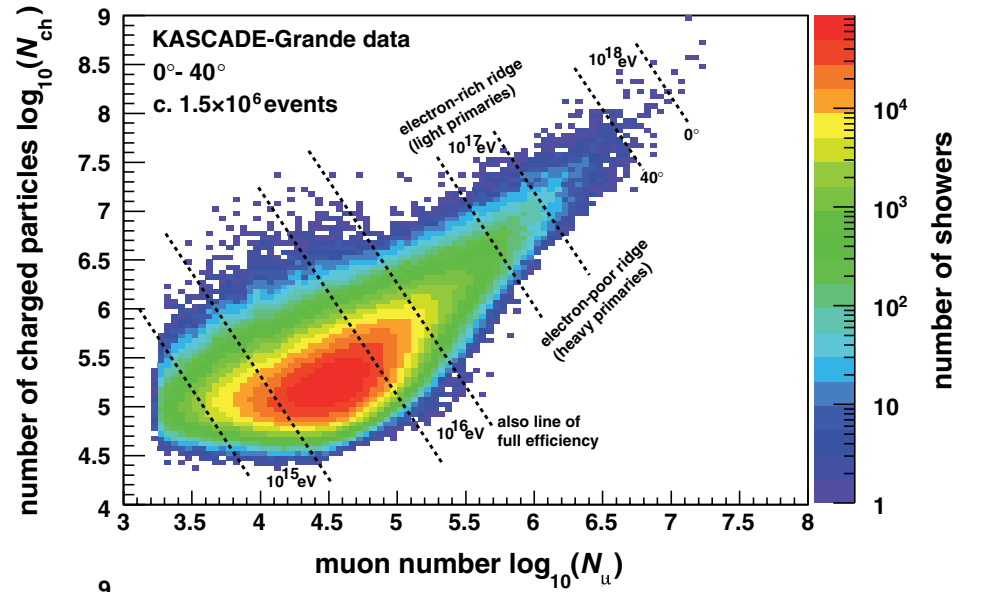
Figure 4. Knee position and spectral indices.

knee observed in all components, electromagnetic, muonic, and hadronic!

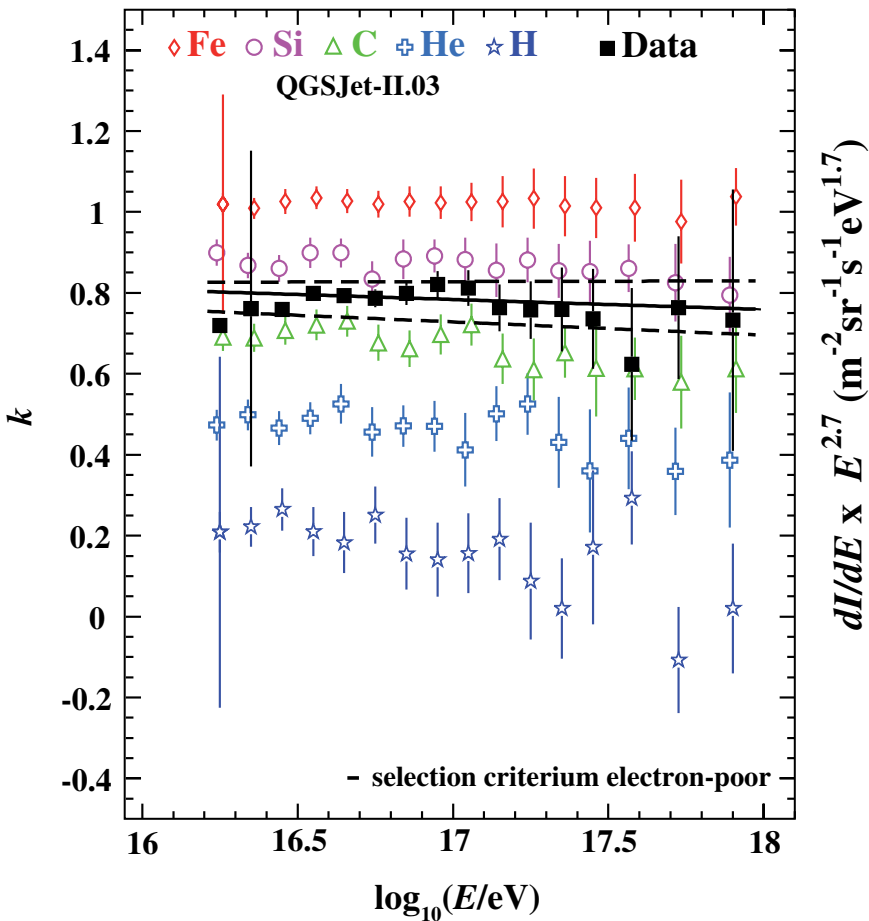
KASCADE-Grande



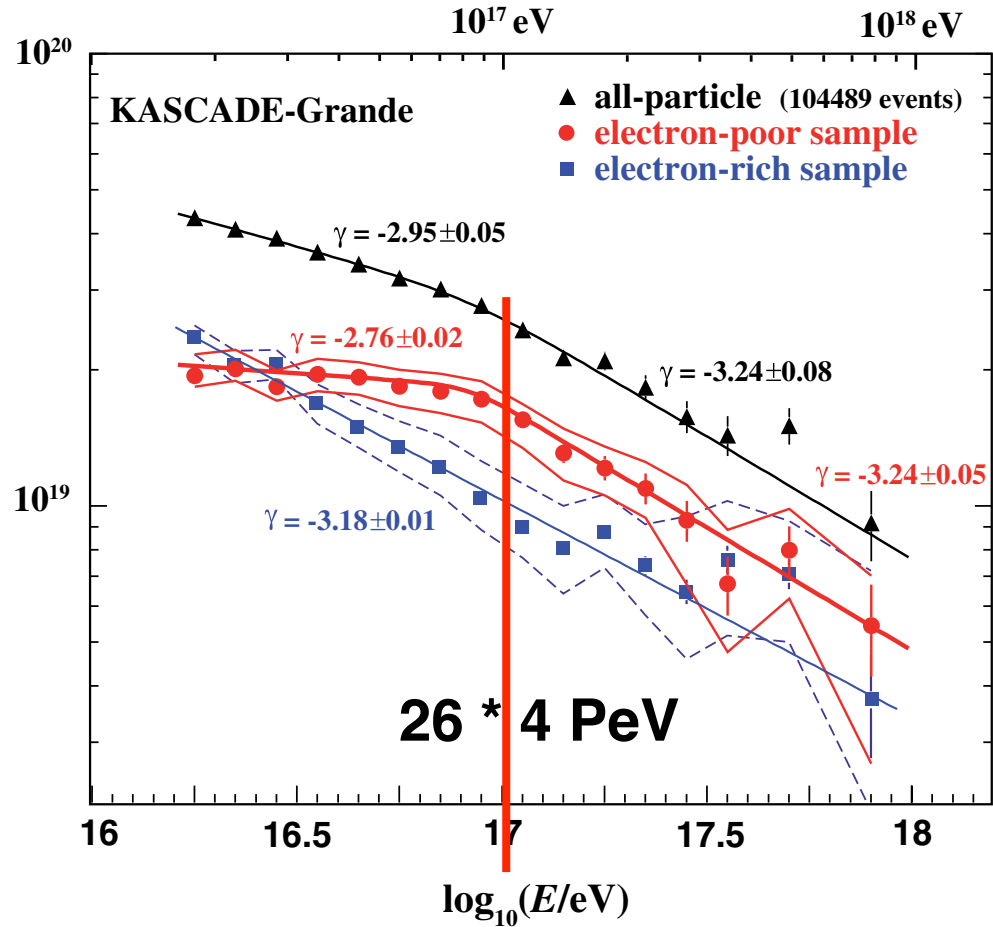
$$k = \frac{\log_{10}(N_{ch}/N_{\mu}) - \log_{10}(N_{ch}/N_{\mu})_H}{\log_{10}(N_{ch}/N_{\mu})_{Fe} - \log_{10}(N_{ch}/N_{\mu})_H}$$



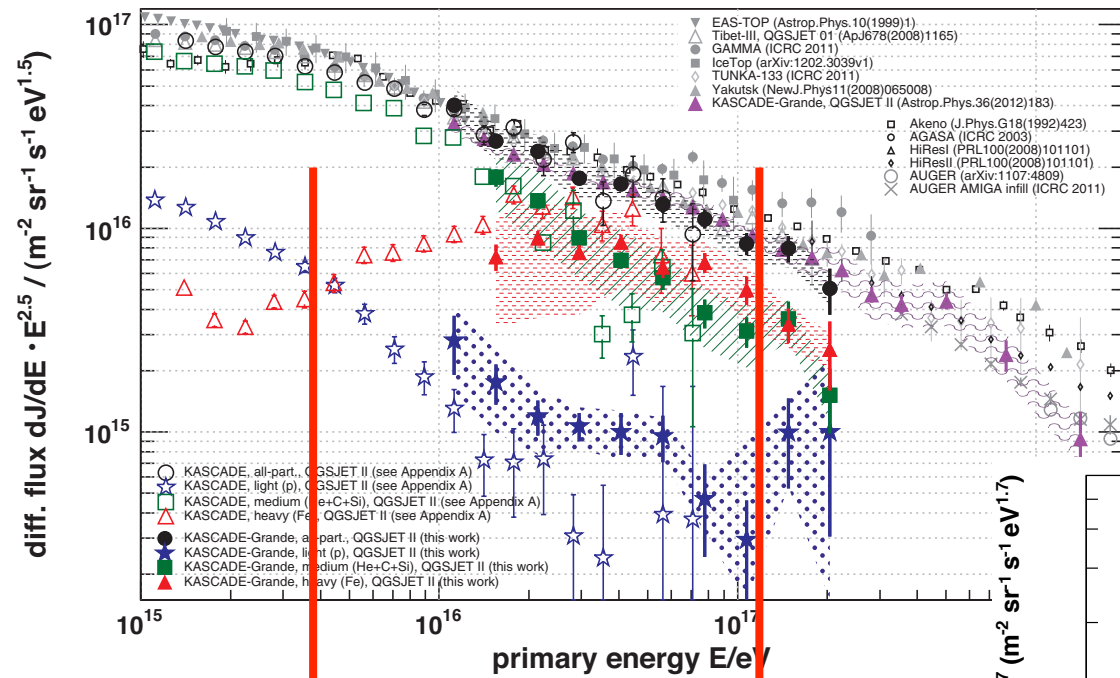
KASCADE-Grande



$$k = \frac{\log_{10}(N_{\text{ch}}/N_{\mu}) - \log_{10}(N_{\text{ch}}/N_{\mu})_{\text{H}}}{\log_{10}(N_{\text{ch}}/N_{\mu})_{\text{Fe}} - \log_{10}(N_{\text{ch}}/N_{\mu})_{\text{H}}}$$



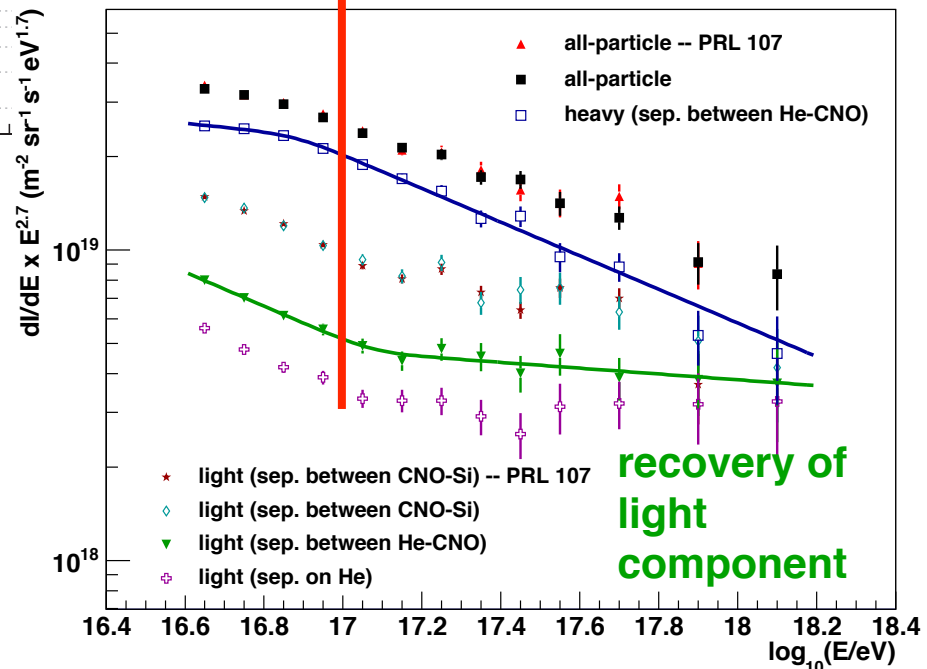
KASCADE-Grande



knee 4 PeV

26 * 4 PeV

26 * 4 PeV

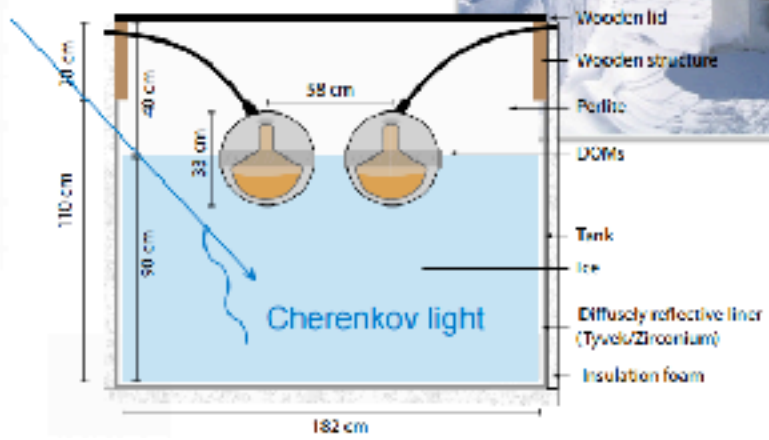
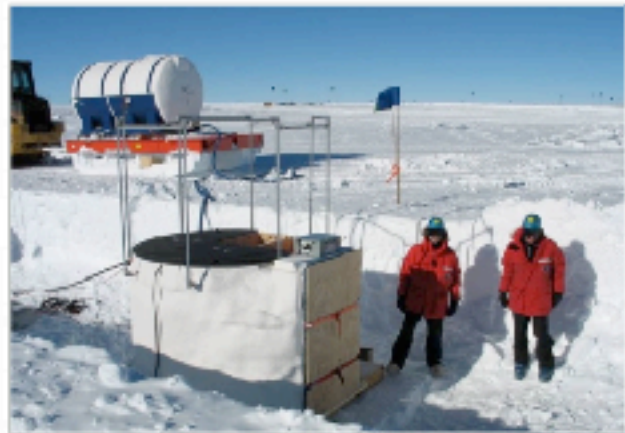
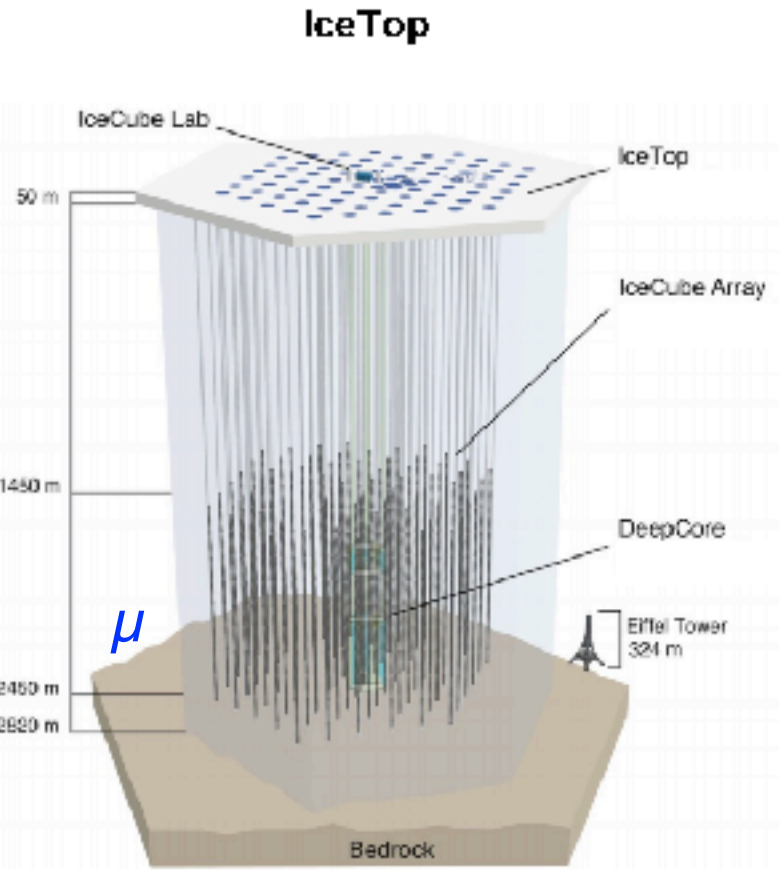


W.D. Apel et al., PRD 87 (2013) 081101

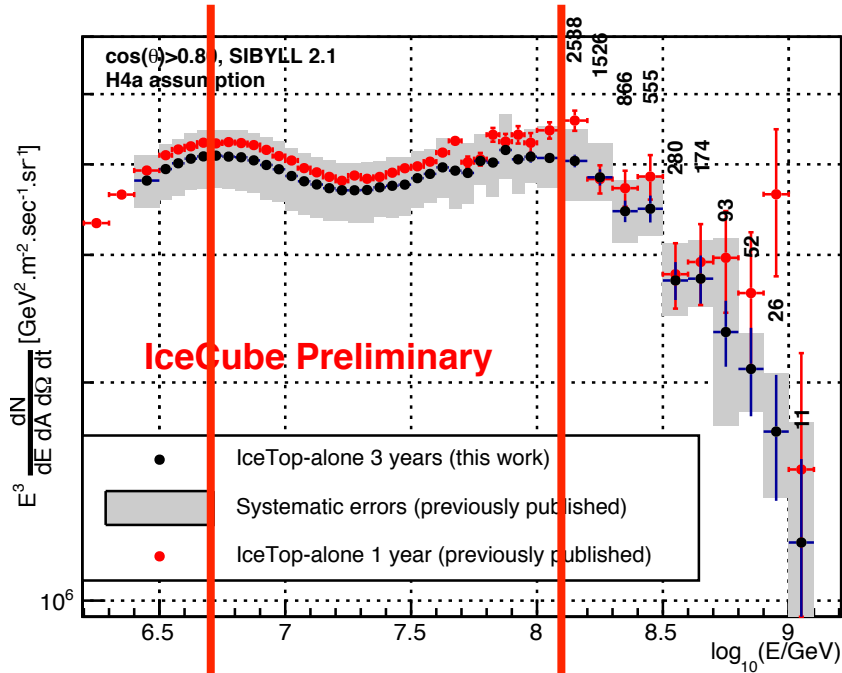
W.D. Apel et al., Astropart. Phys 47 (2013) 54

Ice Cube - Ice Top

$e/m (+\mu)$



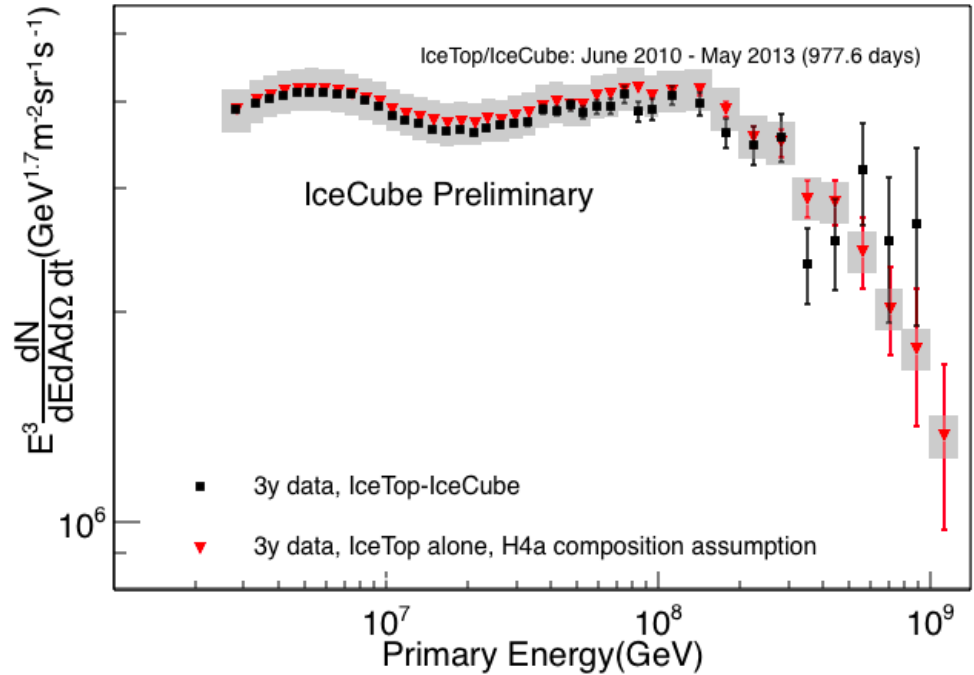
Ice Cube - Ice Top



5 PeV

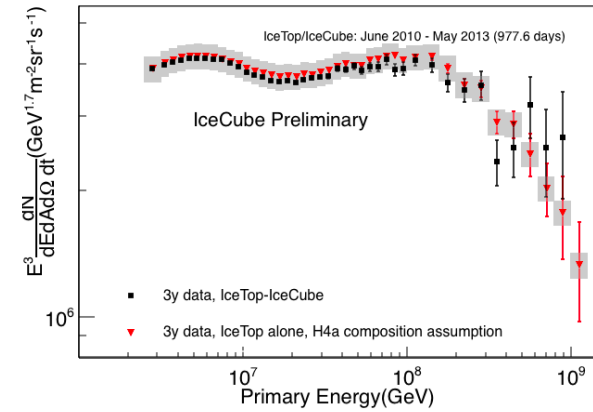
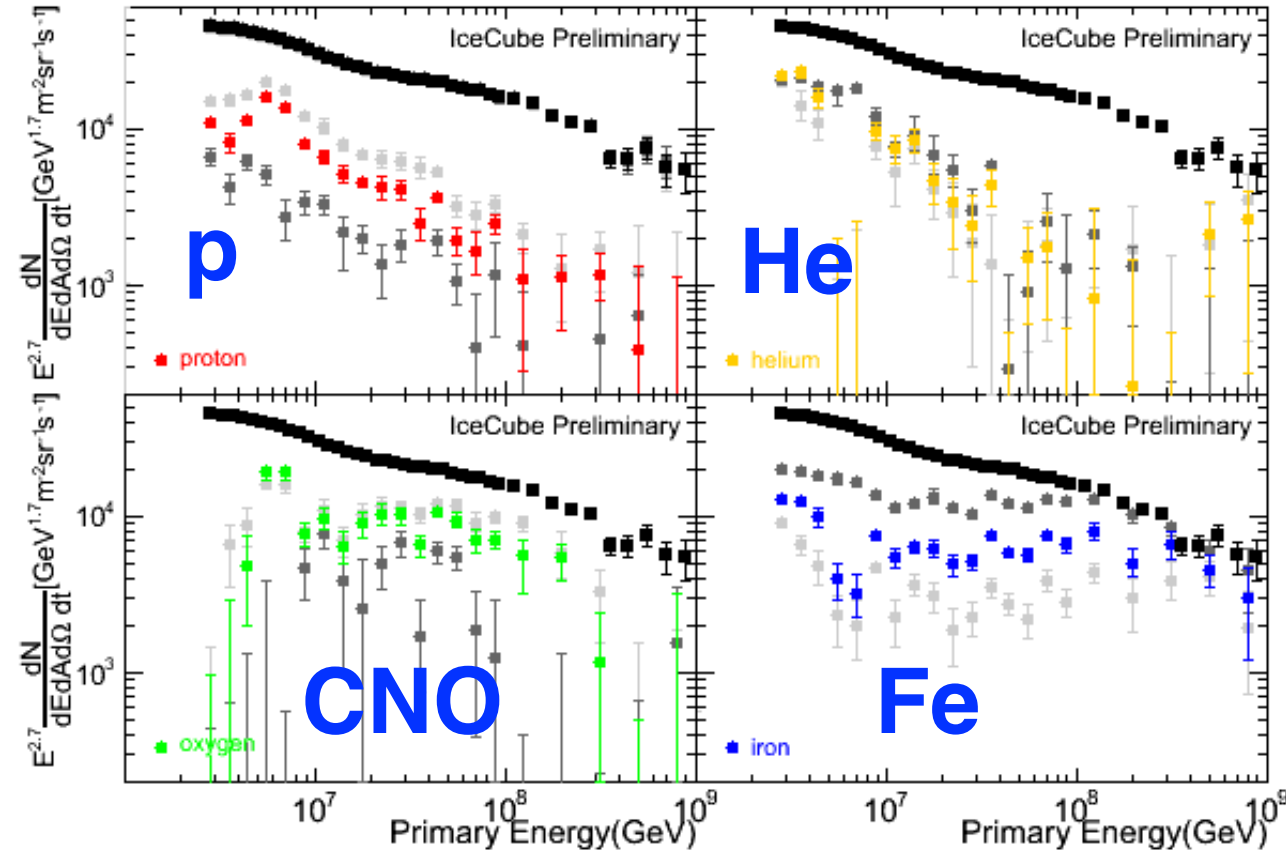
26*5 PeV

Ice Top only

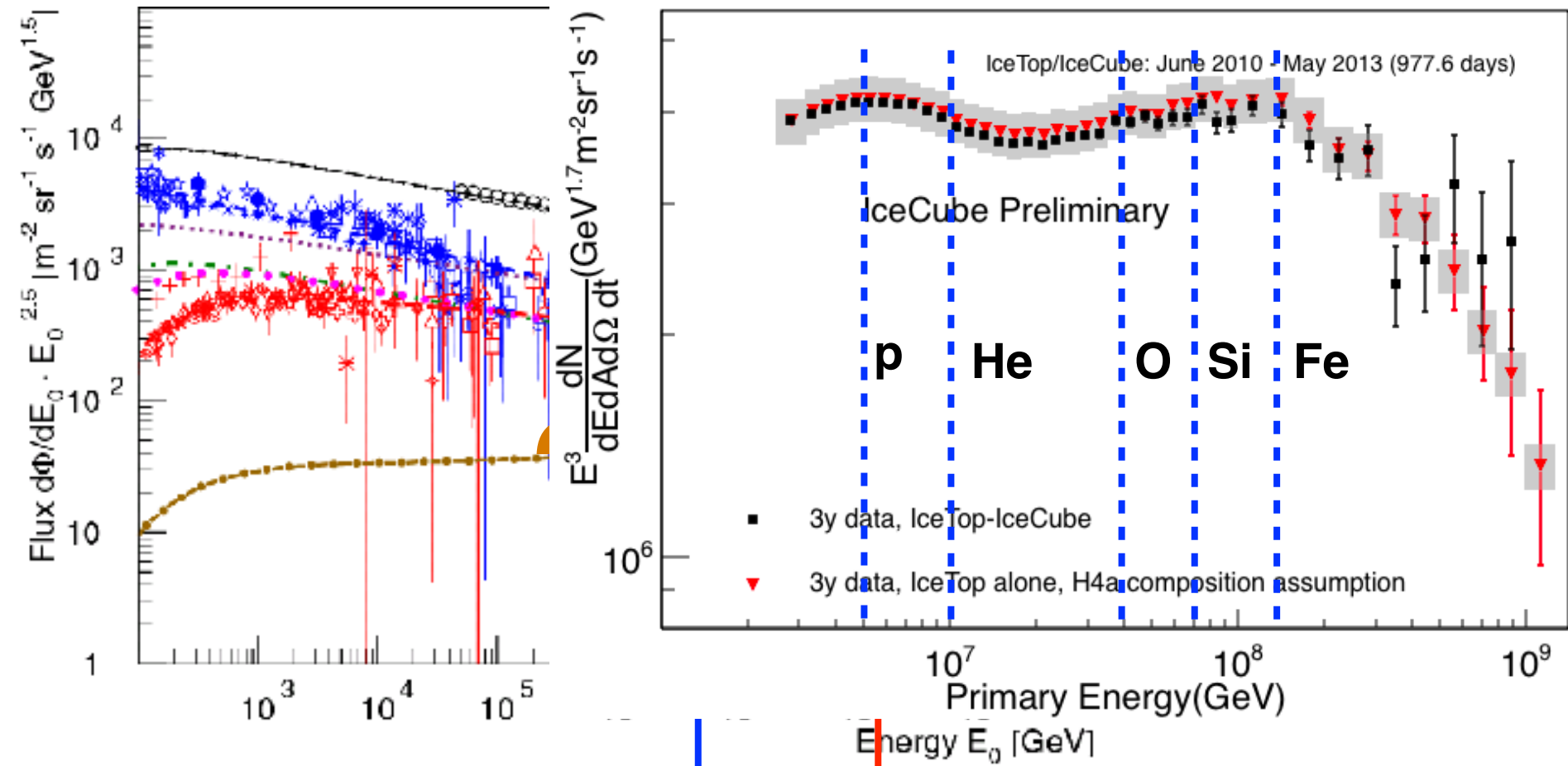


combined

Ice Cube - Ice Top



Cosmic-ray energy spectrum



TALE (TA low-energy extension)

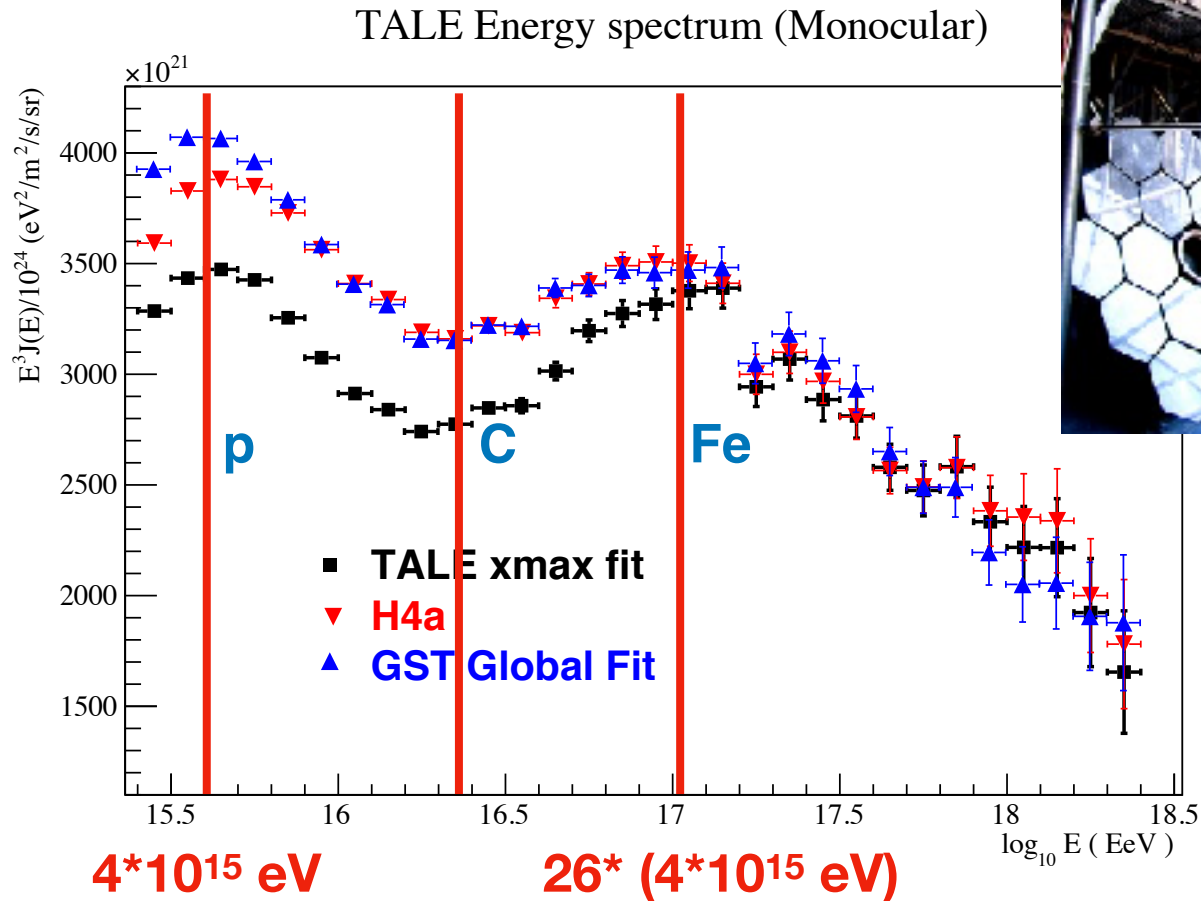
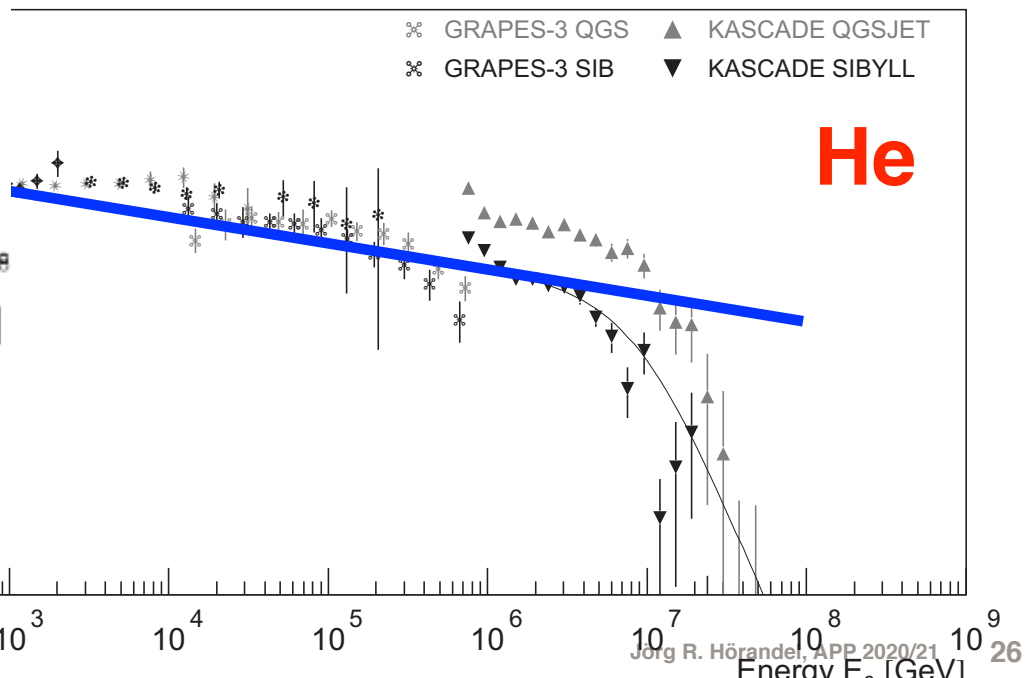
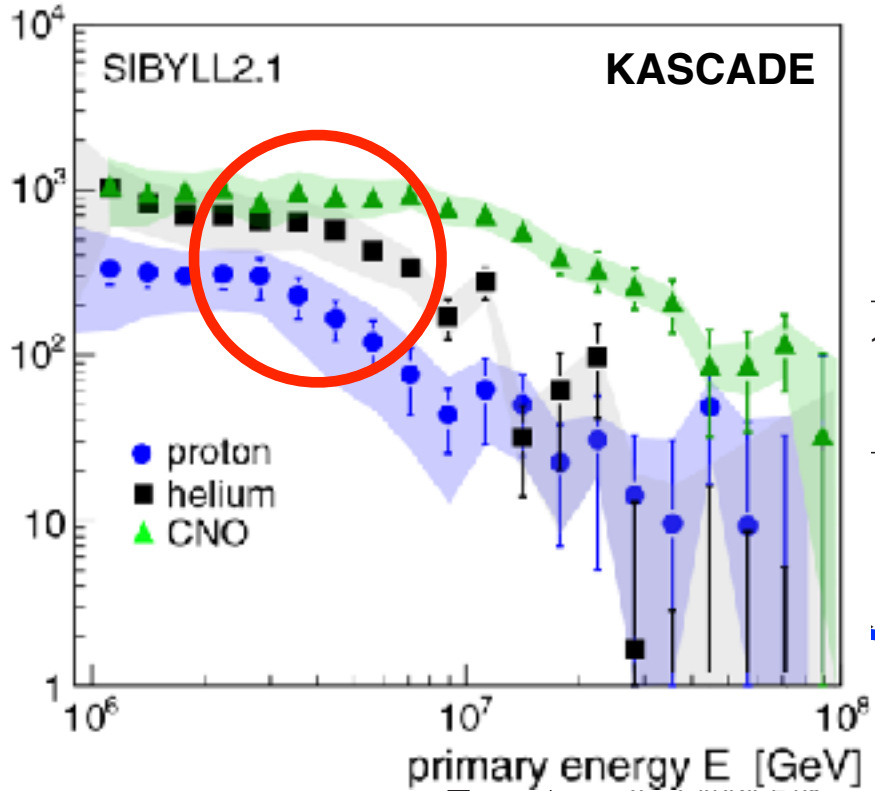
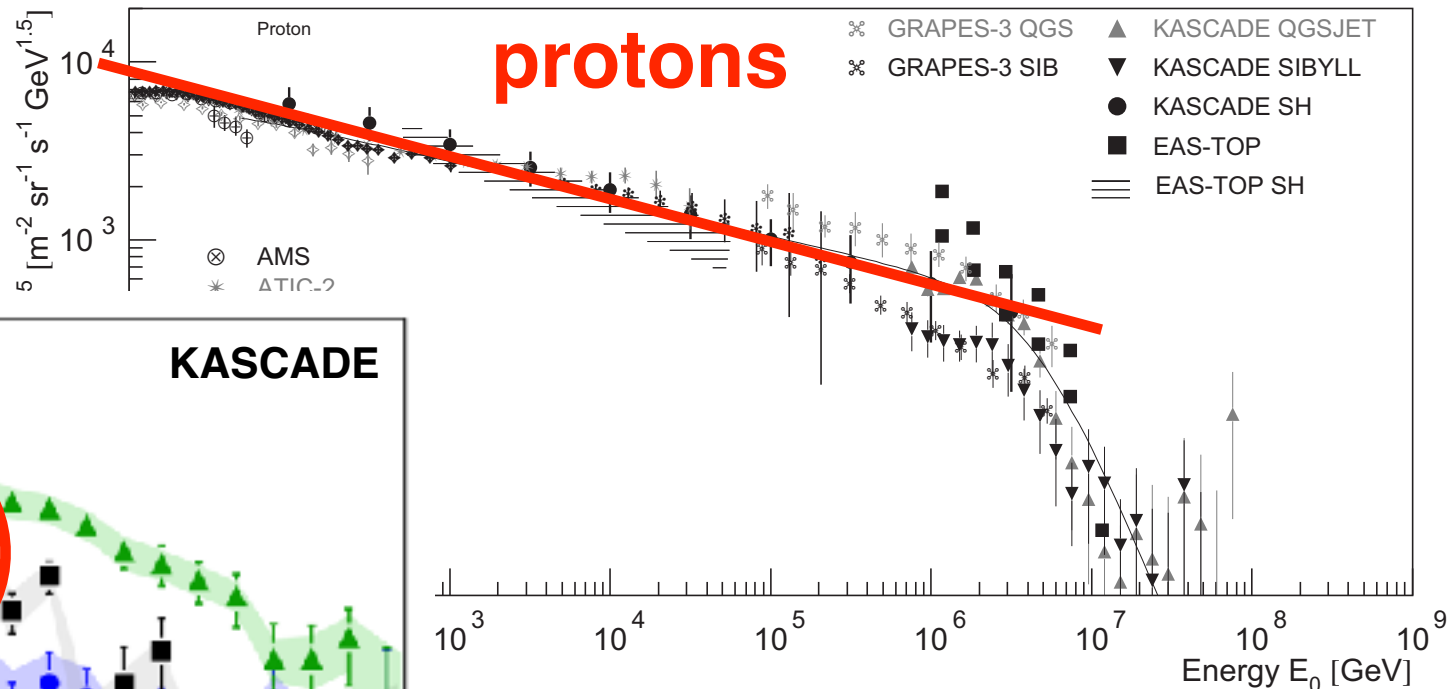
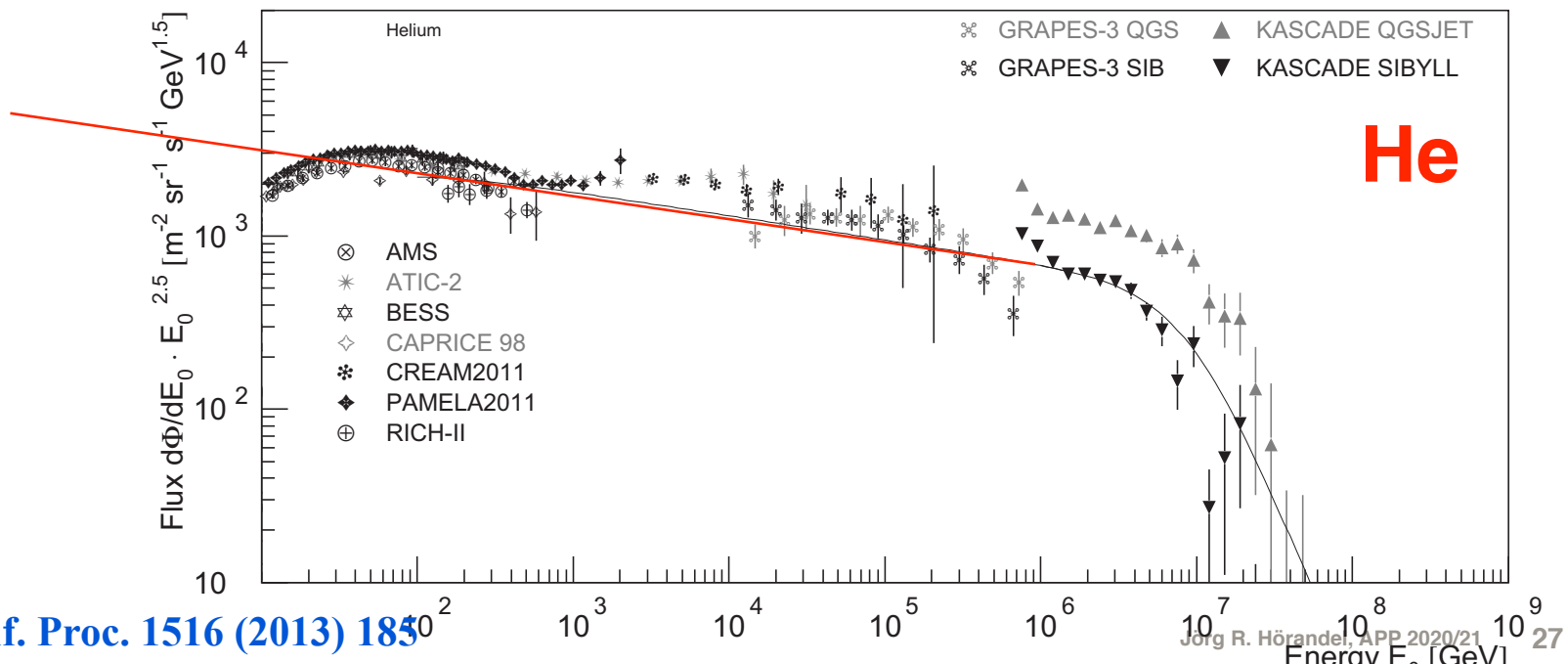
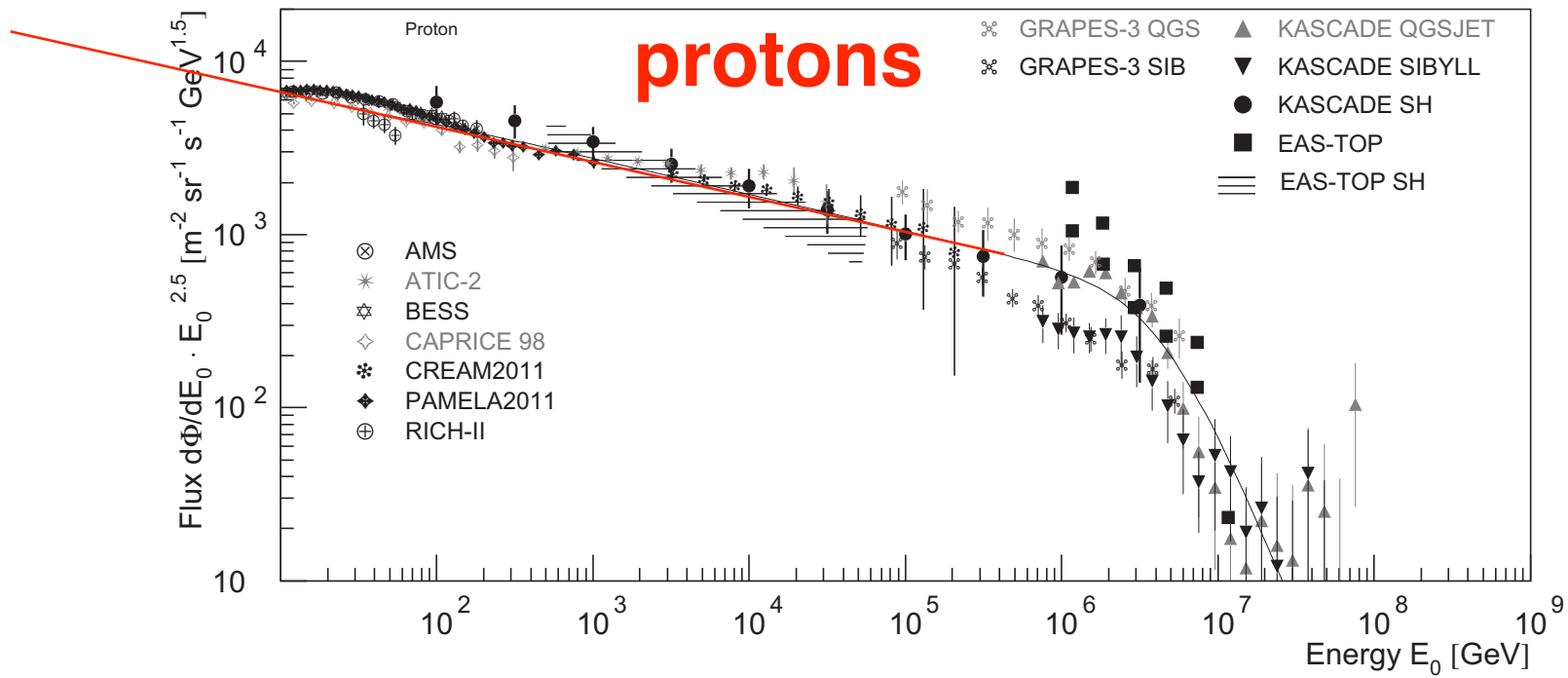


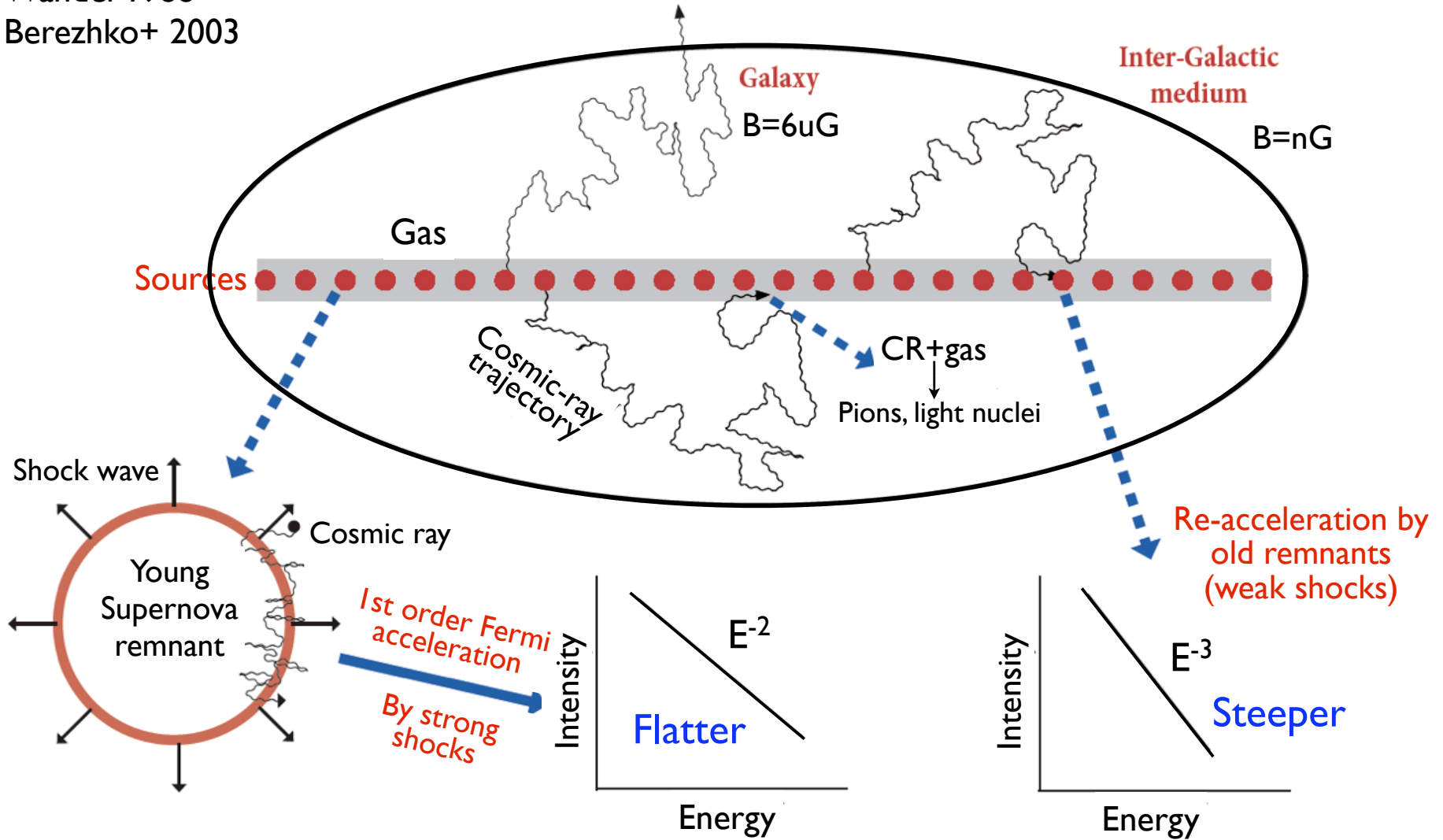
Figure 3: TALE Cosmic rays energy spectrum measured with TALE. The result is based on a QGSJet II-3 hadronic model assumption. A mixed primary composition given by the H4a, and “global fit” models, as well as a TALE derived mix was used in the calculations.

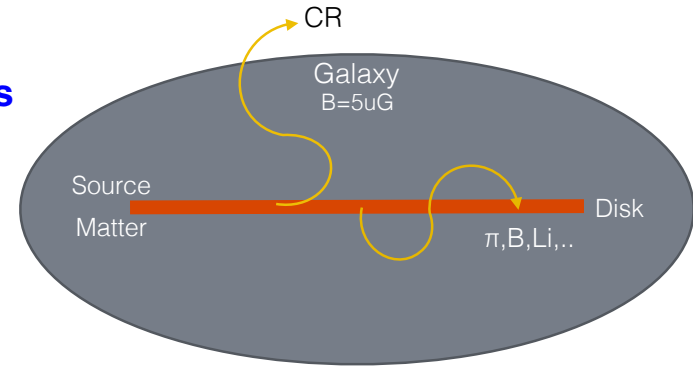




The re-acceleration model

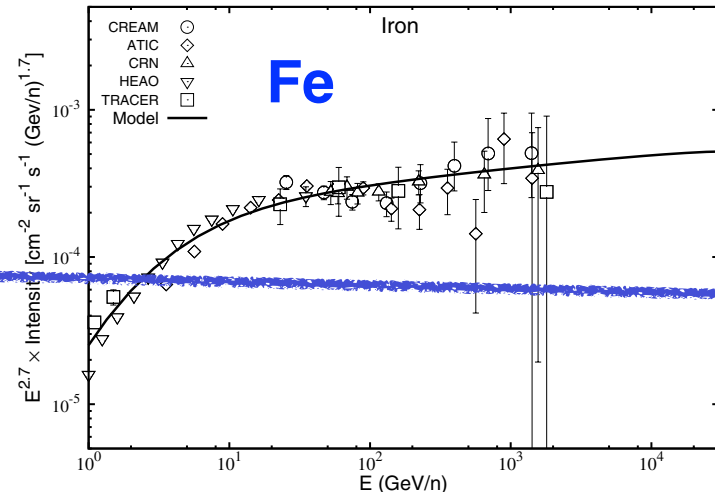
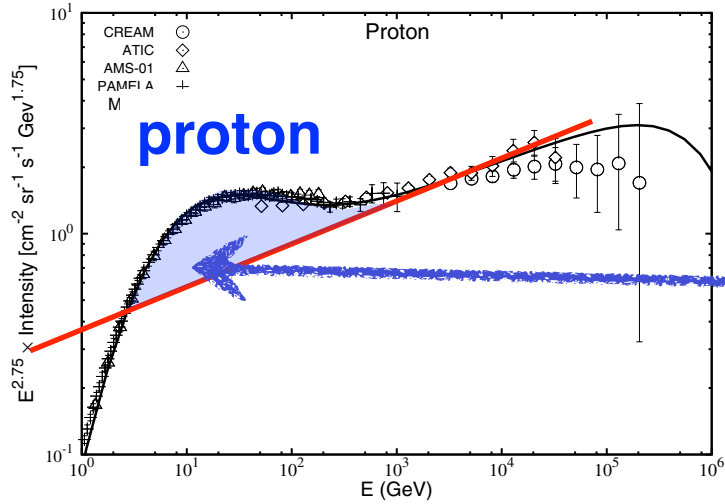
Wandel 1988
Berezhko+ 2003



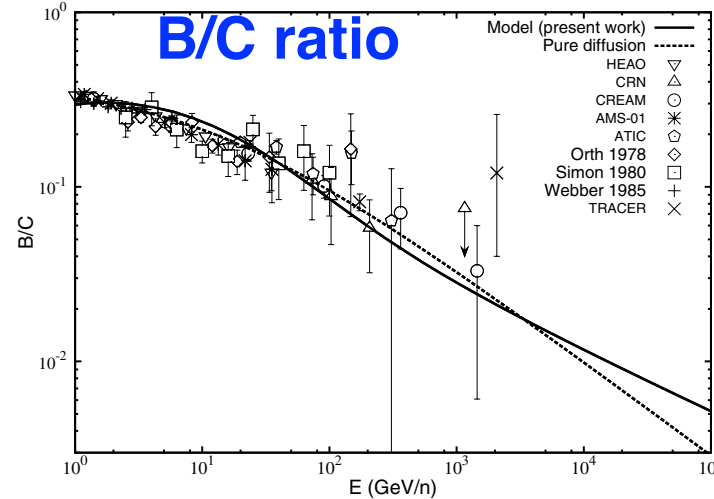
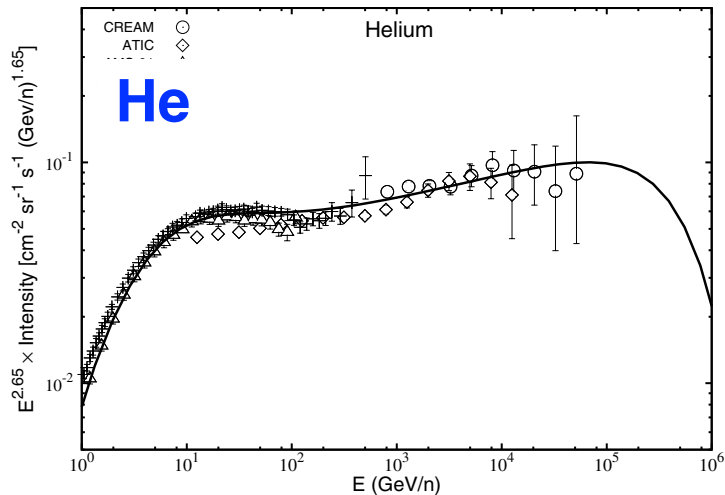


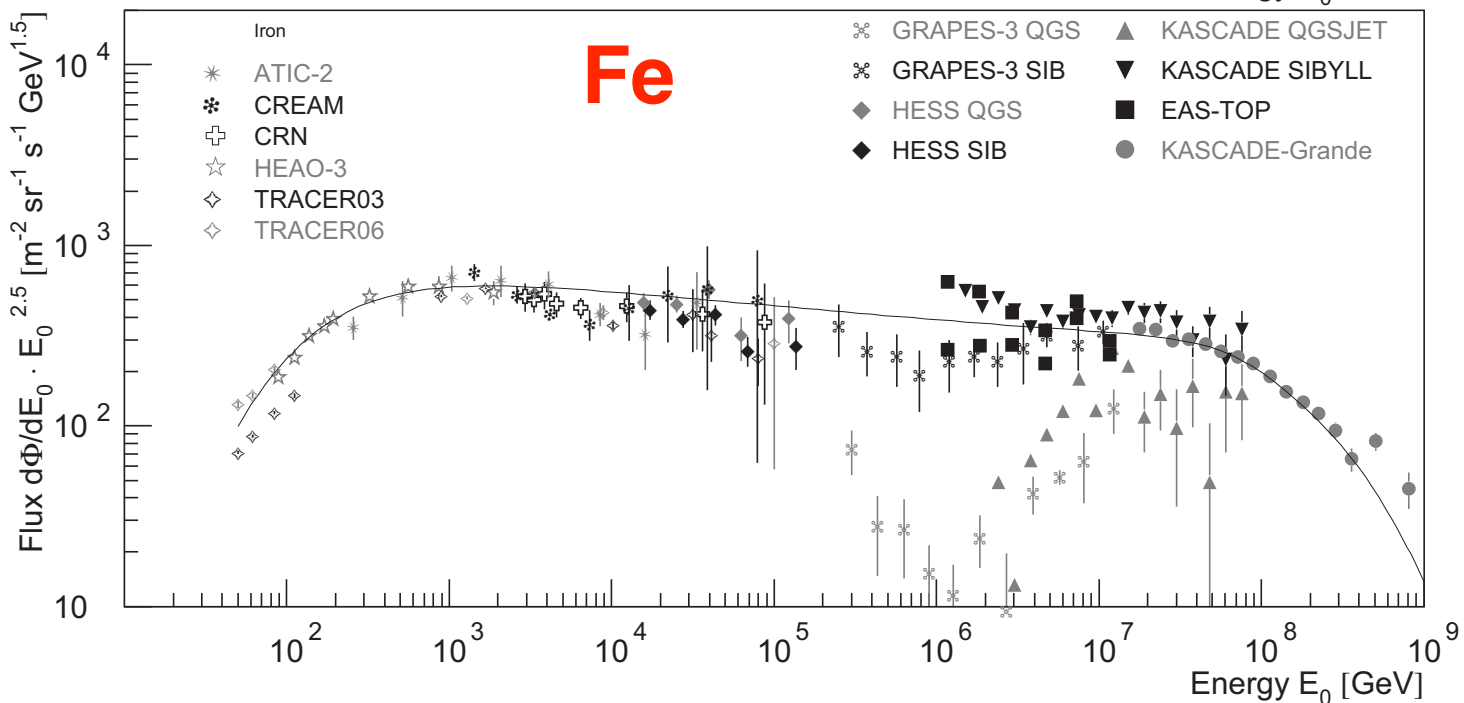
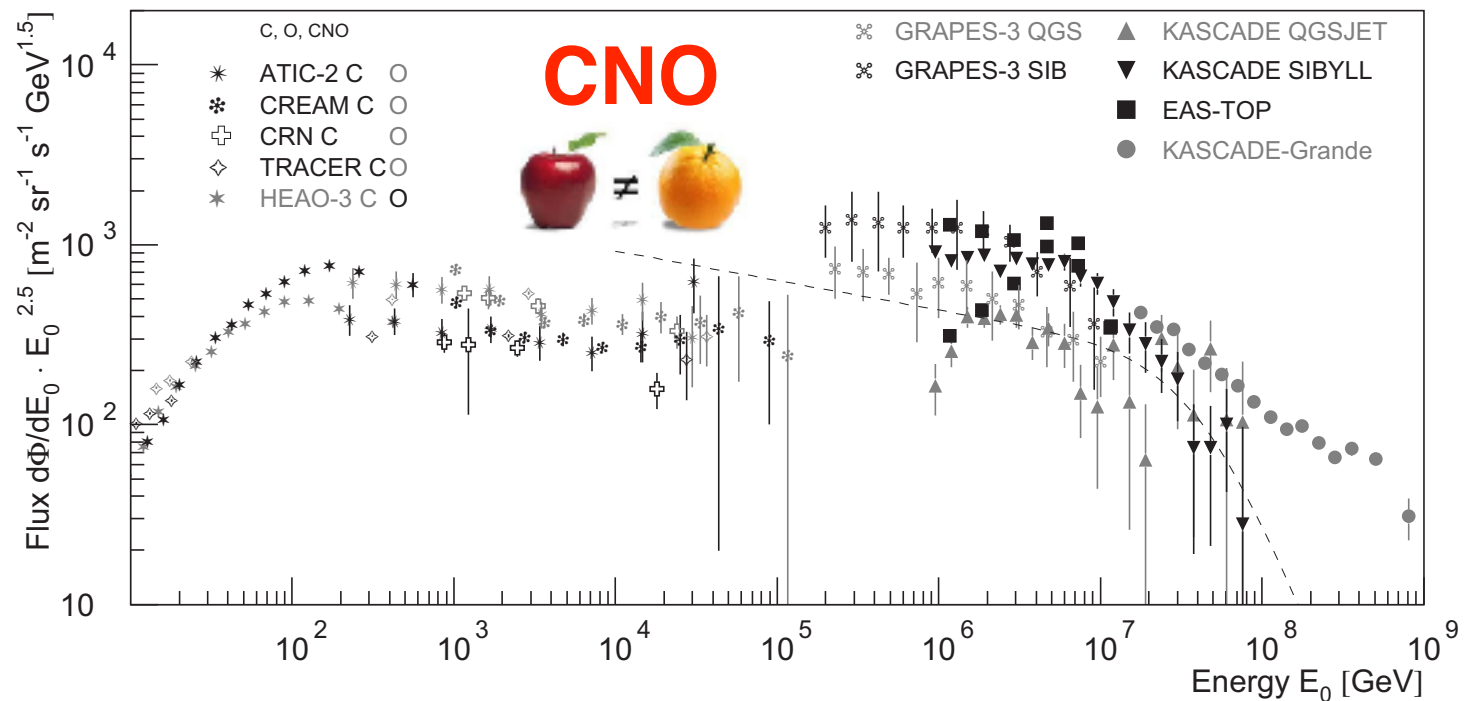
GeV-TeV cosmic-ray spectral anomaly as due to reacceleration by weak shocks in the Galaxy*

Satyendra Thoudam and Jörg R. Hörandel



bump due to reacceleration by weak shocks in the Galaxy





Transport equation for cosmic rays in the Galaxy

diffusion

energy loss (Bethe Bloch)

loss through interactions with ISM (spallation)

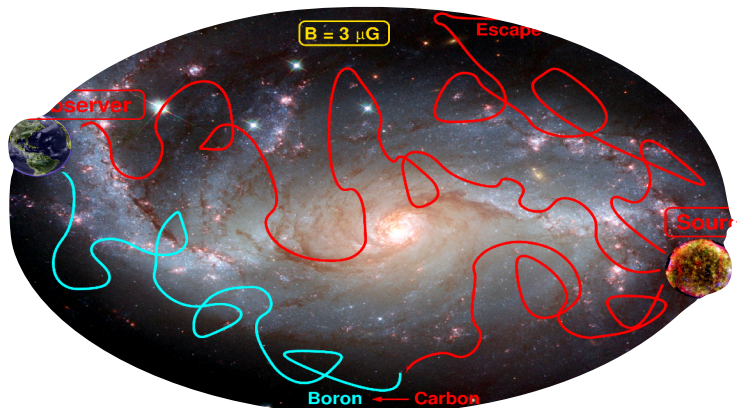
loss through radioactive decay

$$\frac{\partial N_i}{\partial t} = \nabla(D_i \nabla N_i) - \frac{\partial}{\partial E}(b_i N_i) - n\nu\sigma_i N_i - \frac{N_i}{\gamma\tau_i} + Q_i + \sum_{j>i} n\nu\sigma_{ij} N_j + \sum_{j>i} \frac{N_j}{\gamma_j\tau_{ij}}$$

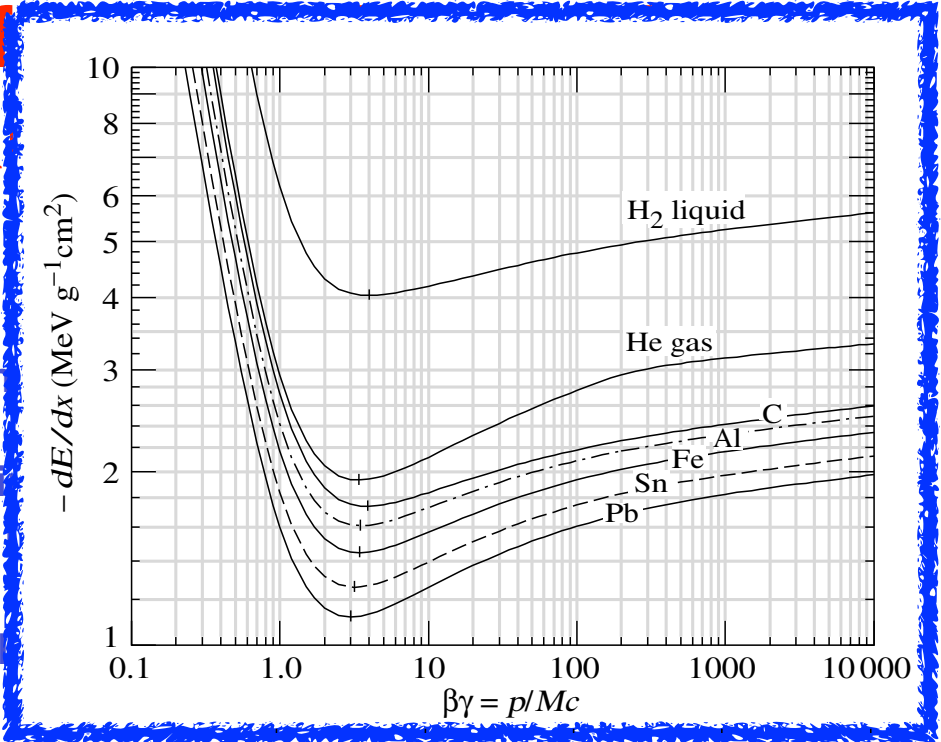
source term

production through spallation of heavy nuclei

production through decay of heavy nuclei



Transport equation for Galaxies



diffusion

energy loss (Bethe Bloch)

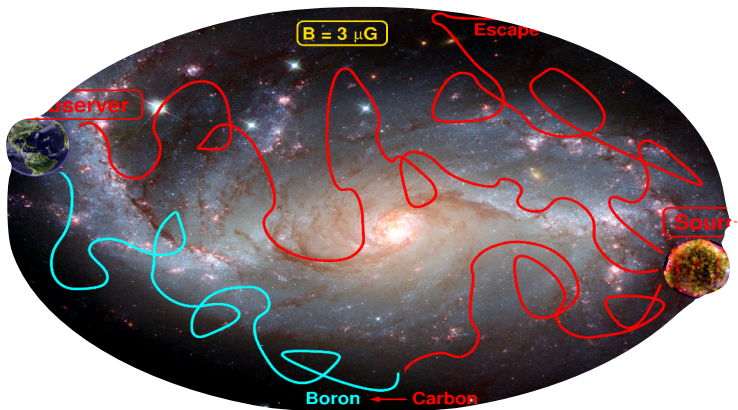
loss through with

$$\frac{\partial N_i}{\partial t} = \nabla(D_i \nabla N_i) - \frac{\partial}{\partial E} (\cancel{v_i N_i}) - n \nu \sigma_i N_i - \frac{N_i}{\gamma \tau_i} + Q_i + \sum_{j>i} n \nu \sigma_{ij} N_j + \sum_{j>i} \frac{N_j}{\gamma_j \tau_{ij}}$$

source term

production through spallation of heavy nuclei

production through decay of heavy nuclei



Transport equation for cosmic rays in the Galaxy

diffusion

energy loss (Bethe Bloch)

loss through interactions with ISM (spallation)

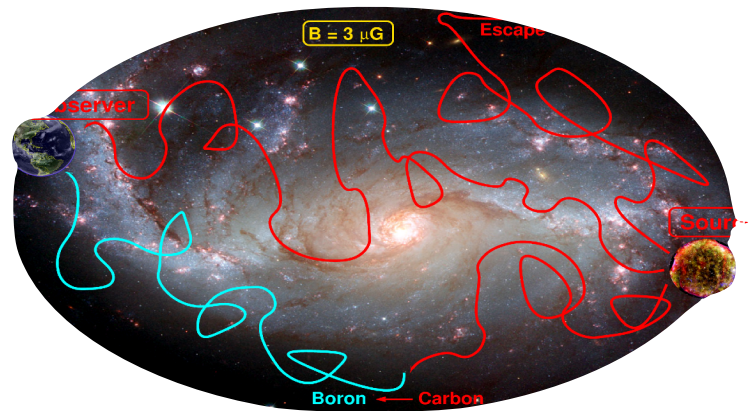
loss through radioactive decay

$$\frac{\partial N_i}{\partial t} = \nabla(D_i \nabla N_i) - \frac{\partial}{\partial E} (E N_i) - n \nu \sigma_i N_i - \frac{N_i}{\gamma \tau_i} + Q_i + \sum_{j>i} n \nu \sigma_{ij} N_j + \sum_{j>i} \frac{N_j}{\gamma_j \tau_{ij}}$$

source term

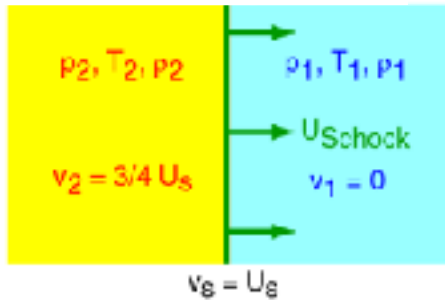
production through spallation of heavy nuclei

production through decay of heavy nuclei

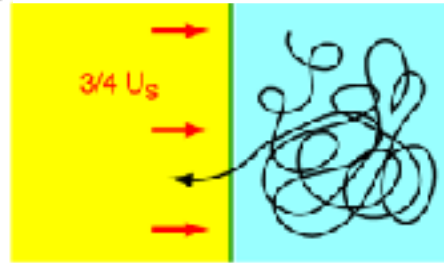


1st order Fermi acceleration at strong shock

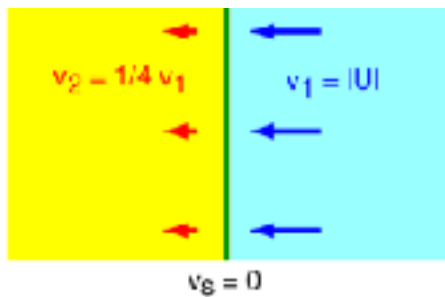
a) rest system of unshocked ISM



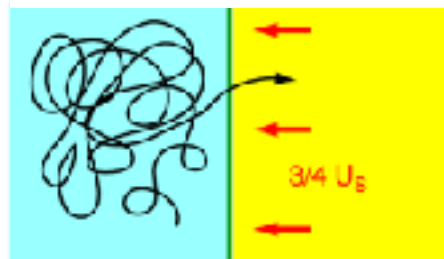
b) rest system of unshocked ISM



c) rest system of shock front



d) rest system of shocked ISM

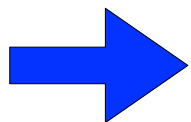


Bell, Blanford, Ostriker (1978)



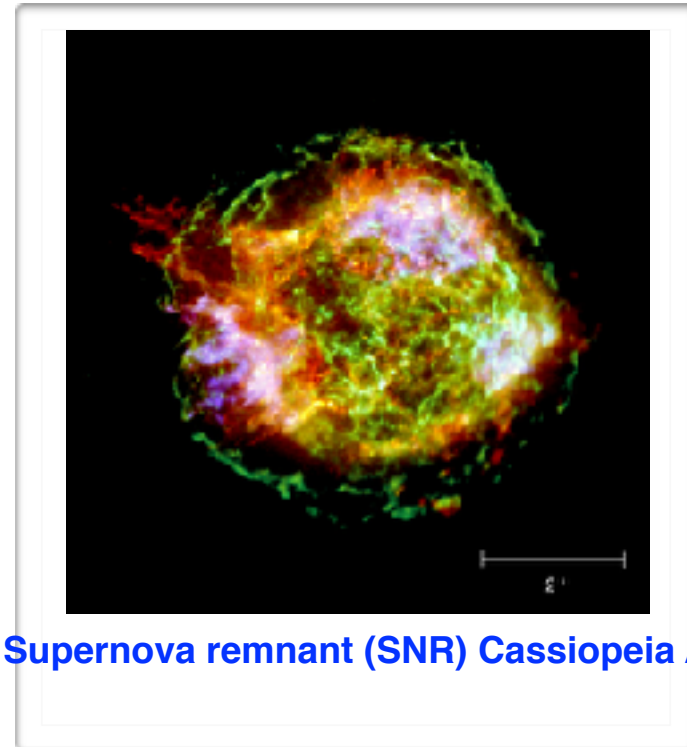
energy gain

$$\frac{\Delta E}{E} \propto \frac{U_s}{c}$$



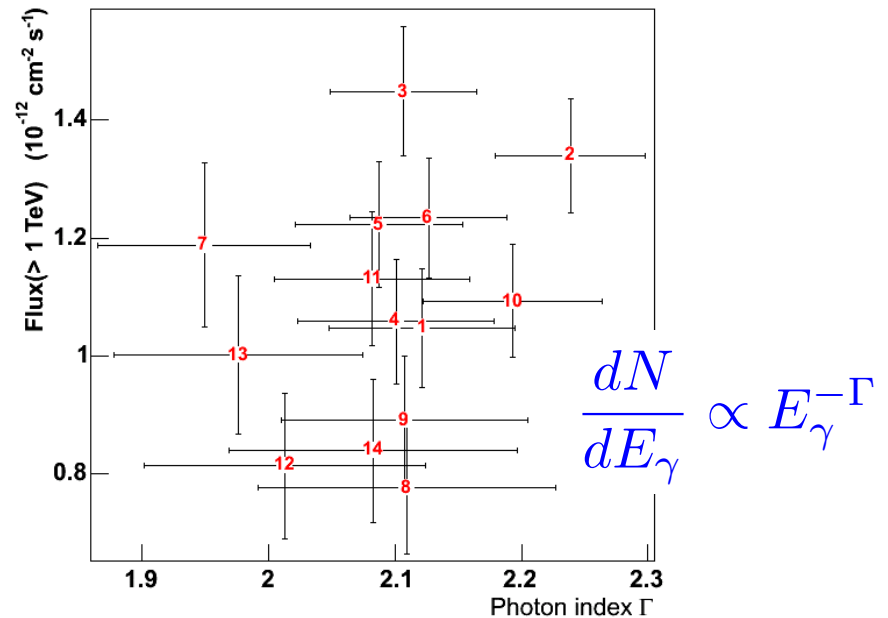
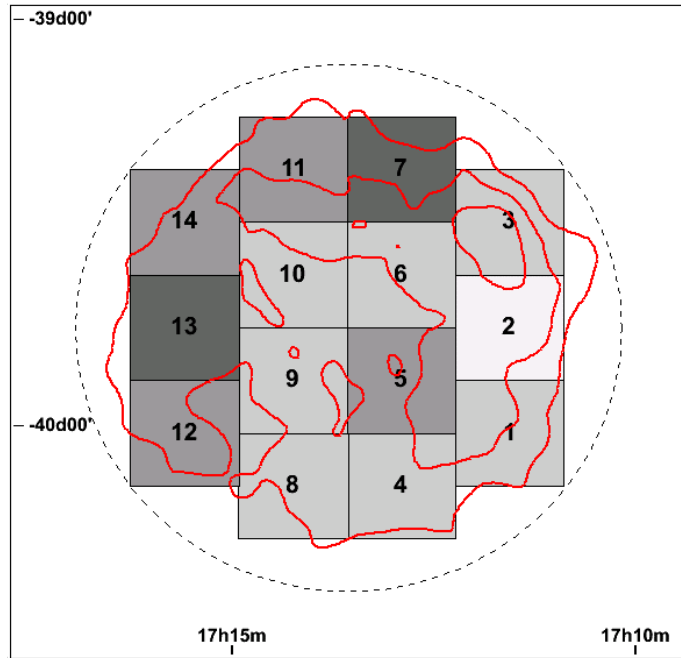
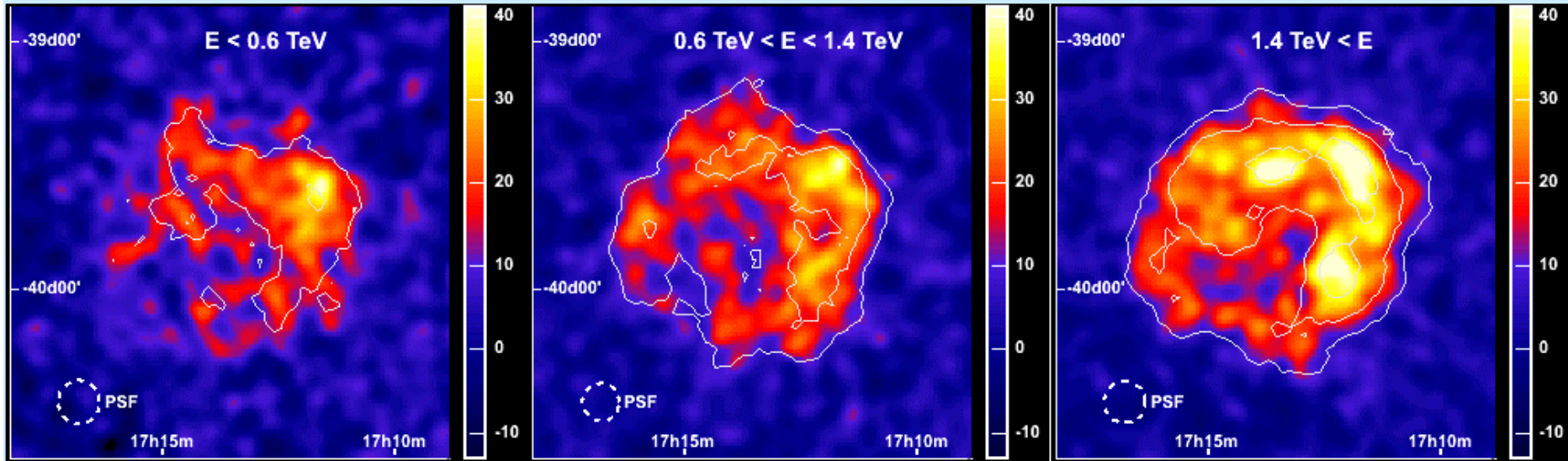
$$N(E) dE \propto E^{-2} dE$$

power law with spectral index
-2.0 ... -2.1



Supernova remnant (SNR) Cassiopeia A

H.E.S.S. supernova remnant RXJ 1713



Acceleration of cosmic rays at SNR

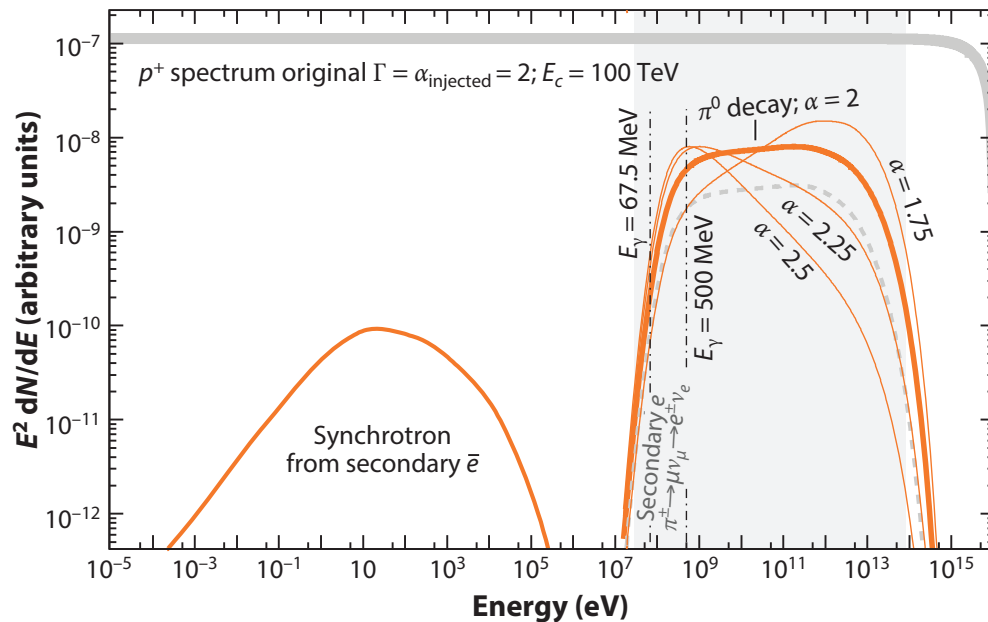


Figure 3
Spectral energy distribution of accelerated protons (power-law index $\alpha_{\text{injected}} = 2.0$ and cutoff at 100 TeV) and γ -rays resulting from inelastic collisions with interstellar material. The dominant emission into photons occurs via the decay $\pi^0 \rightarrow \gamma\gamma$ (solid orange curves). The γ -ray spectrum follows the parent protons' spectrum rather closely in the midenergy range and in the high-energy cutoff region. For all proton indices, the low-energy turnover is a characteristic feature of the pion-decay emission. Also shown is the spectrum of electrons resulting from the inelastic proton-proton interactions via the decay chain $\pi^\pm \rightarrow \mu + \nu_\mu \rightarrow e^\pm \nu_e$ (dashed gray curve). For the synchrotron emission from these so-called secondary electrons, a source with age $t_{\text{age}} = 1,000$ years and $B = 30 \mu\text{G}$ have been assumed. The shaded gray region shows the sensitive range of current γ -ray detectors (*Fermi*-LAT, imaging atmospheric Cherenkov detectors).

of energy in protons E_{pr} at the interaction site can be directly inferred from the γ -ray flux F_γ .

flux at Earth by an astrophysical accelerator that puts a fraction ϵ_{CR} of the energy output into the acceleration of protons:

$$F_\gamma(>100 \text{ MeV}) = 4.4 \times 10^{-7} \frac{\epsilon_{\text{CR}}}{\text{kpc}^2 \text{ s}^{-1}} \frac{E_{\text{pr}}}{10^{51} \text{ erg}}$$

In other words, if the distance d and the density of the interaction region n are known, the proton index α can be directly inferred from the γ -ray flux F_γ .

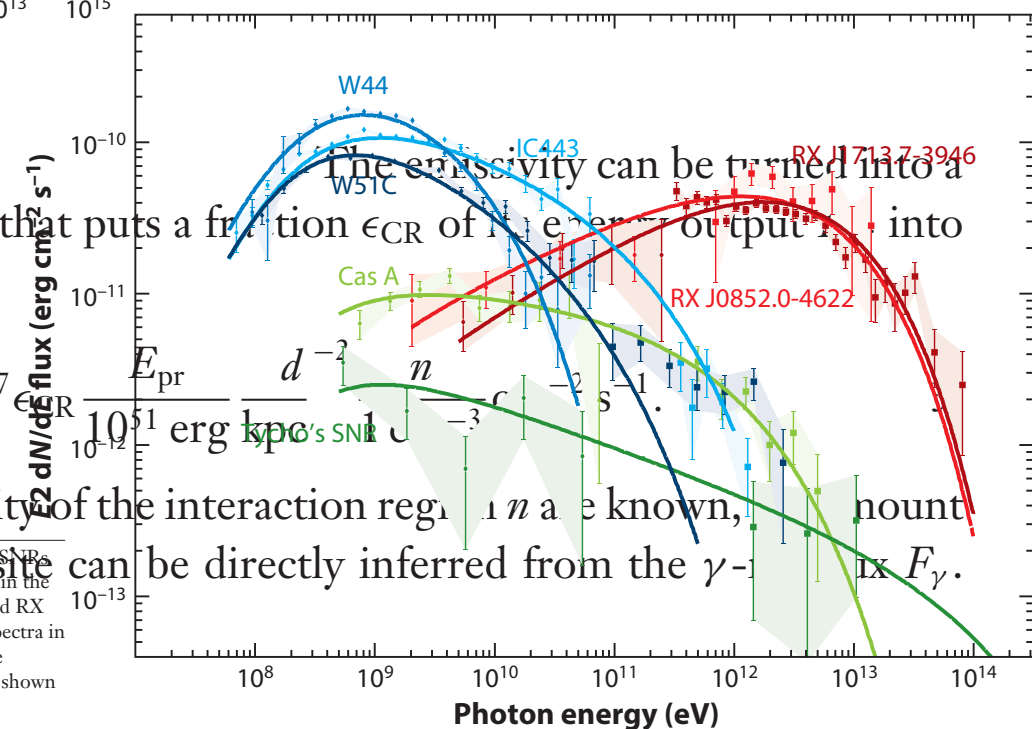


Figure 6
Typical γ -ray energy spectra for several of the most prominent supernova remnants (SNRs). Young SNRs ($< 1,000$ years) are shown in green. These typically show smaller γ -ray fluxes but rather hard spectra in the GeV and TeV bands. The older (but still referred to as young) shell-type SNRs RX J1713.7-3946 and RX J0852.0-4622 (Vela Junior) of ages $\sim 2,000$ years are shown in shades of red. These show very hard spectra in the GeV band ($\Gamma = 1.5$) and a peak in the TeV band with an exponential cutoff beyond 10 TeV. The middle-aged SNRs ($\sim 20,000$ years) interacting with molecular clouds (W44, W51C, and IC443) are shown in blue. Also shown are hadronic fits to the data (solid lines).

general considerations about accelerators

trajectory of particle in B field

centripetal force = Lorentz force

$$m \frac{v^2}{r} = q \cdot v \cdot B \quad m \cdot v = p \quad \text{momentum}$$

$$\frac{p}{r} = Z \cdot e \cdot B$$

$$r_L = \frac{p}{z \cdot e \cdot B} \quad \text{Larmor radius}$$

L dimension of accelerator

$$L > 2 r_L$$

closer look (Hillas 1984):

$$L > \frac{2r_L}{\beta}$$

velocity of scattering centers

$$\beta = \frac{v}{c}$$



$$L > \frac{2 \cdot p}{z \cdot e \cdot B \cdot \beta}$$

$$B \cdot L > \frac{2 \cdot p}{z \cdot e \cdot \beta}$$

Hillas criterion

in astrophysical units

$$r_L = 1.08 \text{ pc} \frac{E_{15}}{Z \cdot B_{\mu G}}$$

$$B_{\mu G} \cdot L_{pc} > \frac{2 \cdot E_{15}}{Z \cdot \beta}$$

**necessary condition
not sufficient**

$$E_{15} < Z \cdot B_{\mu G} \cdot L_{pc} \cdot \frac{\beta}{2}$$

Contribution of (regular) SNR-CR

$$E_c = Z \cdot 4.5 \cdot 10^6 \text{ GeV}$$

$$Q(p) = A Q_0 (Ap)^{-q} \exp\left(-\frac{Ap}{Z p_c}\right),$$

Table 1. Source spectral indices, q , and energy injected per supernova, f , for the different species of cosmic rays used in the calculation of the SNR-CRs spectra shown in Figures 1 and 2.

Particle type	q	f ($\times 10^{49}$ ergs)
Proton	2.24	6.95
Helium	2.21	0.79
Carbon	2.21	2.42×10^{-2}
Oxygen	2.25	2.52×10^{-2}
Neon	2.25	3.78×10^{-3}
Magnesium	2.29	5.17×10^{-3}
Silicon	2.25	5.01×10^{-3}
Iron	2.25	4.95×10^{-3}

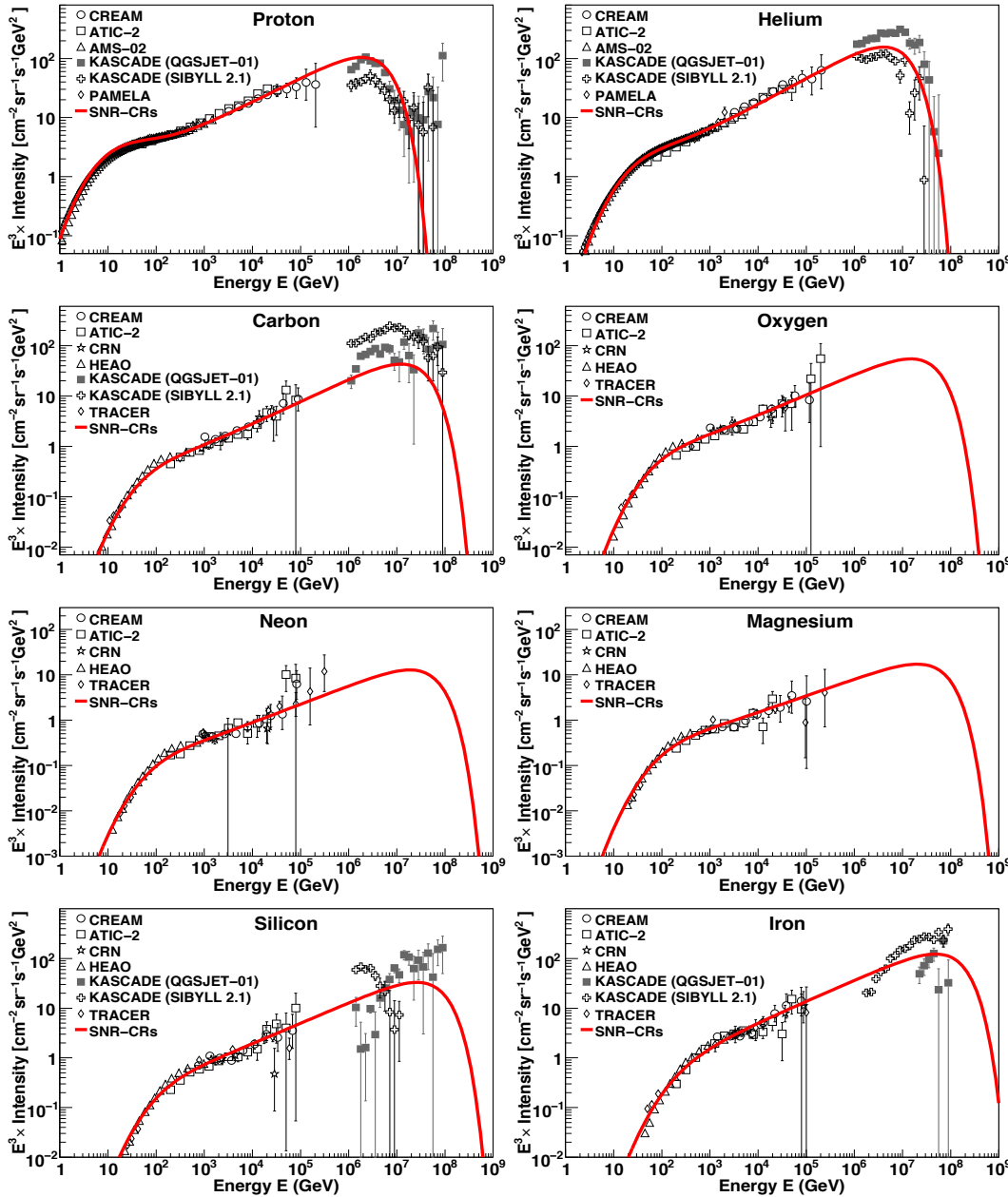
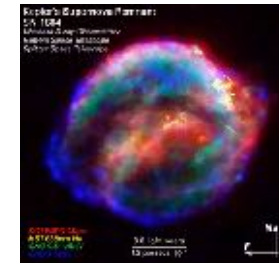


Fig. 1. Energy spectra for different cosmic-ray elements. *Solid line:* Model prediction for the SNR-CRs. *Data:* CREAM (Ahn et al. 2009; Yoon et al. 2011), ATIC-2 (Panov et al. 2007), AMS-02 (Aguilar et al. 2015a,b), PAMELA (Adriani et al. 2011), CRN (Müller et al. 1991; Swordy et al. 1990), HEAO (Engelmann et al. 1990), TRACER (Obermeier et al. 2011), and KASCADE (Antoni et al. 2005). Cosmic-ray source parameters (q, f) used in the calculation are given in Table 1. For the other model parameters (D_0, a, η, s), see text for details.



Transport equation for cosmic rays in the Galaxy

diffusion

energy loss (Bethe Bloch)

loss through interactions with ISM (spallation)

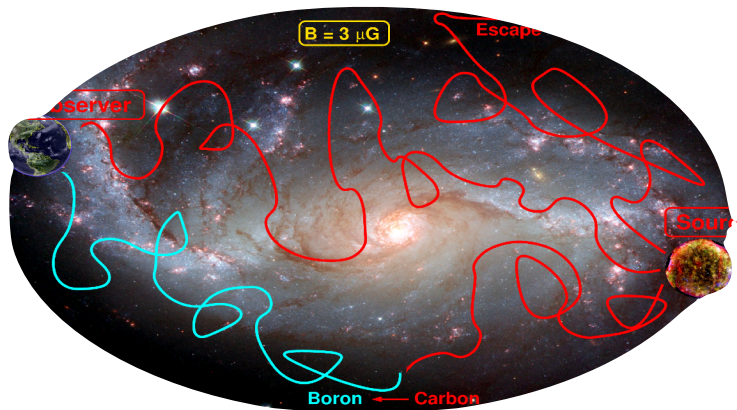
loss through radioactive decay

$$\frac{\partial N_i}{\partial t} = \nabla(D_i \nabla N_i) - \frac{\partial}{\partial E} (v_i N_i) - n \nu \sigma_i N_i - \frac{N_i}{\gamma \tau_i} + Q_i + \sum_{j>i} n \nu \sigma_{ij} N_j + \sum_{j>i} \frac{N_j}{\gamma_j \tau_{ij}}$$

source term

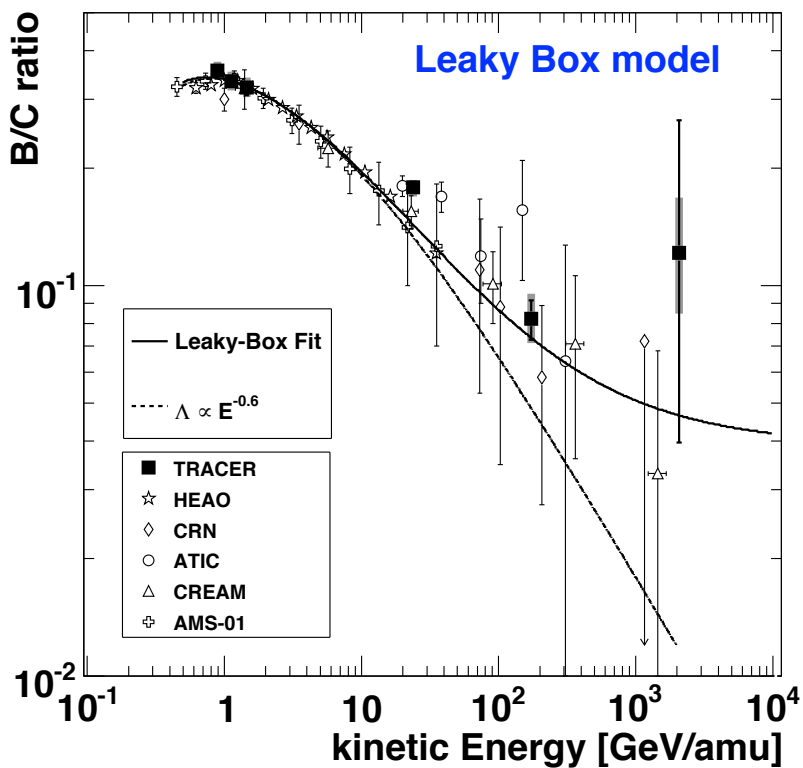
production through spallation of heavy nuclei

production through decay of heavy nuclei



Pathlength of cosmic rays in Galaxy

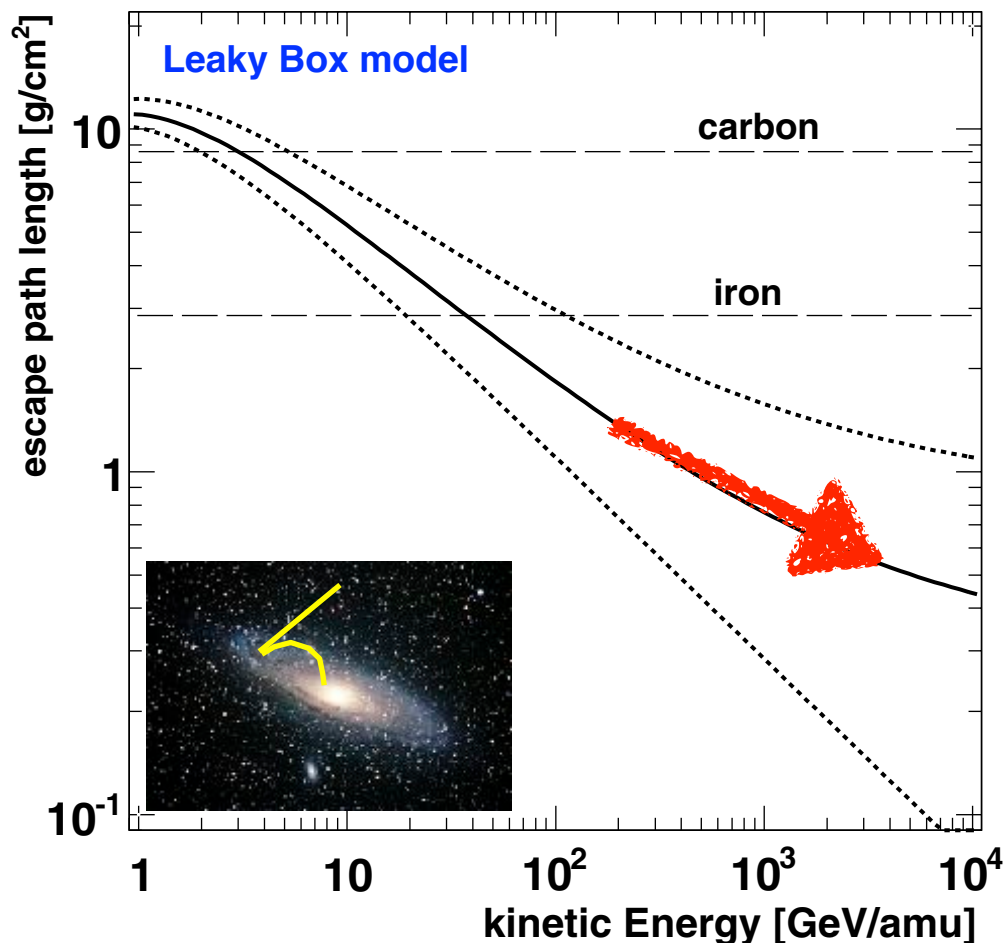
TRACER balloon experiment



Escape Path Length:

$$\Lambda_{esc}(E) = CE^{-\delta} + \Lambda_0$$

- ▶ Propagation index:
 $\delta = 0.64 \pm 0.02$.
- ▶ Residual path length:
 $\Lambda_0 = 0.7 \pm 0.2 \text{ g/cm}^2$.



$$\Lambda(R) = \frac{26.7\beta}{(\beta R)^\delta + (0.714 \cdot \beta R)^{-1.4}} + \Lambda_0 \text{ g/cm}^2,$$

Pathlength vs. interaction length

pathlength in Galaxy

$$\lambda_{esc} = 5 - 10 \text{ g/cm}^2$$

interaction length

nuclear radius

$$r = r_0 A^{1/3} \quad r_0 = 1.3 \cdot 10^{-13} \text{ cm}$$

cross section

$$\sigma_{p-A} = \pi (r_p + r_0 A^{1/3})^2$$

ISM: protons

$$n = 1/\text{cm}^3 \quad \rho = 1.67 \cdot 10^{-24} \text{ g/cm}^3$$

interaction length

$$\lambda_{p-A} = \frac{\rho}{\sigma_{p-A} \cdot n}$$

$$\lambda_{p-p} = 21 \text{ g/cm}^2 > \lambda_{esc}$$

$$\lambda_{p-Fe} = 1.6 \text{ g/cm}^2 < \lambda_{esc}$$

Shape of energy spectrum

$$\frac{dN}{dE} \propto E_0^\gamma$$

at source $\gamma \sim -2.1$

at Earth $\gamma \sim -2.6$ to -2.7

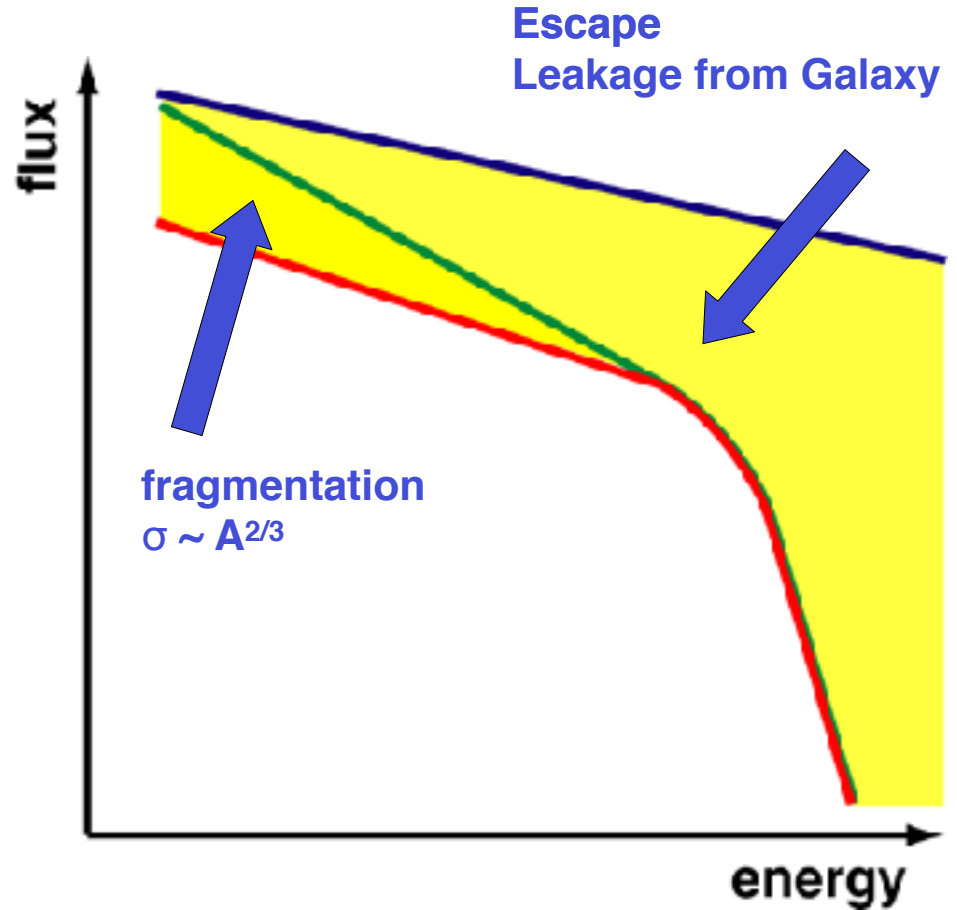
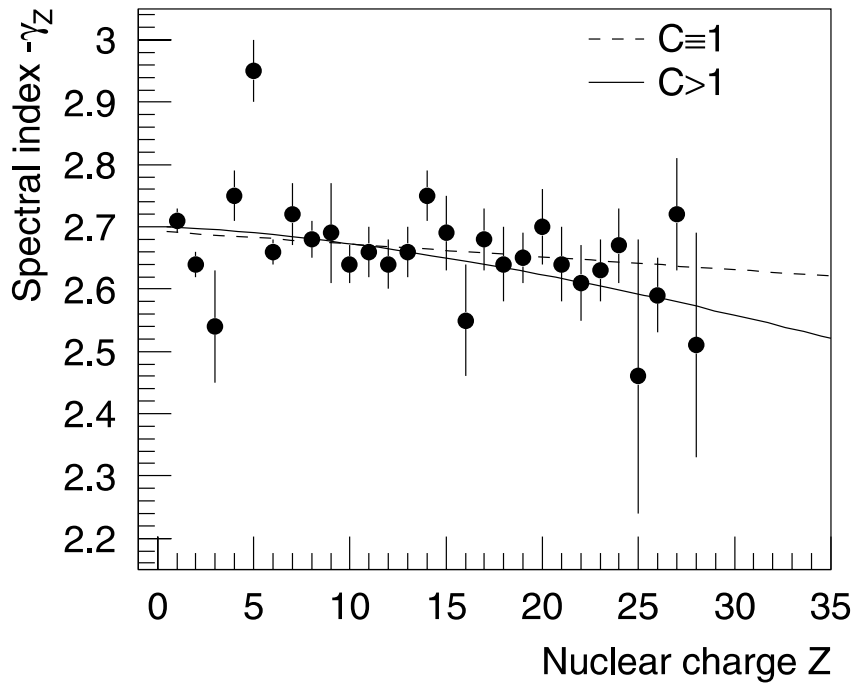


Fig. 5. Spectral index γ_Z versus nuclear charge Z (see Table 1). The solid line represents a three parameter fit according to Eq. (6), the dashed graph a linear fit.

Transport equation for cosmic rays in the Galaxy

diffusion

energy loss (Bethe Bloch)

loss through interactions with ISM (spallation)

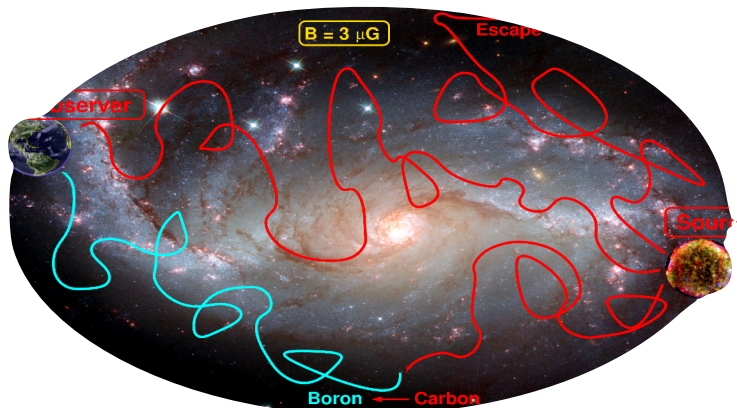
loss through radioactive decay

$$\frac{\partial N_i}{\partial t} = \nabla(D_i \nabla N_i) - \frac{\partial}{\partial E} (b_i N_i) - n \nu \sigma_i N_i - \frac{N_i}{\tau_i} + Q_i + \sum_{j>i} n \nu \sigma_{ij} N_j + \sum_{j>i} \frac{N_j}{\gamma_j \tau_{ij}}$$

source term

production through spallation of heavy nuclei

production through decay of heavy nuclei



Propagation of super-high-energy cosmic rays in the Galaxy

Jörg R. Hörandel ^{a,*}, Nikolai N. Kalmykov ^b, Aleksei V. Timokhin ^c

The steady-state diffusion equation for the cosmic-ray density $N(r)$ is (neglecting nuclear interactions and energy losses)

$$-\nabla_i D_{ij}(r) \nabla_j N(r) = Q(r). \quad (1)$$

$Q(r)$ is the cosmic-ray source term and $D_{ij}(r)$ the diffusion tensor.

Under the assumption of azimuthal symmetry and taking into account the predominance of the toroidal component of the magnetic field, Eq. (1) is presented in cylindrical coordinates as

$$\left[-\frac{1}{r} \frac{\partial}{\partial r} r D_{\perp} \frac{\partial}{\partial r} - \frac{\partial}{\partial z} D_{\perp} \frac{\partial}{\partial z} - \frac{\partial}{\partial z} D_A \frac{\partial}{\partial r} + \frac{1}{r} \frac{\partial}{\partial r} r D_A \frac{\partial}{\partial z} \right] N(r, z) = Q(r, z), \quad (2)$$

where $N(r, z)$ is the cosmic-ray density averaged over the large-scale fluctuations with a characteristic scale $L \sim 100$ pc [3]. $D_{\perp} \propto E^m$ is the diffusion coefficient, where m is much less than one ($m \approx 0.2$), and $D_A \propto E$ the Hall diffusion coefficient. The influence of Hall diffusion becomes predominant at high energies ($>10^{15}$ eV). The sharp

The magnetic field of the Galaxy consists of a large-scale regular and a chaotic, irregular component $\vec{B} = \vec{B}_{\text{reg}} + \vec{B}_{\text{irr}}$. A purely azimuthal magnetic field was assumed for the regular field

$$B_z = 0, \quad B_r = 0, \quad B_{\phi} = 1 \mu\text{G} \exp\left(-\frac{z^2}{z_0^2} - \frac{r^2}{r_0^2}\right),$$

where $z_0 = 5$ kpc and $r_0 = 10$ kpc are constants [3].

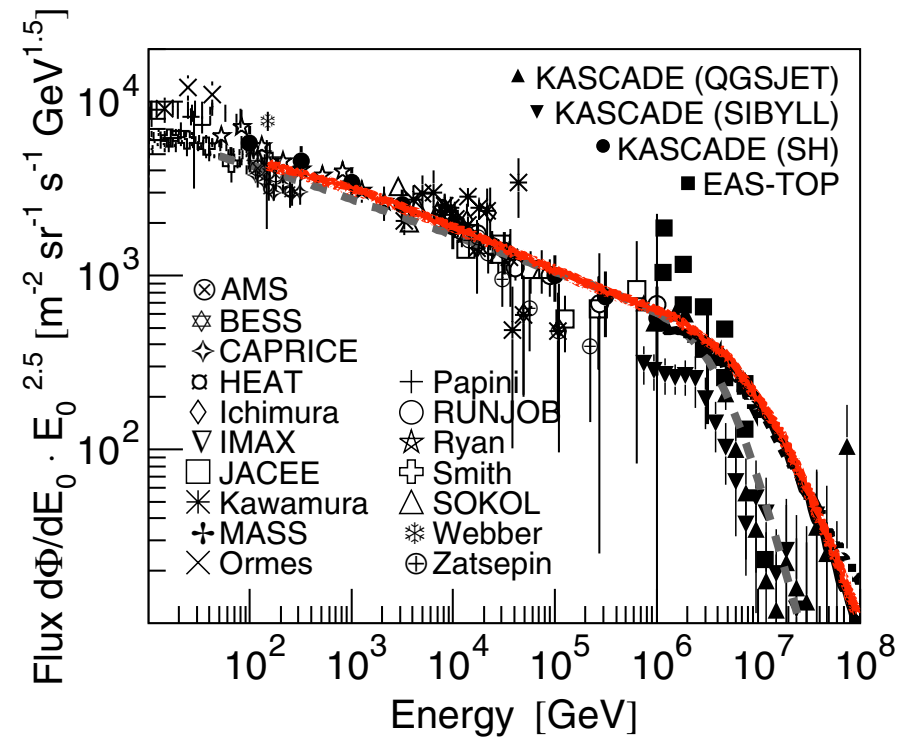


Fig. 7. Proton flux as obtained from various measurements, for references see [28], compared to the spectra shown in Fig. 6 (black lines) and the *polygonoato* model [26] (grey, dashed line).

Origin of the knee?

Acceleration (SNR)

- .. in SNR
- .. in SNR + radio galaxies
- .. in oblique shocks
- .. in variety of SNR
- Single source model
- Reacceleration in galactic wind

Leakage from Galaxy

- Minimum pathlength model
- Anomalous diffusion model
- Hall diffusion model
- Diffusion in turbulent magnetic fields
- Diffusion and drift

~~γ -ray bursts~~

- ~~Cannonball model~~
- ~~Acceleration in GRB + diffusion~~
- ~~Acceleration in GRB $E_{max} \sim A$~~

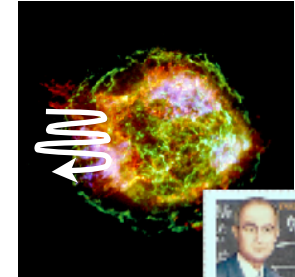
~~Interaction with background particles~~

- ~~Diffusion model + photo-disintegration - Tkaczyk~~
- ~~Interaction with neutrinos in galactic halo~~
- ~~Photo-disintegration (optical and UV photons)~~

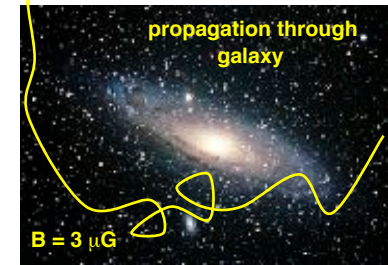
~~Particle physics in atmosphere~~

- ~~Gravitons, SUSY~~

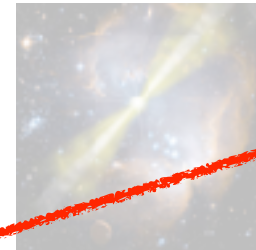
- Berezhko & Ksenofontov
- Stanev ..
- Kobayakawa ..
- Sveshnikova
- Erlykin & Wolfendale
- Völk & Zirakashvili



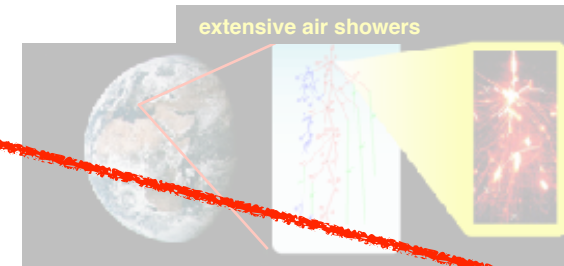
- Swordy
- Lagutin ..
- Ptuskin .., Kalmykov ..
- Ogio & Kakimoto
- Roulet ..



- Plaga
- Wick ..
- Dar ..



- Dova ..
- Candia ..

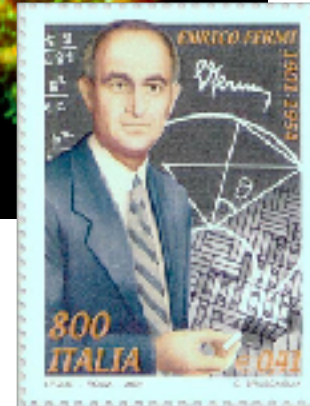
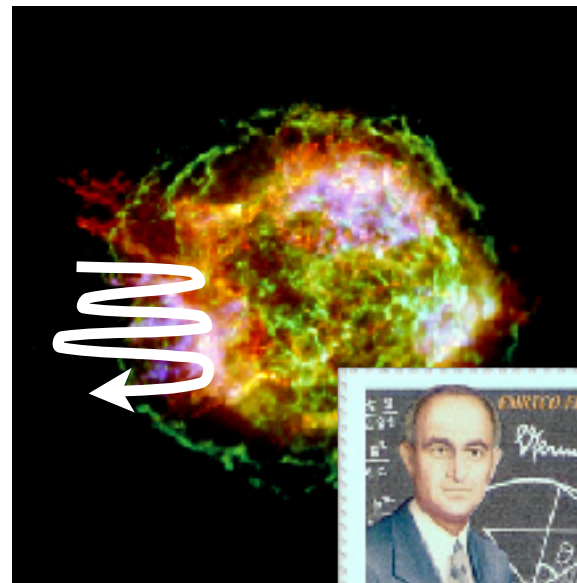
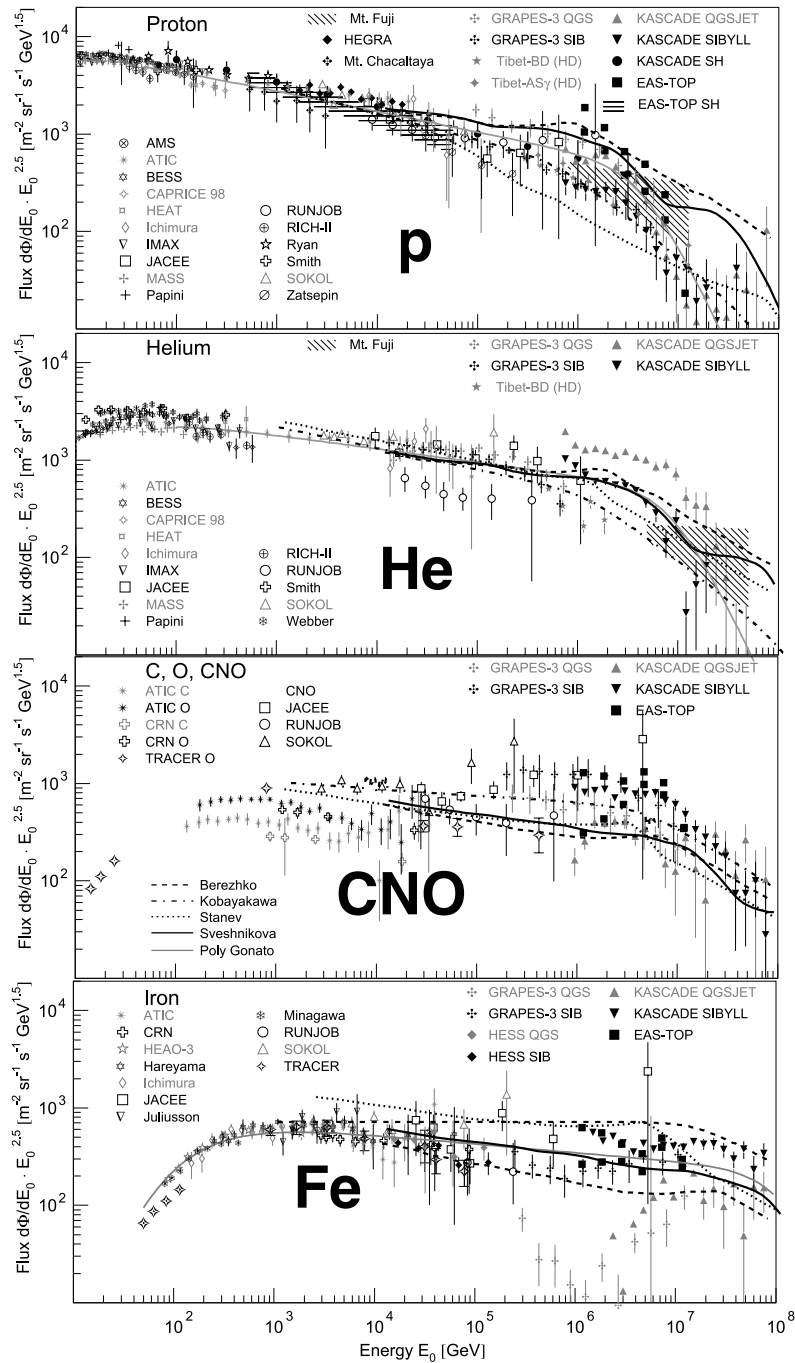


- Kazanas & Nicolaidis

maximum energy

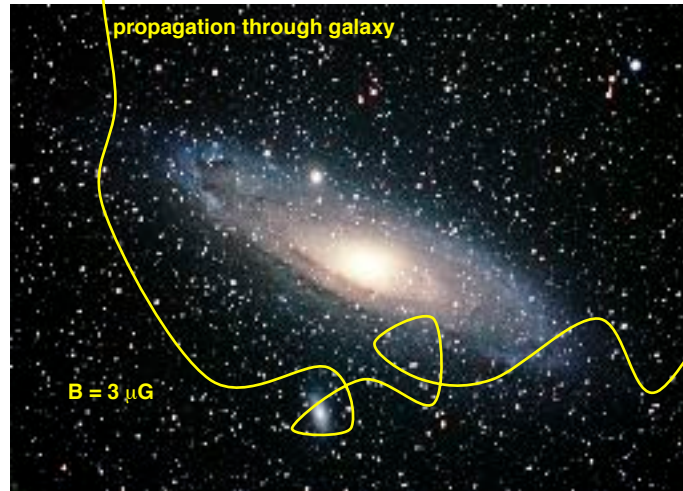
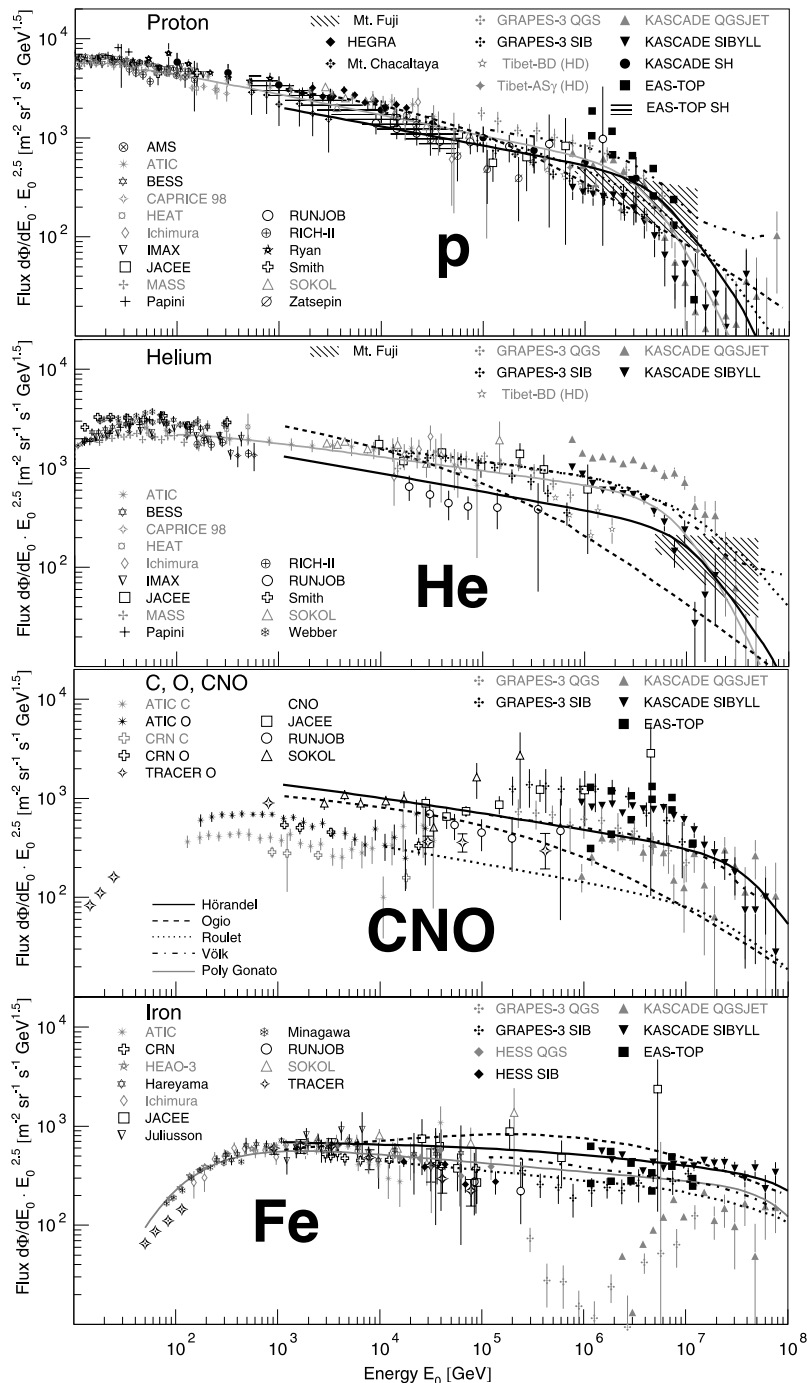
$$E_{max} \propto B \cdot Z$$

$$E_{max} \approx Z \cdot 100 \text{ TeV} \dots Z \cdot 5 \text{ PeV}$$



leakage from Galaxy

$$E_k \propto Z$$



Transition to extragalactic CR component

J. Blümer et al. / Progress in Particle and Nuclear Physics 63 (2009) 293–338

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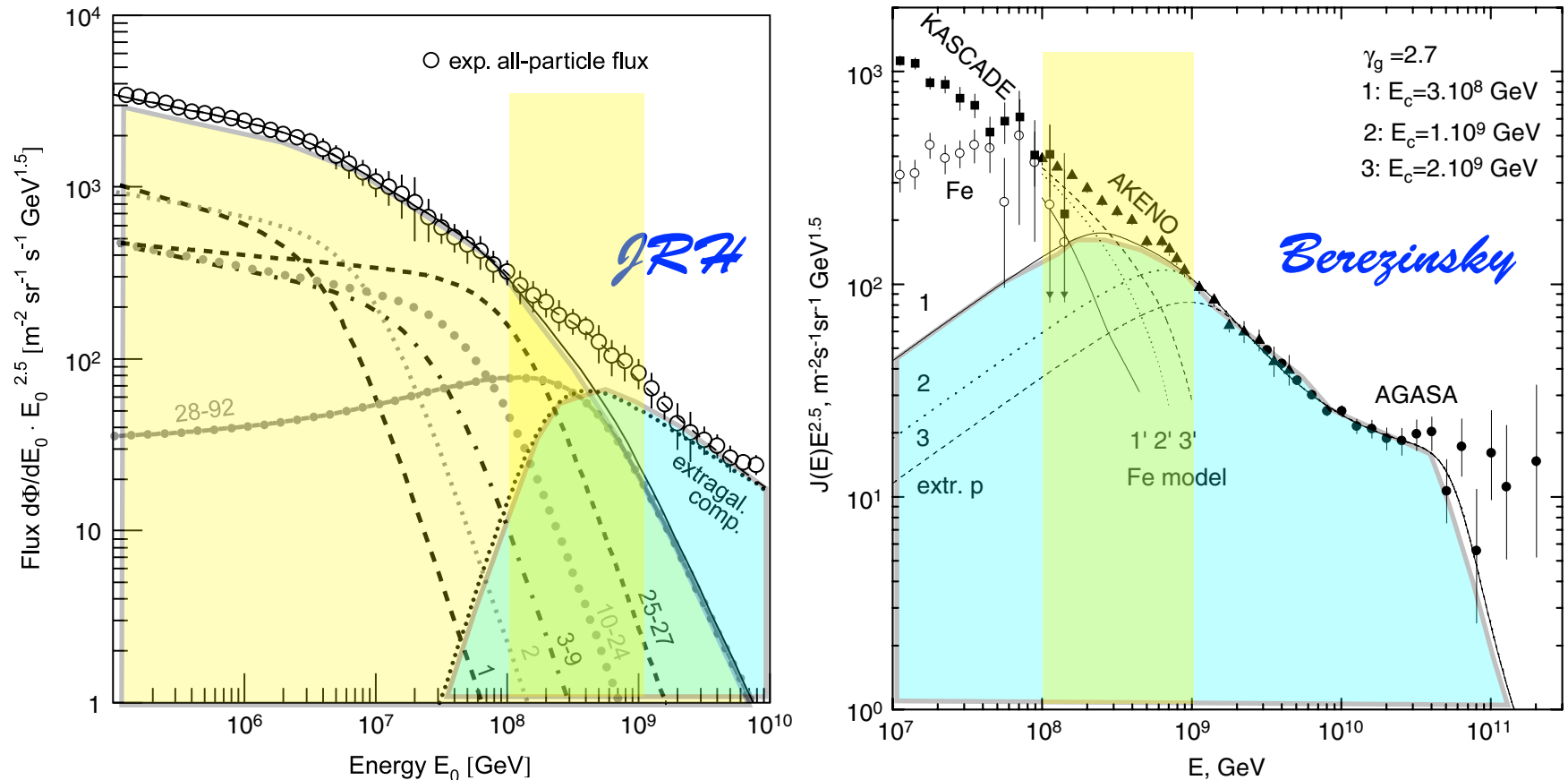
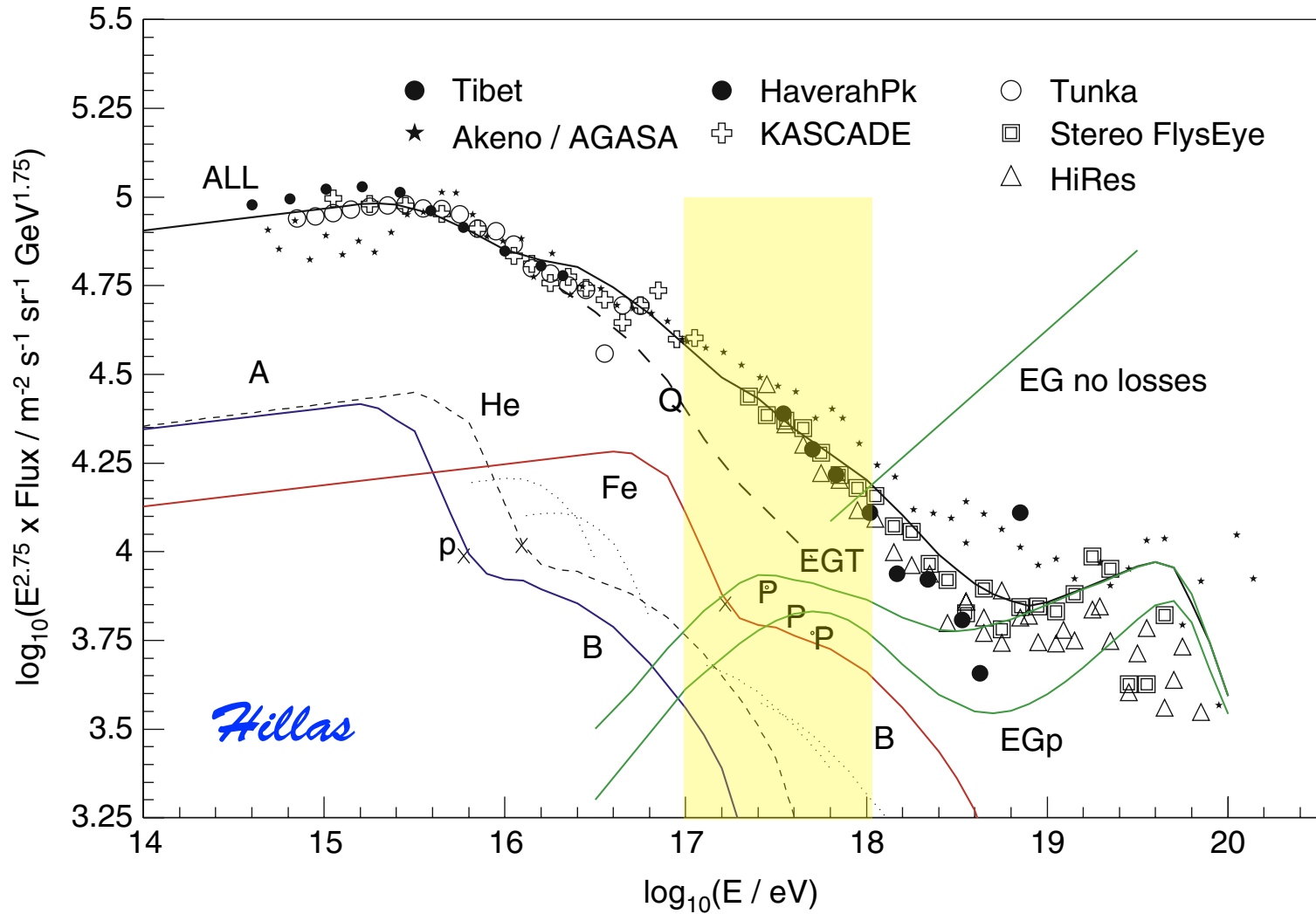


Fig. 26. *Left panel:* Cosmic-ray energy spectra according to the poly-gonato model [2]. The spectra for groups of elements are labeled by their respective nuclear charge numbers. The sum of all elements yields the galactic all-particle spectrum (—) which is compared to the average measured flux. In addition, a hypothetical extragalactic component is shown to account for the observed all-particle flux (---). *Right panel:* Transition from galactic to extragalactic cosmic rays according to Berezhinsky et al. [451]. Calculated spectra of extragalactic protons (curves 1, 2, 3) and of galactic iron nuclei (curves 1', 2', 3') are compared with the all-particle spectrum from the Akeno and AGASA experiments. KASCADE data are shown as filled squares for the all-particle flux and as open circles for the flux of iron nuclei.

Transition to extragalactic CR component



„classical“ supernovae + additional component

Contribution of (regular) SNR-CR to all-particle spectrum

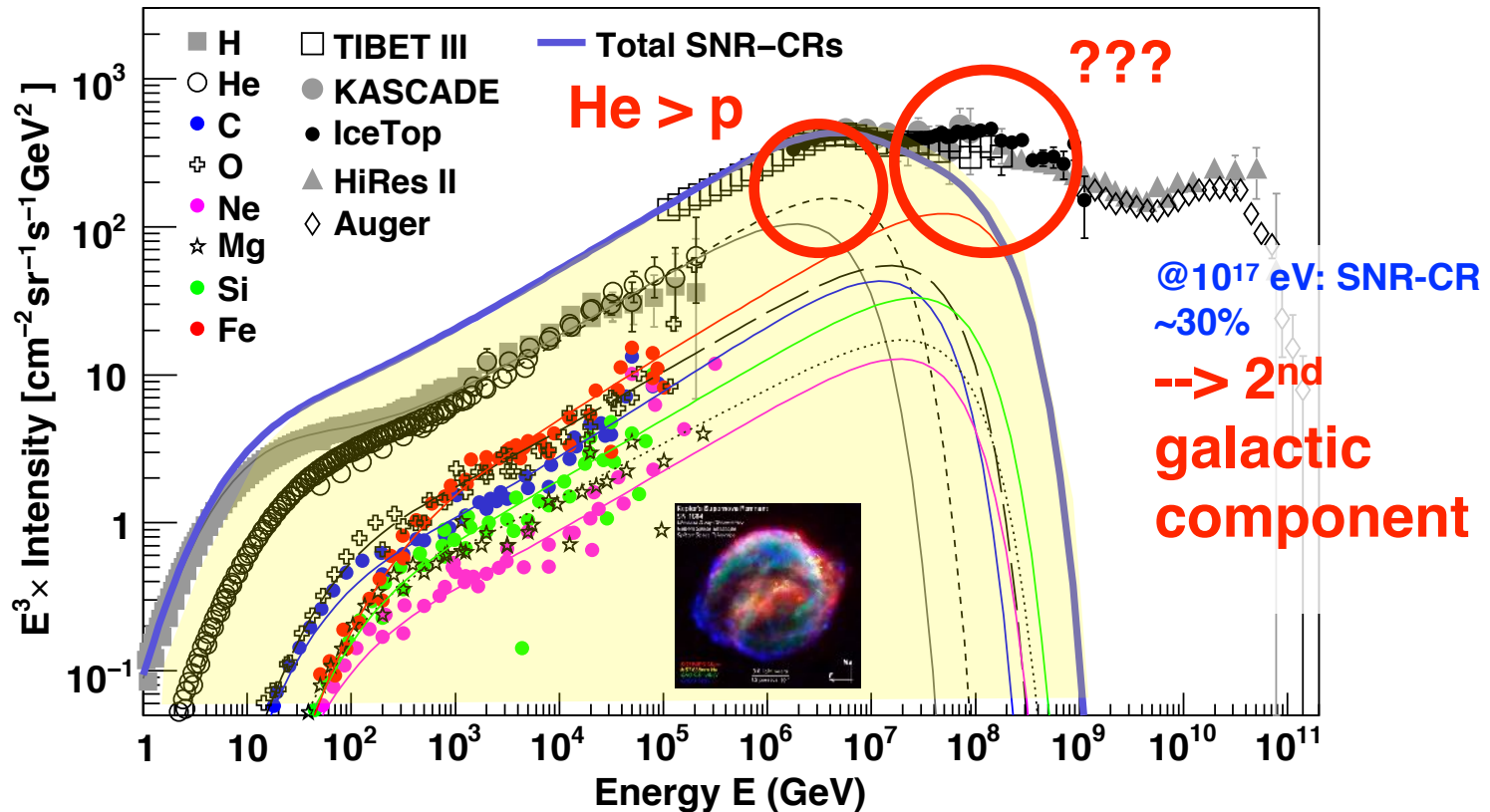
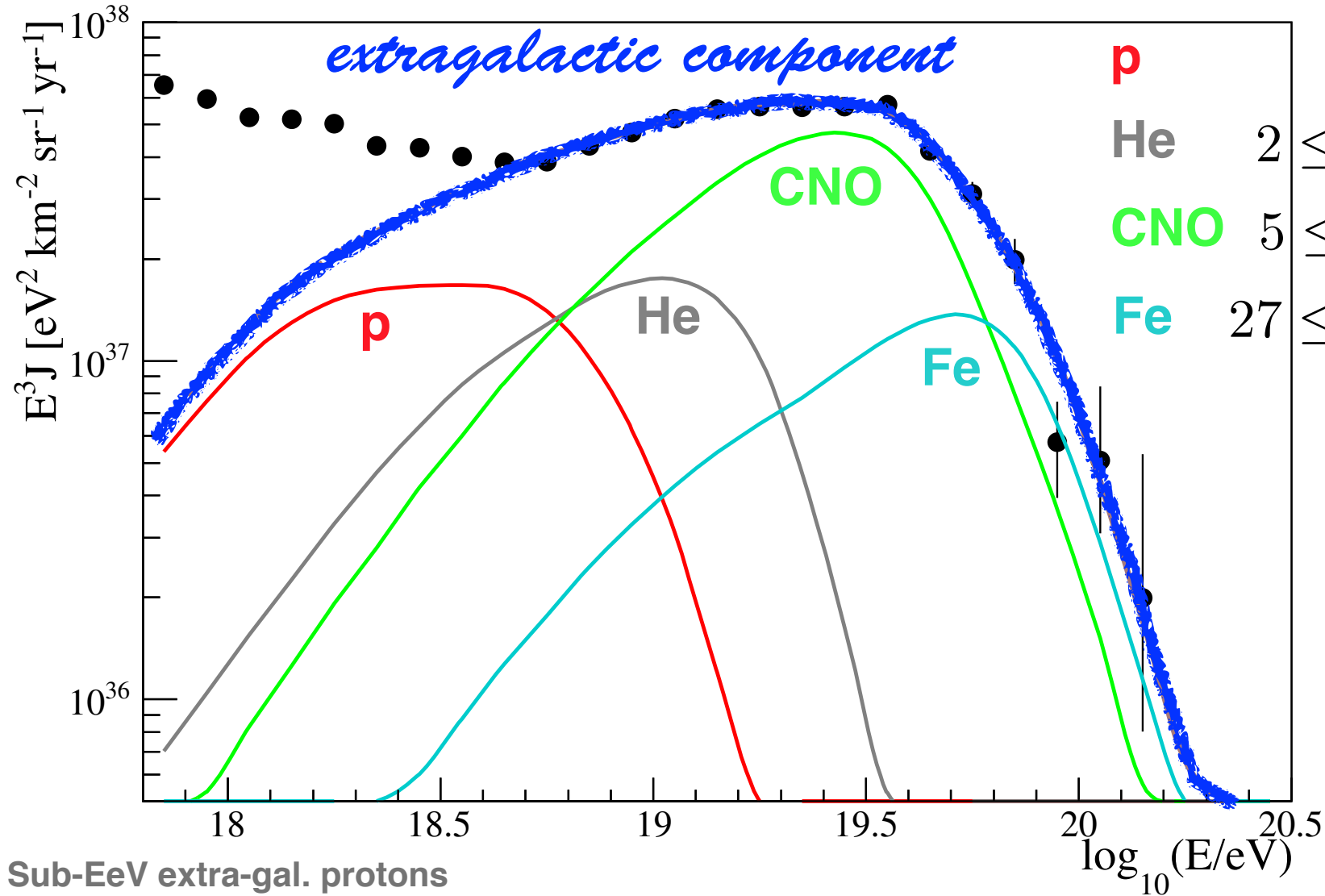


Fig. 2. Contribution of SNR-CRs to the all-particle cosmic-ray spectrum. The thin lines represent spectra for the individual elements, and the thick-solid line represents the total contribution. The calculation assumes an exponential cut-off energy for protons at $E_c = 4.5 \times 10^6$ GeV. Other model parameters, and the low-energy data are the same as in Figure 1. Error bars are shown only for the proton and helium data. High-energy data: KASCADE (Antoni et al. 2005), IceTop (Aartsen et al. 2013), Tibet III (Amenomori et al. 2008), the Pierre Auger Observatory (Schulz et al. 2013), and HiRes II (Abbasi et al. 2009).

~8% of mechanical power of SN --> CRs

Combined fit of spectrum and composition data as measured by the Pierre Auger Observatory



p	$A = 1$
He	$2 \leq A \leq 4$
CNO	$5 \leq A \leq 22$
Fe	$27 \leq A \leq 56$

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Sub-EeV extra-gal. protons from interactions of heavier nuclei

see lecture on UHECR later

all-particle spectra including 2nd galactic component

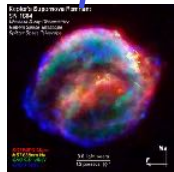
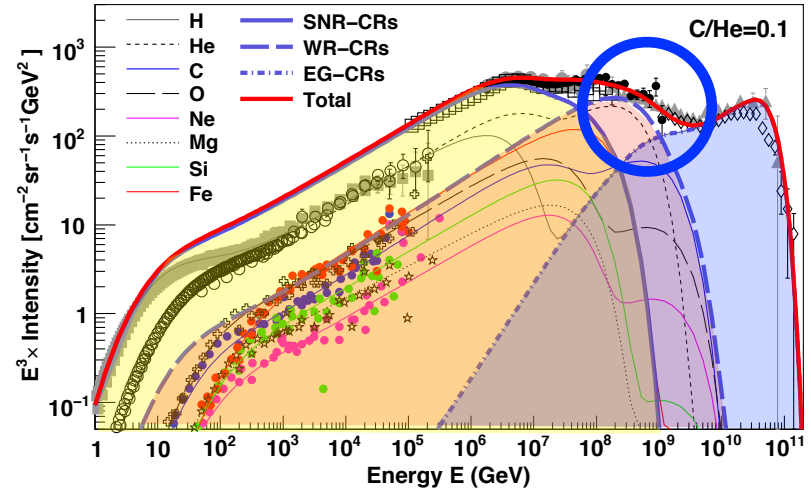
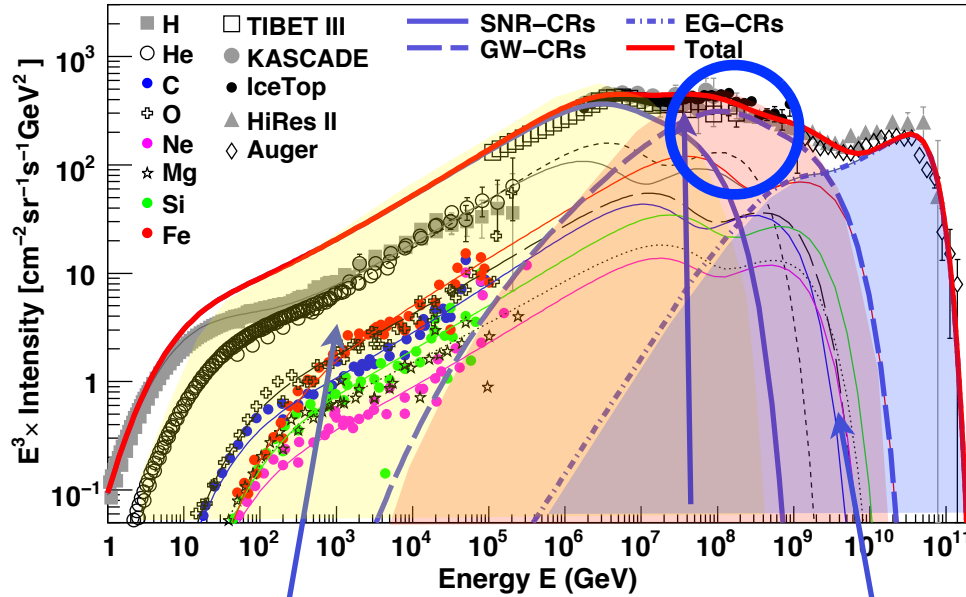
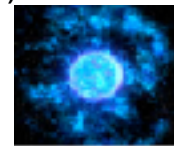
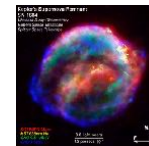
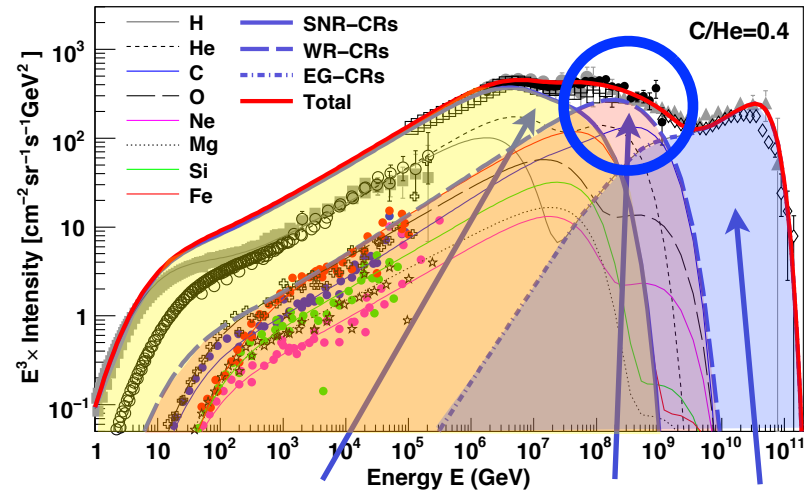


Table 3. Injection energy of SNR-CRs used in the calculation of all-particle spectrum in the WR-CR model (Figure 6).

Particle type	C/He = 0.1 $f(\times 10^{49} \text{ ergs})$	C/He = 0.4 $f(\times 10^{49} \text{ ergs})$
Proton	8.11	8.11
Helium	0.67	0.78
Carbon	2.11×10^{-2}	0.73×10^{-2}
Oxygen	2.94×10^{-2}	2.94×10^{-2}
Neon	4.41×10^{-3}	4.41×10^{-3}
Magnesium	6.03×10^{-3}	6.03×10^{-3}
Silicon	5.84×10^{-3}	5.84×10^{-3}
Iron	5.77×10^{-3}	5.77×10^{-3}



Mean logarithmic mass ($\ln A$)

WR-CR (C/He=0.4) + EG scenarios

$$\langle \ln A \rangle \equiv \sum_i r_i \ln A_i, \quad \langle \ln A \rangle = \frac{X_{\max}^{\text{meas}} - X_{\max}^p}{X_{\max}^{\text{Fe}} - X_{\max}^p} \cdot \ln A_{\text{Fe}}.$$

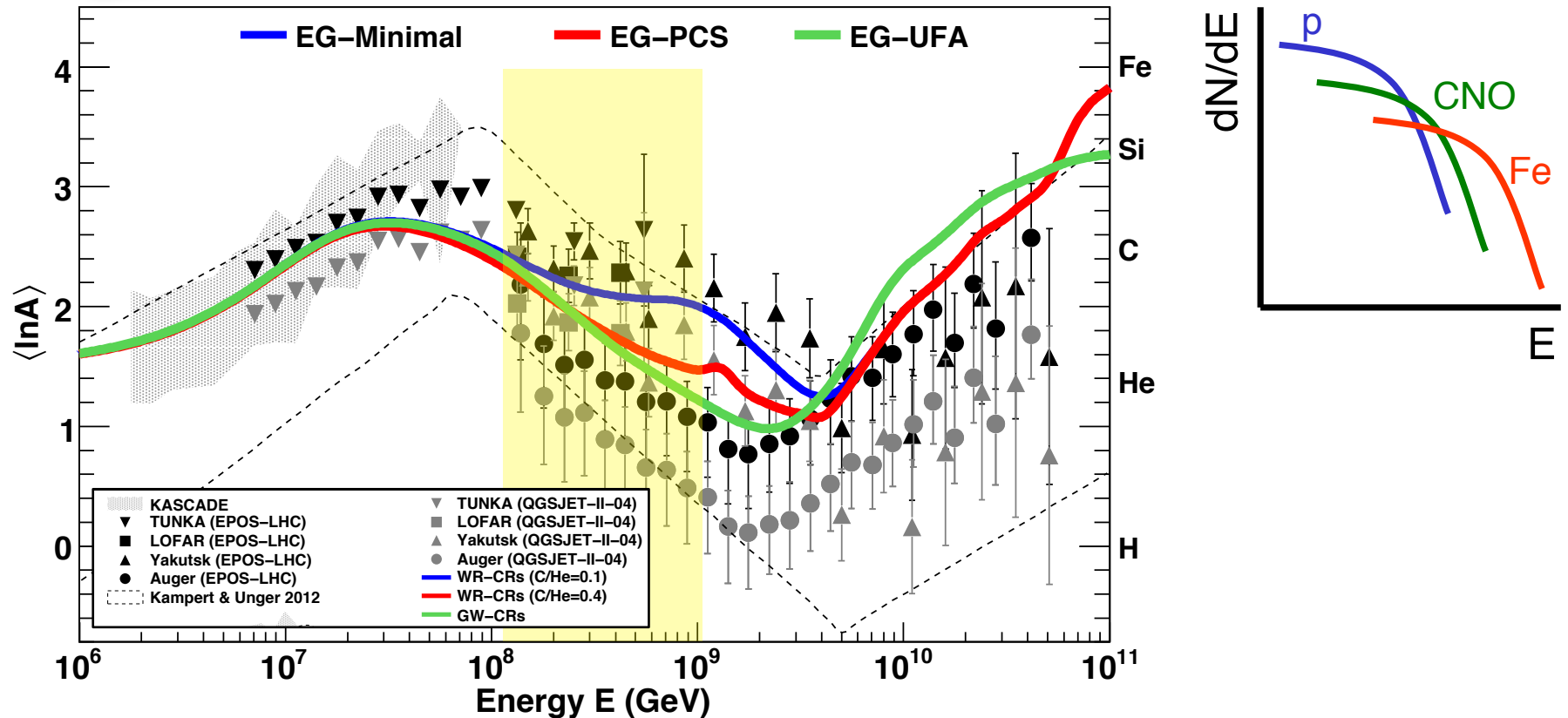


Fig. 11. Mean logarithmic mass for the three different EG-CR models combined with the WR-CR (C/He = 0.4) model. Data are the same as in Figure 8. Results obtained using WR-CR (C/He = 0.1) model are shown in Appendix B.

Cosmic rays at the knee

Results and implications

- knee in all-particle spectrum at ~ 4.5 PeV caused by fall-off of light elements (p, He)
- experimental (world) data indicate rigidity-dependent fall-off of individual elements
(in particular unfolding by KASCADE[-Grande] and IceCube/Top)
- spectrum above knee is superposition of individual spectra (elemental knees)
 - > fine structure in all-particle spectrum
 - > end of galactic CR component
- astrophysical origin of knee:
combination of maximum energy attained in sources (Supernovae)
(Hillas criterion)
and leakage from Galaxy
- 2nd galactic component at $\sim 10^{17}$ eV?
- extra-galactic origin $> 10^{18}$ eV